The ultraviolet luminosity function and star formation rate of the Coma cluster

L. Cortese,1⋆ G. Gavazzi2 and A. Boselli3
1School of Physics and Astronomy, Cardiff University, Cardiff CF24 3AA
2Università degli Studi di Milano-Bicocca, Piazza della Scienza 3, 20126 Milano, Italy
3Laboratoire d’Astrophysique de Marseille, UMR 6110 CNRS, 38 rue F. Joliot-Curie, F-13388 Marseille, France

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ABSTRACT
We present estimates of the GALEX near-ultraviolet (NUV) and far-ultraviolet (FUV) luminosity functions (LFs) of the Coma cluster, over a total area of ∼9 deg² (∼25 Mpc²), i.e. from the cluster centre to the virial radius. Our analysis represents the widest and deepest UV investigation of a nearby cluster of galaxies made to date. The Coma UV LFs show a faint-end slope steeper than the one observed in the local field. This difference, more evident in NUV, is entirely due to the contribution of massive quiescent systems (e.g. ellipticals, lenticulars and passive spirals), more frequent in high-density environments. On the contrary, the shape of the UV LFs for Coma star-forming galaxies does not appear to be significantly different from that of the field, consistently with previous studies of local and high-redshift clusters. We demonstrate that such similarity is only a selection effect, not providing any information on the role of the environment on the star formation history of cluster galaxies. By integrating the UV LFs for star-forming galaxies (corrected for the first time for internal dust attenuation), we show that the specific star formation rate (SSFR) of Coma is significantly lower than the integrated SSFR of the field and that Coma-like clusters contribute only <7 per cent of the total SFR density of the local universe. Approximately two-thirds of the whole star formation in Coma is occurring in galaxies with $M_{\text{star}} < 10^{10} M_\odot$. The vast majority of star-forming galaxies has likely just started its first dive into the cluster core and has not yet been affected by the cluster environment. The total stellar mass accretion rate of Coma is $(0.6–1.8) \times 10^{12} M_\odot \text{Gyr}^{-1}$, suggesting that a significant fraction of the population of lenticular and passive spirals observed today in Coma could originate from infalling galaxies accreted between $z \sim 1$ and 0.

Key words: galaxies: clusters: individual: Coma (Abell1656) – galaxies: evolution – galaxies: fundamental parameters – galaxies: luminosity function, mass function – ultraviolet: galaxies.

1 INTRODUCTION
Star formation is probably the most fundamental of all astrophysical processes regulating galaxy evolution: not only the properties of galaxies depend on how their stars were formed, but also the properties of the interstellar medium (ISM) are controlled to a large extent by the various feedback effects of star formation. The overall evolution of galaxies depends on the rate at which their interstellar gas is converted into stars and the star formation rate (SFR) depends on the rate at which diffuse interstellar matter is collected into star-forming (SF) regions. Consequently, it is of great importance for galaxy evolution studies to understand how stars form and what determines their properties.

The cluster environment is well known to affect both the star formation activity and morphology of the member galaxies, resulting in the morphology–density (e.g. Dressler 1980; Whitmore, Gilmore & Jones 1993) and star formation–density (e.g. Gavazzi et al. 2002b; Lewis et al. 2002; Gómez et al. 2003; Boselli & Gavazzi 2006) relations. The scale over which the number density of elliptical galaxies increases towards the centre of rich clusters is very small (<0.5 Mpc), while the one over which the star formation activity and the fraction of late-type galaxies decrease is much larger (a few Mpc). This indicates that separate processes contribute to shaping these relations. The mechanism involved in the early-type central enhancement is probably connected to the early phases of structure

⋆E-mail: luca.cortese@astro.cf.ac.uk

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formation (e.g. De Lucia et al. 2006; Ellis et al. 1997; Stanford, Eisenhardt & Dickinson 1998; Renzini 2006), while the progressive quenching of the star formation and lack of spirals towards the centre of rich clusters could represent a separate, diluted in space and time, process probably related to the cluster accretion history (e.g. Dressler 2004; Boselli & Gavazzi 2006; Cortese et al. 2006b; Poggianti et al. 2006).

Although the effects of the environment on the star formation activity of cluster galaxies are gradually being unveiled (e.g. Moore et al. 1996; Poggianti et al. 1999; Vollmer et al. 2001; Dressler 2004; Boselli & Gavazzi 2006; Boselli et al. 2006; Poggianti et al. 2006), we still know very little about the contribution of clusters of galaxies to the whole star formation history (SFH) and the stellar mass growth of the universe. The cosmic SFR dramatically decreases from $z \sim 1$ to the present epoch (Lilly et al. 1996; Madau et al. 1996), but we do not know if this drop varies with the environment (Kodama & Bower 2001; Cooper et al. 2008; Finn et al. 2008), what fraction of the cosmic SFR takes place in cluster galaxies and how this number evolves with redshift.

In addition, a significant increase from $z \sim 1$ to 0 in the stellar mass density of the universe has been observed (e.g. Drory et al. 2005; Borch et al. 2006; Fontana et al. 2006). Interestingly very little growth is found in the stellar mass of blue sequence galaxies, and the vast majority of stars formed since $z \sim 1$ appears to inhabit galaxies which, at the present epoch, lie in the red sequence (e.g. Cimatti, Daddi & Renzini 2007; Bell et al. 2007; Brown et al. 2007; Faber et al. 2007). This automatically implies that a significant fraction of SF galaxies has migrated from the blue to the red sequence. The mechanism responsible for the quenching of the star formation in blue galaxies is still unknown and remains one of the greatest challenges for modern extragalactic astronomy. However, it is interesting to note that a similar migration, from SF to passive galaxies, is usually observed in high-density environments (e.g. Boselli et al. 2008), perhaps suggesting that the environment might play an important role.

In order to clarify these issues, it is therefore crucial to quantify the contribution of galaxy clusters to the total SFR and stellar mass budget of the universe, and their evolution with redshift. To reach this goal we need to determine the SFR, luminosity and mass functions of the cluster population up to the virial radius, avoiding bias in our determination of the global cluster properties by observing a population that is not representative of the whole cluster environment. This kind of study has become possible only recently thanks to the advent of wide field surveys. In particular, the Galaxy Evolution Explorer (GALEX; Martin et al. 2005) is providing for the first time precise ultraviolet (UV) photometry of galaxies over large stretches of the sky, opening a new era of extragalactic UV astronomy. The UV emission is dominated by young stars of intermediate masses ($2 < M < 5 \, \text{M}_\odot$, e.g. Kennicutt 1998) and thus represents an ideal tracer to identify and quantify the star formation activity of clusters. Surprisingly, very few statistical studies of galaxy clusters in UV have been carried out so far and great part of what we know about UV properties of cluster galaxies is still based on pioneering observations by the FAUST telescope (e.g. Deharveng et al. 1994; Brosch et al. 1997) and the balloon-borne UV telescopes SCAP and FOCA (e.g. Donas et al. 1987; Donas, Milliard & Laget 1995; Andreon 1999; Cortese et al. 2003a). For example, we still lack an accurate determination of the UV luminosity functions (LFs) in high-density environments. The FOCA UV LFs were in fact affected by large statistical errors due to the insufficient redshift coverage for UV-selected galaxies and to the uncertainty in the UV background counts. Moreover, the only GALEX UV LF for a nearby cluster presented so far (Abell 1367; Cortese et al. 2005) does not sample the whole cluster region and suffers from low number statistics at high luminosities. Thus, in order to investigate the UV properties of nearby clusters we carried out a wide panoramic survey in near-UV (NUV) and far-UV (FUV) of the nearby Coma cluster of galaxies.

