Towards a complete census of active galactic nuclei in nearby galaxies: the incidence of growing black holes

A. D. Goulding, D. M. Alexander, B. D. Lehmer, and J. R. Mullaney

ABSTRACT
We investigate the local supermassive black hole (SMBH) density function and relative mass accretion rates of all active galactic nuclei (AGNs) identified in a volume-limited sample of infrared (IR) bright galaxies \((L_{\text{IR}} > 3 \times 10^8 L_\odot)\) to \(D < 15\) Mpc. A data base of accurate SMBH mass \(M_{\text{BH}}\) estimates is compiled from literature sources using physically motivated AGN modelling techniques (reverberation mapping, maser mapping and gas kinematics) and well-established indirect \(M_{\text{BH}}\) estimation methods (the \(M-\sigma\) and \(M_{\text{BH}}-L_{\text{Bul}}\) relations). For the three sources without previously published \(M_{\text{BH}}\) estimates, we use Two Micron All Sky Survey (2MASS) \(K\)-band imaging and GALFIT to constrain the bulge luminosities, and hence SMBH masses. In general, we find the AGNs in the sample host SMBHs which are spread over a wide mass range \([M_{\text{BH}} \approx (0.1–30) \times 10^7 M_\odot]\), but with the majority in the poorly studied \(M_{\text{BH}} \approx 10^6–10^7 M_\odot\) region. Using sensitive hard X-ray (2–10 keV) and mid-IR constraints we calculate the bolometric luminosities of the AGNs \(L_{\text{Bol,AGN}}\) and use them to estimate relative mass accretion rates. We use these data to calculate the volume-averaged SMBH growth rate of galaxies in the local Universe and find that the AGNs hosting SMBHs in the mass range \(M_{\text{BH}} \approx 10^6–10^7 M_\odot\) are dominated by optically unidentified AGNs. These relatively small SMBHs are acquiring a significant proportion of their mass in the present day, and are amongst the most rapidly growing in the local Universe (SMBH mass-doubling times of \(\approx 6\) Gyr). Additionally, we find tentative evidence for an increasing volume-weighted AGN fraction with decreasing SMBH mass in the \(M_{\text{BH}} \approx 10^6–10^8 M_\odot\) range. Overall, we conclude that significant mass accretion on to small SMBHs may be missed in even the most sensitive optical surveys due to absent or weak optical AGN signatures.

Key words: galaxies: active – galaxies: evolution – galaxies: nuclei – galaxies: Seyfert – infrared: galaxies.

1 INTRODUCTION
It is now well established that all massive galaxies \((M_* \approx 10^{10–12} M_\odot)\) in the local Universe harbour central supermassive black holes (SMBHs), with masses proportional to those of their stellar spheroids (hereafter bulge; e.g. Kormendy & Richstone 1995; Magorrian et al. 1998). Comparisons between the SMBH mass density in the local Universe and the total energy produced by active galactic nuclei (AGNs) across cosmic time have shown that these SMBHs were primarily grown through mass accretion events (e.g. Soltan 1982; Rees 1984; Marconi et al. 2004). The space density of high-luminosity AGNs appears to have peaked at higher redshifts than lower luminosity AGNs, suggesting that the most massive SMBHs \((M_{\text{BH}} \approx 10^9–10^9 M_\odot)\) grew first, a result commonly referred to as ‘AGN cosmic downsizing’ (e.g. Cowie et al. 2003; Ueda et al. 2003; McLure & Dunlop 2004; Hasinger, Miyaji & Schmidt 2005; Alonso-Herrero et al. 2008). Extrapolation of these results imply that the most rapidly growing SMBHs in the nearby Universe should be of comparatively low mass \((M_{\text{BH}} \approx 10^6 M_\odot)\).

To determine the characteristic masses of these growing SMBHs requires a complete census of AGN activity and SMBH masses in the local Universe.

Using data from the Sloan Digital Sky Survey (SDSS; York et al. 2000) in conjunction with the well-established SMBH–stellar velocity dispersion relation (hereafter \(M-\sigma\); e.g. Gebhardt et al. 2000; Tremaine et al. 2002), Heckman et al. (2004, hereafter H04) deduced that relatively low-mass SMBHs \((M_{\text{BH}} \approx 3 \times 10^7 M_\odot)\) residing in moderately massive bulge-dominated galaxies host the majority of present-day accretion on to SMBHs. However, the space density of SMBHs derived from the \(M-\sigma\) relation in the optical survey of H04 was limited by the spectral resolution of the SDSS \((\sigma_* > 70 \text{ km s}^{-1})\)
to SMBHs of $M_{\text{BH}} \gtrsim 3 \times 10^6 \, M_\odot$ (assuming the $M - \sigma$ relation of Gebhardt et al. 2000). Furthermore, due to attenuation of optical emission by dust, source selection and AGN classification at optical wavelengths will be biased against gas-rich, dust-obscured objects. These surveys are unlikely to include galaxies hosting the smallest bulges, and consequently the lowest mass SMBHs, and may therefore be missing a significant proportion of SMBH growth in the local Universe. Indeed, the nearby Scd galaxy, NGC 4945, hosting a low-mass SMBH ($M_{\text{BH}} \approx 1.4 \times 10^6 \, M_\odot$; Greenhill, Moran & Herrnstein 1997) only displays evidence for AGN activity in X-ray (Iwasawa et al. 1993) and mid-infrared (mid-IR) observations (Goulding & Alexander 2009). By contrast, the AGN in NGC 4945 (accreting at $\lesssim 30$ per cent of the predicted Eddington limit; Itoh et al. 2008) is completely hidden at optical wavelengths, and classified as a starburst galaxy. Clearly, using optical data alone, the intrinsic AGN properties of sources similar to NGC 4945 cannot be derived.

While optical emission-line diagnostics alone cannot reliably characterize the properties of a non-negligible fraction of the AGN population, they are readily identified at obscuration-independent wavelengths (e.g. X-ray, mid-IR). Hence, the identification of AGNs made at X-ray and mid-IR wavelengths complements traditional ultraviolet (UV)/optical methods to yield a more complete census of AGN activity. Indeed, using the high-resolution mid-IR spectrograph on board the NASA Spitzer Space Telescope (Spitzer-IRS), Goulding & Alexander (2009, hereafter GA09) found using the first complete volume-limited sample of all ($\approx 94$ per cent) local ($D < 15$ Mpc) bolometrically luminous galaxies ($L_{\text{IR}} > 3 \times 10^9 \, L_\odot$) that $\approx 50$ per cent of local AGNs are not identified in sensitive optical surveys. At least 30 per cent of these AGNs were previously identified as pure optical starburst galaxies, similar to NGC 4945 (i.e. not even otherwise known to be transition-type objects as defined by Kauffmann et al. 2003). Furthermore, $\approx 30$ per cent of the optically unidentified AGNs were found to reside in late-type spiral galaxies (Sc–Sd; e.g. similar to NGC 4945). Complimentary to