Coma (Abell 1656) is one of the nearest richest clusters, and it is considered as the prototype of evolved, relaxed cluster of galaxies. In the last decades, it has been one of the main laboratory for environmental studies at all wavelengths (e.g. see Boselli & Gavazzi 2006 and references therein) and more recently it has become the focus of a Hubble/ACS Treasury survey (Carter et al. 2008) making it an ideal target for a deep census of UV properties in a nearby cluster up to its virial radius.

In this paper, we present the determination of the NUV and FUV LFs, the total SFR of the Coma cluster and its contribution to the SFR density in the local universe.

We assume a distance modulus of $m - M = 35.0$ mag for the Coma cluster, corresponding to a distance of 100 Mpc and a scale of 1.67 Mpc deg$^{-1}$ for $H_0 = 70$ Mpc$^{-1}$ km s$^{-1}$.

2 THE DATA

GALEX provides FUV ($\lambda = 1539$ Å, $\Delta \lambda = 442$ Å) and NUV ($\lambda = 2316$ Å, $\Delta \lambda = 1069$ Å) images with a circular field of view of ~0.6 radius. The spatial resolution is ~4–5 arcsec. Details of the GALEX instrument can be found in Martin et al. (2005) and Morrissey et al. (2005, 2007).

The data analysed in this paper are part of a Cycle 1 Guest Investigator proposal (P.I. G. Gavazzi, proposal number: GI1-039). 12 pointings have been originally approved but only nine have been completed so far. In addition, in order to increase the GALEX coverage of the Coma cluster, we include in our analysis two more fields publicly available as part of the Nearby Galaxy Survey: NGC1514 and ComaSPECA. The coordinates and exposure times of the 11 fields are shown in Table 1. Their typical exposure time is ~1500 s in FUV and it varies between 1500 and 3000 s in NUV. The total area covered is ~9 deg$^2$ (see Fig. 1). The major limitation of these observations is that they do not include the centre of Coma. The presence of an UV bright star near the cluster centre makes this region unobservable in NUV. FUV observations of this region are present in the GALEX schedule but they have not yet been executed. The total area not observed near the centre of Coma is ~0.26 deg$^2$ (see Fig. 1).

Each field has been reduced using the standard GALEX Data Analysis Pipeline (Morrissey et al. 2007). Sources were detected and measured using SEXTTRACTOR (Bertin & Arnouts 1996). As the NUV images are significantly deeper than the FUV images, sources were selected, and their parameters determined, in the NUV. FUV parameters were extracted within the same apertures. To avoid artefacts present at the edge of the field, we considered only the central 0.58 radius from the field centre. We used a variable SEXTTRACTOR deblending parameter contrary to the standard GALEX pipeline, providing reliable magnitudes (MAGAUTO) also for very extended sources (Cortese et al. 2005). By comparing the flux estimates for objects detected in more than one field we find an average uncertainty in the NUV and FUV magnitudes of ~0.15 mag. This value decreases to ~0.07 mag for objects brighter than $m_{\text{AB}}(\text{NUV}) \sim 19$ mag. Magnitudes are corrected for Galactic extinction using the Schlegel, Finkbeiner & Davis (1998) reddening map and the Galactic extinction curve of Cardelli, Clayton & Mathis.
Table 1. GALEX observations of the Coma cluster.

<table>
<thead>
<tr>
<th>Field</th>
<th>RA (J.2000)</th>
<th>Dec. (J.2000)</th>
<th>FUV integration time (s)</th>
<th>NUV integration time (s)</th>
</tr>
</thead>
<tbody>
<tr>
<td>NGA, DDO154</td>
<td>12:54:13.0</td>
<td>26:55:31</td>
<td>1451</td>
<td>1451</td>
</tr>
<tr>
<td>Coma_MOS03</td>
<td>12:55:10.0</td>
<td>28:24:00</td>
<td>1689</td>
<td>2547</td>
</tr>
<tr>
<td>Coma_MOS04</td>
<td>12:57:12.0</td>
<td>26:42:00</td>
<td>1692</td>
<td>2630</td>
</tr>
<tr>
<td>Coma_MOS05</td>
<td>12:57:12.0</td>
<td>29:00:36</td>
<td>1686</td>
<td>2606</td>
</tr>
<tr>
<td>Coma_SPECA</td>
<td>12:57:36.0</td>
<td>27:27:00</td>
<td>1350</td>
<td>1350</td>
</tr>
<tr>
<td>Coma_MOS06</td>
<td>13:00:00.0</td>
<td>29:00:36</td>
<td>1693</td>
<td>3524</td>
</tr>
<tr>
<td>Coma_MOS08</td>
<td>13:01:58.0</td>
<td>28:36:00</td>
<td>1693</td>
<td>2598</td>
</tr>
<tr>
<td>Coma_MOS09</td>
<td>13:02:46.0</td>
<td>27:48:00</td>
<td>1696</td>
<td>1696</td>
</tr>
<tr>
<td>Coma_MOS10</td>
<td>13:04:00.0</td>
<td>29:36:54</td>
<td>1694</td>
<td>2595</td>
</tr>
<tr>
<td>Coma_MOS11</td>
<td>13:05:19.2</td>
<td>28:18:00</td>
<td>1699</td>
<td>1699</td>
</tr>
<tr>
<td>Coma_MOS12</td>
<td>13:06:14.4</td>
<td>27:18:00</td>
<td>1701</td>
<td>2604</td>
</tr>
</tbody>
</table>

Figure 1. The Coma cluster region analysed in this paper. Red dots show all the NUV sources with $m_{AB}$(NUV) < 21.5. Black contours indicate the Coma X-ray emission as observed by XMM and circles indicate the apertures used to determine the LF in this paper. The dotted circle is at the virial radius of Coma $\sim$2.9 Mpc, as determined by Lokas & Mamon (2003).

The GALEX resolution is not sufficient to accurately separate stars and galaxies. To address this issue we matched the GALEX catalogue against the Sloan Digital Sky Survey Data Release 6 (SDSS-DR6) (Adelman-McCarthy et al. 2008) observations of the Coma cluster. Stellar objects and artefacts in the SDSS plates are sometimes erroneously present in both star and galaxy SDSS catalogue. For this reason, firstly we cross-matched our sample with the SDSS stellar catalogue using a search radius of 6 arcsec; secondly we cross-correlated the remaining UV detections with the SDSS galaxy catalogue, using the same impact parameter. Finally we cross-matched our catalogue with NASA/IPAC Extragalactic Database (NED) in order to obtain additional redshifts for the UV-selected sample. A total number of 762 galaxies in FUV and of 1640 galaxies in NUV with $m_{AB} < 21$ mag have been detected in the $\sim$9 deg$^2$ analysed in this paper.
The NUV and FUV luminosity functions

The determination of the cluster UV LF requires a reliable estimate of the contribution from background/foreground objects to the UV counts. This can be accurately obtained for $m_{AB} \leq 18.0$, since at this limit our redshift completeness is $\geq 90$ per cent (see Fig. 2).

The redshift completeness drops rapidly at fainter magnitudes, thus requiring a statistical estimate of the contamination. Three different methods have been here adopted for the computation of the cluster LFs. The first one is based on the statistical subtraction of field galaxies, per bin of UV magnitude, that are expected to be randomly projected on to the cluster area, as derived by Xu et al. (2005). In the second one, cluster members are identified on morphological grounds. Although sometimes subjective, in particular in the case of faint and/or compact objects, this technique can be applied to nearby clusters of galaxies when the redshift coverage is inadequate (e.g. Sandage & Tammann 1985; Ferguson & Sandage 1988; Trentham & Tully 2002). To do so, we have visually inspected SDSS–RGB images and we have used the criteria discussed in Binggeli, Sandage & Tammann (1985) to discriminate between Coma and background galaxies. In ambiguous cases, we adopted the colour cut described below to select possible cluster members. The third one is the completeness-corrected method proposed by De Propris et al. (2003). This method is based on the assumption that the spectroscopic sample (membership confirmed spectroscopically) is ‘representative’ of the entire cluster. That is, the fraction of galaxies that are cluster members is the same in the (incomplete) spectroscopic sample as in the (complete) photometric one. However, this is not always the case. For example, redshift estimates of galaxies in the Coma region are obtained from different sources and the selection criteria are not always well defined. Moreover, spectroscopic targets are often selected according to their colour (e.g. Mobasher et al. 2003) or narrow-band emission (e.g. Iglesias-Páramo et al. 2002) in order to maximize the number of members in the final sample, thus overestimating the real fraction of members. It is therefore likely that the membership fraction in our sample represents an upper limit to the real value (in particular at faint apparent magnitudes where the redshift completeness and the fraction of members are low). The only way to reduce this bias is to try to increase the completeness of our sample by excluding those galaxies which are likely to be background objects. We thus used stellar population synthesis models and a (NUV − $i$) versus ($g − i$) colour diagram to reject galaxies with colours not representative of the Coma members ($4000 \text{ km s}^{-1} < V < 10000 \text{ km s}^{-1}$).

In Fig. 3 we compare the distributions of Coma confirmed members (squares), background galaxies (triangles) and galaxies without redshift estimate (circles) with the spectral energy distribution (SED) library typical of nearby galaxies described in Cortese et al. (2008), whose models span a wide range of ages, metallicities and dust attenuations. Every confirmed Coma cluster member lies within the parameter space admitted by the models, allowing us to reduce the number of possible members in our sample. We excluded all galaxies with $(g − i) < 0.34(\text{NUV} − i)$ (dashed line in Fig. 3), reducing our sample from 1640 to 853 galaxies in NUV and from 762 to 500 galaxies in FUV. However, the improvement in the redshift completeness for our sample appears significant only in the last bin, as shown by the dotted and dashed lines in Fig. 2.

In addition, it is important to remember that the spectroscopic sample is optically selected whereas the photometric sample is UV selected. Therefore, objects bright in optical but faint in UV, like...
elliptical galaxies, have a completeness distribution significantly different from that of spiral galaxies (see Fig. 2, middle and bottom panel). Thus we decided to modify the completeness-corrected method and to treat separately quiescent (Q) and SF galaxies. In particular, we used an observed colour $\text{NUV} - r = 4.5$ mag (which corresponds to a specific star formation rate $\text{SSFR} \sim 10^{-11.5}$ yr$^{-1}$) to separate blue/SF galaxies from quiescent objects. For each type and magnitude bin $i$, we then counted the number of cluster members $N_M$ (i.e. galaxies with velocity in the range $4000 < V < 10000$ km s$^{-1}$), the number of galaxies with a measured recessional velocity $N_Z$ and the total number of galaxies $N_T$. The completeness-corrected number of cluster members for each type in each bin is

$$N_i(\text{SF, Q}) = \frac{N_M(\text{SF, Q}) \times N_T(\text{SF, Q})}{N_Z(\text{SF, Q})}.$$  

(1)

$N_T$ is a Poisson variable and $N_M$ is a binomial variable (the number of successes in $N_Z$ trials with probability $N_M/N_Z$) and the errors associated with $N_i(\text{SF, Q})$ are given by

$$\frac{\delta N_i(\text{SF, Q})^2}{N_i(\text{SF, Q})} = \frac{1}{N_T(\text{SF, Q})} + \frac{1}{N_M(\text{SF, Q})} - \frac{1}{N_Z(\text{SF, Q})}.$$  

(2)

The total number of galaxies in each bin of the LF and its uncertainty are therefore

$$N_i = N_i(\text{SF}) + N_i(\text{Q}), \delta N_i = \sqrt{\delta N_i(\text{SF})^2 + \delta N_i(\text{Q})^2}.$$  

(3)

The NUV and FUV LFs obtained using the three methods described above, within different circular apertures centred on the cluster centre, are shown in Fig. 4. All methods show a reasonable agreement for a cluster-centric distance $R < 1^\circ$ (1.67 Mpc) whereas for larger areas the statistical subtraction of the field provides number counts significantly lower than the other two techniques. This is likely due to the fact that, for scales larger than $\sim 1^\circ$, the cluster overdensity in UV starts to disappear. No significant difference is observed between the completeness-corrected and the morphology-based methods. For these reasons, in the following we will measure the UV LF of Coma using the completeness corrected method only.$^1$

The NUV and FUV LFs obtained for the whole $\sim 9$ deg$^2$ here investigated are shown in Fig. 5 (upper panels). The LFs are not well described by a Schechter function (Schechter 1976), for $M(\text{UV}) < -17$ mag making difficult an accurate estimate of $M_*$, However, in order to compare our results with previous determinations of field and cluster UV LFs, we fitted our points with a Schechter function by minimizing the $\chi^2$. The best-fitting values for $M_*$ and $\alpha$ so obtained are presented in Table 2 and compared with previous determinations for local field and cluster galaxies. The $\chi^2$ contours are shown in the bottom panels of Fig. 5. The faint-end slopes of the Coma LFs are consistent with those recently obtained by Cortese et al. (2005) in Abell 1367 using GALEX observations (see Table 2) and with previous determination of the UV LF in nearby clusters based on balloon-borne experiments (Cortese et al. 2003a,b). On the contrary, Coma shows a bright end $\sim 1.5$ mag fainter than the one observed in Abell 1367. This apparently remarkable difference is due to the presence in Abell 1367 of a few galaxies with enhanced star formation, among which the ultraluminous UV galaxy UGC6697 (Gavazzi et al. 2001; Gil de Paz et al. 2007). Excluding the brightest point in the LF of Abell 1367 we find a good agreement between the LFs of the two clusters. When compared to the field (Wyder et al. 2005, see Table 2), the Coma UV LFs have approximately the same bright end ($M_*$) but show a considerably steeper faint-end slope ($\alpha$), in particular in the NUV (see Fig. 5). By dividing our sample into SF and quiescent galaxies according to their observed NUV $- r$ colour (see Fig. 6, filled symbols), it clearly appears that the difference between field and cluster LFs is mainly due to the significant contribution of quiescent systems at low UV luminosities. These are not only elliptical and lenticular galaxies (e.g. Boselli et al. 2005; Donas et al. 2007), but also passive spirals whose star formation has been recently quenched by the cluster environment (e.g. Boselli et al. $^1$)

1 All the results presented in the following do not change if the background subtraction (for $R < 1^\circ$) or the morphology are used to estimate the LFs.
Figure 5. Upper panels: The FUV (left-hand side) and NUV (right-hand side) LFs for the whole ∼9 deg² observed by GALEX in the Coma cluster. The dotted lines show the GALEX LFs for local field galaxies as obtained by Wyder et al. (2005). The field LFs have been normalized in order to match the total number of galaxies detected for $M \leq -14$ mag. The solid lines indicate the best-fitting Schechter functions to the data. Lower panels: Contour plots of the 68, 95 and 99 per cent confidence levels for the values of $M_*$ and $\alpha$. The results obtained for the field LF and for Abell 1367 (Cortese et al. 2005) are indicated with an asterisk and a triangle, respectively.

Table 2. Best-fitting parameters for the NUV and FUV LFs.

<table>
<thead>
<tr>
<th>Band</th>
<th>Sample</th>
<th>$M_*$</th>
<th>$\alpha$</th>
</tr>
</thead>
<tbody>
<tr>
<td>NUV</td>
<td>Coma</td>
<td>$-18.50 \pm 0.50$</td>
<td>$-1.77^{+0.16}_{-0.13}$</td>
</tr>
<tr>
<td>NUV</td>
<td>Abell 1367</td>
<td>$-19.77 \pm 0.42$</td>
<td>$-1.64 \pm 0.21$</td>
</tr>
<tr>
<td>NUV</td>
<td>Field</td>
<td>$-18.23 \pm 0.11$</td>
<td>$-1.16 \pm 0.07$</td>
</tr>
<tr>
<td>FUV</td>
<td>Coma</td>
<td>$-18.20 \pm 0.80$</td>
<td>$-1.61^{+0.19}_{-0.25}$</td>
</tr>
<tr>
<td>FUV</td>
<td>Abell 1367</td>
<td>$-19.86 \pm 0.50$</td>
<td>$-1.56 \pm 0.19$</td>
</tr>
<tr>
<td>FUV</td>
<td>Field</td>
<td>$-18.04 \pm 0.11$</td>
<td>$-1.22 \pm 0.07$</td>
</tr>
</tbody>
</table>

aThis paper; bCortese et al. (2005); cWyder et al. (2005).

2006, 2008), as shown in Fig. 7. In fact, if we use morphology instead of colour to separate late and early-type galaxies (shaded and dotted regions in Fig. 6), the number of red galaxies contributing to the faint end of the LF decreases. The presence of passive spirals is mainly evident in NUV; e.g. for $M$(NUV) $> -16$ mag nearly half of the cluster UV emitting objects are not forming stars, ∼70 per cent early-types and ∼30 per cent quiescent late-type galaxies. This is probably due to the different stellar populations responsible for the NUV and FUV emission in quiescent systems (Boselli et al. 2005; Donas et al. 2007). On the contrary, it is interesting to note that the UV LFs for Coma blue/SF galaxies are not significantly different from that of the field. We can therefore conclude that (1) the faint-end slope of the UV LFs are steeper for Coma than for the local field and (2) this difference is almost entirely due to the contribution of red/quiescent galaxies at low UV luminosities, consistently with previous studies (Cortese et al. 2003a, 2005).

Finally, we note that the lack of GALEX observations in the central ∼0.26 deg² of Coma, mentioned in Section 2, does not affect our conclusions. In fact, Cortese et al. (2003b) used UV (2000 Å) FOCA (Milliard, Donas & Laget 1991) observations to determine the UV LF of the central ∼0.7 deg² of Coma for $M_*$($2000 \text{ Å}$) $< -15.5$ mag. They find a faint-end slope ($\alpha$ $\sim -1.65 \pm 0.30$), consistent with our results. Moreover, Cortese et al. (2003a) combined FOCA observations of the central part of Coma and Abell 1367 with FAUST (Lampton, Deharveng & Bowyer 1990) observations of the Virgo cluster and determined a composite local UV LF for galaxy clusters, finding a shape of the LF very similar to the one obtained in this paper ($M_* = -18.79 \pm 0.40$ mag, $\alpha$ $\sim -1.50 \pm 0.10$).

4 THE RADIAL VARIATION OF THE UV LUMINOSITY FUNCTIONS

The wide area surveyed by GALEX can be used to investigate the variation in the shape of the LF as a function of the cluster-centric distance. Ideally, this should be done by comparing the LF obtained in different circular annuli centred on the cluster centre. In reality,
Figure 6. The contribution of different galaxy types to the FUV (left-hand panel) and NUV (right-hand panel) LFs. The LFs of SF and quiescent galaxies, classified according to their observed NUV $-$ r colour, are indicated with squares and circles, respectively. The LFs of late and early-types, classified from visual inspection of SDSS images, are the shaded and dotted regions, respectively. The dotted lines show the GALEX LFs for local field galaxies. The field LFs have been normalized in order to match the total number of cluster SF galaxies brighter than $M = -14$ mag.

Figure 7. SDSS colour images of quiescent disc galaxies in our UV-selected sample. Each image is 1.2 arcmin per side, corresponding to $\sim 33$ kpc at the distance of Coma.

this is not always possible given the poor statistics of current samples. For example, in our case the small number of galaxies at high UV luminosities and the uncertainty in the membership fraction at low luminosities make impossible any estimate of the differential LF. Any difference observed could in fact only reflect statistical fluctuations in our sample. For these reasons, we decided to investigate the spatial variation of the LFs focusing our attention on the integral UV LFs, i.e. calculated within different cluster-centric apertures. Inevitably, by using this method, any difference in the shape of the LFs will be less evident since the central part of the cluster will contribute to all the LFs. The FUV and NUV LFs obtained within circular regions having radius 0.5, 0.75, 1', 1.5'(\~0.84, 1.25, 1.67, 2.5 Mpc or \~0.3, 0.4, 0.6, 0.9 virial radii) and for the whole sample are shown in Fig. 8. At all radii the LFs in both NUV and FUV are inconsistent with a Schechter function. Our data are better described by a Gaussian at bright luminosities and a power law at faint luminosities, complicating the quantification of any difference between LFs at various radii. As discussed in the previous section, this is the result of the different contribution of SF objects and quiescent galaxies to the UV LFs. Excluding a variation in the normalization, no significant change is observed in the shape for different cluster-centric radii. The apparent flattening in the faint-end slope at smaller radii marginally visible in NUV is likely due to statistical fluctuations. We note that a fit of the faint-end slope in the range $-16 \leq M_{\text{UV}} \leq -14$ mag would result in a significant steepening with increasing radius in NUV (from $\alpha \sim -1.2$ to $\sim -1.8$) and flattening with increasing radius in FUV (from $\alpha \sim -2.5$ to $\sim -1.7$). This points out the danger of blindly fitting faint-end slopes to compare the shape of the LFs obtained with different samples.

In Fig. 9 we show the integrated NUV and FUV LF as a function of the cluster-centric distance for SF and quiescent objects separately. Also in this case we find that, within statistical errors, the shape of the LFs does not vary with cluster-centric distance. Similar results are obtained if we divide our sample in late and early-types accordingly to their morphology (not shown). Only by integrating the LFs in the range in the range $-17 < M_{\text{NUV}} < -14$ mag [$-16.5 < M_{\text{FUV}} < -14$ mag], we find that the ratio of the UV emission of SF to that of quiescent systems monotonically increases...
Figure 8. The NUV (upper panel) and FUV (bottom panel) LFs of the Coma cluster within different circular apertures: 0.5° (circles), 0.75° (empty triangles), 1° (filled triangles), 1.5° (stars) and the whole region (squares). The dotted line shows the local field LF, normalized in order to match the total number of cluster galaxies brighter than $M = -14$ mag.

We can therefore conclude that, at least in the range $0.5 \leq R \leq 2°$, the shape of the integrated NUV and FUV LFs of the Coma cluster remains constant. We remind the reader that our results are only valid for integrated LFs and that the small number statistic makes impossible any quantification of the shape of the differential UV LFs.

5 THE EXTINCTION CORRECTED LF AND SFR OF THE COMA CLUSTER

Previous determinations of UV LFs in all environments and at all redshifts have never taken internal extinction corrections into account. This is mainly due to the fact that methods for estimating dust attenuation usually require FIR photometry, quite rare for large samples, and are calibrated only on starbursts and active SF galaxies. Only recently, Cortese et al. (2006a, 2008) have provided UV dust attenuation correction recipes, valid for galaxies with different SFHs. These recipes only require two broad-band colours and can be used to estimate an UV LF corrected for dust attenuation and to determine the total SFR in Coma. Since UV dust attenuation is important only for SF galaxies, in the following we restrict our analysis to the subsample including only blue galaxies (i.e. $NUV - r \leq 4.5$). For each galaxy, we determined $A(FUV)$ and $A(NUV)$ using the observed $FUV - NUV$ and $NUV - i$ colours as described in Cortese...
et al. (2008). We then corrected the NUV and FUV magnitudes and determined the NUV and FUV LF using the completeness corrected method described in Section 3. The LFs obtained for the whole sample are shown in Fig. 10: after dust attenuation correction, the LFs shift to brighter luminosity by ~1.5 mag in FUV and ~1 mag in NUV, but their shapes do not change significantly. We also determined the unreddened LFs within different apertures and found no significant radial variation in their shapes (not shown).

By integrating the observed and unreddened LFs within the magnitude range investigated here ($M \leq -14$ mag), we determined the fraction of UV luminosity absorbed by dust. The total UV dust attenuations of the Coma cluster are $A_{\text{NUV}} = 0.84 \pm 0.28$ mag and $A_{\text{FUV}} = 1.25 \pm 0.25$ mag, implying that two-thirds (half) of the attenuation correction is due to dust absorption.

The attenuations in the two GALEX bands have been obtained independently, without any assumption on the shape of the extinction curve.

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Figure 9. The NUV (right-hand side) and FUV (left-hand side) LFs for quiescent systems (upper panel) and SF galaxies (bottom panel) within different circular apertures. Symbols are as in Fig. 8.

Figure 10. Effect of dust attenuation corrections on FUV (left-hand panel) and NUV (right-hand panel) LFs of SF galaxies. Empty and filled symbols show the LFs before and after internal extinction correction, respectively.
the light emitted in FUV (NUV) by Coma galaxies is absorbed by dust. No variation in the UV dust attenuation with the distance from the cluster centre is observed. These values are consistent with the total FUV dust attenuation observed in the local field $A_{\text{FUV}} = 1.29^{+0.35}_{-0.32}$ mag (Takeuchi, Buat & Burgarella 2005), suggesting that the dust properties of blue/SF galaxies in Coma are not significantly different from those of field SF systems.

The unreddened UV LFs of SF galaxies can be also used to estimate the total SFR of the Coma cluster. To do so we integrated the LFs in the observed range and then converted the total luminosity into total SFR following Iglesias-Páramo et al. (2006):

$$\text{SFR(FUV)}(\mathcal{M}_\odot \text{yr}^{-1}) = \frac{L(\text{FUV})(\text{erg s}^{-1})}{3.83 \times 10^{33}} \times 10^{-9.51},$$

$$\text{SFR(NUV)}(\mathcal{M}_\odot \text{yr}^{-1}) = \frac{L(\text{NUV})(\text{erg s}^{-1})}{3.83 \times 10^{33}} \times 10^{-9.33},$$

where $L(\text{NUV})$ and $L(\text{FUV})$ are the total FUV and NUV luminosities (corrected for dust attenuation) of the Coma cluster.\footnote{We note that, in reality, the calibration constant for the UV luminosity versus SFR relation varies by up to a factor of $\sim 1.5$ with the stellar metallicity; i.e. the UV luminosity increases at lower metallicities (Bicker & Fritze-v. Alvensleben 2005). Unfortunately we do not have accurate metallicity information to take this effect into account.} The total SFR of the Coma cluster results\footnote{We consider the average between the total SFR obtained in NUV and FUV.} $\text{SFR} = 90 \pm 16 \mathcal{M}_\odot \text{yr}^{-1}$, consistent with the SFR (FIR) = 97 $\mathcal{M}_\odot \text{yr}^{-1}$ recently obtained by Bai et al. (2006) using FIR observations. A similar value is obtained if the total SFR is estimated combining the FIR and UV luminosities (Bell 2003; Hirashita, Buat & Inoue 2003; Iglesias-Páramo et al. 2004):

$$\text{SFR} = \text{SFR}_{\text{NUV}}^0 + (1-\eta)\text{SFR(FIR)},$$

where $\eta$ accounts for the IR cirrus emission (here we assume $\eta = 0.3$, Bell 2003) and $\text{SFR}_{\text{NUV}}^0$ is obtained from equation (4) but using the observed NUV luminosity.

This result suggests that totally obscured star formation (i.e. not at all visible in UV), if present, represents a very small fraction of the total star formation in Coma, in agreement with IRAS (e.g. Bica & Giovanelli 1987; Boselli & Gavazzi 2006) and Spitzer (e.g. Muzzin et al. 2008; Saintonge, Tran & Holden 2008) FIR observations of local clusters of galaxies.\footnote{In the near future, when Spitzer and Herschel data will become publicly available, it will also be possible to test this result on a galaxy by galaxy basis.} In Fig. 11 we show the variation of the integrated SFR and SFR surface density as a function of cluster-centric distance and we compare our results with previous estimates based on FIR observations (Bai et al. 2006) and Hα narrow-band imaging (Iglesias-Páramo et al. 2002). As expected, the integrated SFR increases by a factor of $\sim 4$, and the integrated SFR surface density decreases by a factor of $\sim 5$, from the centre to the outskirts of Coma. Although the total SFR obtained integrating over the whole area observed by GALEX is roughly consistent with the value obtained from FIR observations, the SFR obtained in UV for $R < 1^\circ$ is $\sim 2-3$ times lower than that obtained in similar areas from FIR and Hα data. This difference is likely due to the absence of GALEX observations in the central $\sim 0.26 \text{deg}^2$, as supported by the similar SFR surface density obtained at small radii in UV, FIR and Hα (see Fig. 11).

The contribution of Coma-like clusters of galaxies to the SFR density in the local universe can be computed if the cluster volume density is known. This value varies significantly with the properties of the adopted cluster sample (e.g. Scarlomella et al. 1991; Mazure et al. 1996; Bramel, Nichol & Pope 2000; De Propris et al. 2002), but we can roughly assume it to lie in the range $3-9.5 \times 10^{-6} \text{Mpc}^{-3}$. Thus, the total SFR density in local clusters is $\rho_{\text{SFR}}(cl) \sim 2.5-8.5 \times 10^{-4} \mathcal{M}_\odot \text{yr}^{-1} \text{Mpc}^{-3}$, i.e. between 1.5 and 5 per cent of the total SFR density in the local universe $\rho_{\text{SFR}} \sim 10^{-1.80 \pm 0.16} \mathcal{M}_\odot \text{yr}^{-1} \text{Mpc}^{-3}$ (Hanish et al. 2006). This value mainly depends on the density of local clusters and does not provide any information on the different SFH of cluster and field galaxies. A more interesting exercise is the comparison of the SSFR of Coma and the field. In order to quantify the total mass for our sample, we converted the unreddened $g-i$ band magnitude to stellar mass using the $g-i$ colour following Bell et al. (2003) and estimated the stellar mass distribution by using the completeness-corrected method described in Section 3. The stellar mass distributions for the SF and quiescent galaxies in the UV-selected sample are shown in Fig. 12. The total stellar mass of our sample results $M_{\text{star}}(\text{total}) = 10^{13.1 \pm 0.10} \mathcal{M}_\odot$, great part of which is assembled in quiescent systems; i.e. $M_{\text{star}}(Q) = 10^{13.07 \pm 0.09} \mathcal{M}_\odot$ and $M_{\text{star}}(\text{SF}) = 10^{12.26 \pm 0.20} \mathcal{M}_\odot$. In Fig. 12, it clearly emerges that the vast majority of the SF galaxies are low-mass ($\leq 10^{10} \mathcal{M}_\odot$) objects. Moreover, by estimating the UV LF for low- and high-mass objects it emerges that $\sim 66 \pm 10$ per cent of the total Coma SFR is taking place in galaxies less massive than $\sim 10^{10} \mathcal{M}_\odot$, confirming previous observational evidences that downsizing is also present in galaxy clusters (e.g. Gavazzi, Pierini & Boselli 1996; Boselli et al. 2001; Gavazzi et al. 2002; De Lucia et al. 2004b; Kodama et al. 2004; Poggianti et al. 2004; Boselli & Gavazzi 2006; Poggianti et al. 2006). The integrated SSFR for the Coma cluster is $\sim 10^{-11.18 \pm 0.13} \text{yr}^{-1}$, significantly lower than the typical SSFR observed in the local field, e.g., SSFR$_\text{field} \sim 10^{-10.17} \text{yr}^{-1}$ (Salim et al. 2007) or SSFR$_\text{field} \sim 10^{-10.55 \pm 0.20} \text{yr}^{-1}$ if we assume $\rho_{\text{SFR}} \sim$
$10^{-1.30\pm 0.16} M_\odot \text{yr}^{-1} \text{Mpc}^{-3}$ (Hanish et al. 2006) and a local stellar mass density $\sim 10^{0.76\pm 0.12} M_\odot \text{Mpc}^{-3}$ (Pérez-González et al. 2008). This difference is likely a combination of the morphology–density and star formation–density relations: massive elliptical galaxies are more overdense in clusters (Dressler 1980; Whitmore et al. 1993) but also, for the same morphological type, cluster galaxies have a lower star formation activity than field objects (Gavazzi et al. 2006). Finally, it is interesting to note that the SSFR of blue/SF galaxies is $\sim 10^{-10.31\pm 0.22} \text{yr}^{-1}$, i.e. consistent with that of the field, suggesting that Coma SF galaxies have not been significantly affected by the cluster environment.

### 5.1 Bias and selection effects

The results presented above might be affected by two different selection effects. First, we are underestimating the total SFR of Coma since we lack observations for the cluster core. Secondly, the selection in UV adopted here is strongly biased against red massive systems (e.g. see Fig. 13), suggesting that we could also underestimate the total stellar mass of Coma. We can quantify the first bias by comparing our results with previous determinations of the total SFR (Iglesias-Páramo et al. 2002; Bai et al. 2006; see fig. 11) that include the cluster centre. In the central $\sim 0.5$, we are missing $\sim 30 M_\odot \text{yr}^{-1}$, implying that the total SFR of Coma presented above is underestimated by a factor of $\sim 1.3$. The total SFR density in Coma-like clusters would therefore increase to $\sim 2$–7 per cent of the total SFR density in the local universe. To address the second bias, we can estimate the total stellar mass of the Coma cluster from its total mass (baryonic + dark matter) $M_{\text{tot}} \sim 1.4 \times 10^{15} M_\odot$ (Lokas & Mamon 2003), assuming a mass-to-light ratio $M_{\text{star}}/L_K = 75 M_\odot/L_\odot$ (Rines et al. 2004) and a stellar mass-to-light ratio $M_{\text{star}}/L_K$ ratio of $\sim 1 M_\odot/L_\odot$ (Bell et al. 2003). We obtain a total stellar mass of $M_{\text{star}} \sim 10^{11.3} M_\odot$, i.e. $\sim 1.5$ times higher than the value obtained for our sample. The fact that the underestimates in the total SFR and stellar mass are comparable implies that the SSFR of the Coma cluster decreases only by a factor of $\sim 1.1$, well within the observational errors. We can therefore conclude that the lack of

### 6 DISCUSSION

#### 6.1 Why are the field and cluster LFs of SF galaxies so similar?

Our analysis of GALEX observations of the Coma cluster provides definitive evidence that the shape of the UV LF is not universal but depends on the local environment (e.g. see Figs 5 and 6). Coma has a steeper faint-end due to the high number density of massive, quiescent galaxies and its total SSFR is significantly lower than that of the local field. This is a direct consequence of the morphology–density and star formation–density relations: quiescent galaxies are in fact ellipticals, lenticulars and passive spirals whose star formation has been recently quenched by the cluster environment (see Fig. 7). Only if we restrict our attention to blue/SF cluster galaxies, the UV LFs and SSFR are not significantly different from those observed in low-density environments, in agreement with recent studies of nearby (e.g. Balogh et al. 2004; Cortese et al. 2005) and high-redshift (e.g. Cooper et al. 2008; Finn et al. 2008) clusters. These results could be interpreted as a strong evidence that the environment does not significantly affect the cosmic SFH of galaxies (Cooper et al. 2008; Finn et al. 2008). However, this interpretation is inconsistent with the plethora of evidences showing the strong effects of the environment on the evolution of cluster galaxies (e.g. see Boselli & Gavazzi 2006 and references therein). Here we show how the apparent independence of the properties of SF galaxies from the environment is, in reality, only a selection effect.

Our definition of blue/SF galaxies is based on colour or SSFR (i.e. $NUV - r \leq 4.5$, SSFR $\geq 10^{-11.5} \text{yr}^{-1}$), as often happens for large samples, implying that those galaxies which have been strongly affected by the environment are automatically excluded. This selection effect is evident from Fig. 13, where the relation between

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**Figure 12.** The stellar mass distributions for quiescent (circles) and SF galaxies (squares) in our UV-selected sample.

**Figure 13.** The SSFR as a function of the stellar mass for Coma confirmed members. Circles and squares indicate blue/SF and red/quiescent galaxies, respectively. The dashed line shows our detection limit in UV. The best-fitting relation for field galaxies (solid line) and its scatter (dashed line) as obtained by Salim et al. (2007) is superposed.

**GALEX** observations for the core of Coma does not significantly bias our results.

SSFR and stellar mass for confirmed cluster members in our sample is shown. The solid line shows the typical relation (and its 1σ uncertainty, dotted lines) for field galaxies as obtained by Salim et al. (2007), while the dashed line represents our sensitivity limit (SFR $\sim 0.04\ M_\odot\ yr^{-1}$). For $M_{\text{star}} < 10^9\ M_\odot$ we detect only healthy SF objects and, if a galaxy had its star formation quenched by the environment, it will not be detected by our observations. By comparison, for $M_{\text{star}} > 10^9\ M_\odot$, SF galaxies have to lie within $\sim 3\sigma$ from the field relation, otherwise their NUV $- r$ colour would be too red and they would be classified as quiescent objects. A clear example is represented by NGC4569 in the Virgo cluster, the prototype of anaemic spiral. Its star formation activity has been very recently (<500 Myr ago) quenched by ram pressure stripping (Vollmer et al. 2004; Boselli et al. 2006) and its colour (NUV $- r$) is already almost as red as that of a quiescent galaxy. Passive spiral galaxies are in fact already included in our sample of quiescent systems, as shown in Figs 6 and 7.

In order to demonstrate the validity of this interpretation, we investigated the effects of the environment on the shape of the UV LF for SF galaxies using a very simple toy model. We started with an infalling population of $\sim 1200$ galaxies with $M(\text{FUV}) < -14$ mag, reproducing the field UV LF. For each infalling galaxy, the SFR [i.e. $M(\text{FUV})$] is instantaneously quenched by the cluster environment. The amount of quenching (SFR before/SFR after) is randomly chosen, assuming a uniform distribution. Two different scenarios have been considered, depending on the maximum amount of quenching allowed. In scenario A, the maximum amount of quenching is SFR before/SFR after = 100, independently of galaxy luminosity. In scenario B, the maximum amount of quenching decreases linearly with galaxy luminosity: i.e. from SFR before/SFR after = 100 for $M(\text{FUV}) = -14$ mag to SFR before/SFR after = 10 for $M(\text{FUV}) = -18$ mag. By construction, both models are able to reproduce the parameter space occupied by our sample in Fig. 13. For each scenario, we then reconstruct the cluster UV LF including only those galaxies still matching our definition of SF object, i.e. SSFR $> 10^{-11.5}\ yr^{-1}$ and SFR $> 0.04\ M_\odot\ yr^{-1}$. 1000 simulated LFs have been so obtained and the average result, normalized to the total number of blue galaxies in our sample, is shown in Fig. 14. We note that both scenarios are likely to overestimate the real variation of the LF. In fact, we assume an instantaneous quenching whereas this time-scale is probably longer (see next section). Moreover, a quenching uniformly distributed in the range 1–100 is likely an overestimate of the average decrease in the SFR of cluster galaxies (Lewis et al. 2002; Gómez et al. 2003; Gavazzi et al. 2006). Even so, the variation in the shape of the LF after (filled area in Fig. 14) falling into the cluster does not appear very significant, supporting our interpretation. Therefore, even if the cluster is very efficient in quenching the SFR, a selection based on colour/SSFR automatically excludes those galaxies which have already been strongly affected by the environment. This selection effect, even more important at higher redshift where the dynamical range in stellar mass is usually smaller than ours, implies that cluster and field SF galaxies are characterized by a similar LF by construction, since we are only selecting healthy objects. The steepening of the UV LF for late-type galaxies, once quiescent spirals are included (shaded region in Fig. 6), provides additional support to our interpretation. Thus, the important result does not lie in the shape of the LFs but in the fact that we detect a significant fraction of healthy SF galaxies in the cluster environment. These galaxies likely represent a recent infalling population of field systems, not yet affected by the harsh environment of Coma. This scenario also explains the similar decrease in the SFR of field and cluster galaxies between $z \sim 1$ and 0 (Cooper et al. 2008; Finn et al. 2008); the vast majority of cluster SF galaxies are still infalling for the first time into the cluster centre and therefore their properties are still representative of the field.

6 We remind that we are only interested in the shape of the LF. After the quenching of the star formation the absolute normalization of the LF is significantly reduced.

6.2 The accretion rate of the Coma cluster

As argued in the previous section, the healthy SF galaxies observed in Coma appear to have just started their first dive into the cluster centre. Our observations could therefore allow us to quantify the mass accretion rate of SF galaxies from the field assuming that we...
know the exact time-scale necessary to quench the star formation. This is still an open issue but, assuming reasonable estimates for the typical Coma crossing time ($\sim 1.6 \times 10^8$ yr, Boselli & Gavazzi 2006), the time-scale associated with the various physical processes invoked to suppress star formation ($\sim 0.1$–1 Gyr; Abadi, Moore & Bower 1999; Poggianti et al. 1999; Quilis, Moore & Bower 2000; Vollmer et al. 2001; Shioya et al. 2002; Fujita 2004; Poggianti et al. 2004; Roediger & Hensler 2005; Boselli & Gavazzi 2006; Boselli et al. 2006; Cortese et al. 2007; Boselli et al. 2008) and the time-scale for the evolution of the UV luminosity ($\sim 10^9$ yr), we conclude that 1–3 Gyr (at most) should be sufficient to transform an infalling galaxy into a quiescent system. By combining this time-scale with the total stellar mass of the infalling galaxies we can infer a stellar mass accretion rate of $\sim (0.6$–1.8) $\times 10^{-12}$ M$_{\odot}$ Gyr$^{-1}$, roughly consistent with the value obtained by Adami et al. (2005), when converted to total mass using a total-to-stellar mass ratio $M_{\text{tot}}/M_{\text{star}} = 75$ (Rines et al. 2004). This implies that, if the mass accretion rate has not significantly changed during the age of the universe, the whole quiescent population could have formed by infalling galaxies in a Hubble time. In particular, from $z \sim 1$ until now the stellar mass in the red sequence has increased by a factor of 1.5–5 (depending on the real infalling rate$^5$), a value consistent with the cosmic growth of stellar mass (e.g. Bell et al. 2007; Brown et al. 2007; Cimatti et al. 2006; Zucca et al. 2006; Faber et al. 2007). We note that this simple picture is consistent with the large accretion rate in clusters at $z < 1$ predicted by numerical simulations (e.g. de Lucia et al. 2004a) and with the typical infalling rate obtained from observations (e.g. Ellingson et al. 2001; Andreon et al. 2006) and simulations (e.g. van den Bosch 2002; Berrier et al. 2008). In addition, a steady growth of Coma over cosmic time would not significantly modify the bright end ($M^*$) of the cluster stellar mass function, in agreement with observations of high-redshift clusters (e.g. Andreon 2006; de Propris et al. 2007).

The fact that red sequence galaxies are preferentially found in clusters of galaxies suggests that the cluster environment could play a non negligible role in the building up of the red sequence since $z \sim 1$. If so, we can expect a rapid decrease in the difference between the specific star formation of Coma and the field from $z \sim 1$ to 0. An accurate quantification of the mass accretion rate of Coma, and of its evolution with redshift, is therefore mandatory to determine whether or not the shape of the cosmic SFH depends on the environment.

In spite of their large uncertainty, these simple calculations show that the infall of healthy spirals from low-density environments into the centre of cluster of galaxies cannot only explain part of the morphology–density (Smith et al. 2005; Desai et al. 2007) and star formation–density relations observed in today clusters of galaxies, but could also easily account for the strong evolution in the stellar mass budget of the red sequence from $z \sim 1$ to 0.

### 7 SUMMARY AND CONCLUSIONS

In this paper, we have presented a study of the UV properties of the Coma cluster based on GALEX NUV and FUV observations covering $\sim 9$ deg$^2$ centred on the cluster core. Although the central $\sim 0.26$ deg$^2$ could not be observed, our analysis represents the widest and deepest UV investigation of a nearby cluster of galaxies made to date. Our main results are as follows.

(i) The Coma NUV and FUV LFs show a faint-end slope significantly steeper than the one observed in the field. This difference is more evident in NUV and it is due to the higher number density of massive quiescent/red galaxies (i.e. ellipticals, lenticulars and passive spirals) in Coma compared to the field. The contribution of quiescent galaxies to the total UV emission at low luminosities ($M_{\text{UV}} > -17$ mag) is larger in the cluster centre, however no significant variation in the shape of the UV LFs with cluster-centric distance is observed.

(ii) We estimated for the first time the UV LFs of SF galaxies corrected for internal dust attenuation. We showed that Coma-like clusters contribute only $< 7$ per cent of the total SFR density of the local universe. More interestingly the SSFR of Coma is $\sim 10^{-11.18 \pm 0.13}$ yr$^{-1}$, significantly lower than the integrated SSFR of the local universe. Approximately two-thirds of the whole SF in Coma is occurring in objects with $M_{\text{star}} < 10^{10}$ M$_{\odot}$, confirming that downsizing is also present in high-density environments.

(iii) The shape of the UV LF and the SSFR of blue/SF galaxies are consistent with those of the field, in agreement with previous works. We have shown that this similarity does not imply that the effects of the environment on the evolution of the cosmic SFH are negligible. On the contrary, these results are consistent with a scenario in which cluster SF galaxies are still infalling for the first time into the cluster centre. The stellar mass accretion rate of Coma results $\sim (0.6$–1.8) $\times 10^{-12}$ M$_{\odot}$ Gyr$^{-1}$. At this rate, the whole cluster could have easily formed from infalling galaxies accreted from the field in a Hubble time. More interestingly, a significant fraction of the population of lenticular and passive spirals observed today in Coma could arise from infalling galaxies accreted between $z \sim 1$ and 0, perhaps suggesting that the environment plays a significant role in the mass growth of the red sequence in the universe.

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$^7$ These values are obtained assuming that all the stellar mass in the central part of Coma (not included in this paper) lies in red sequence. These therefore represent lower limits to the real value.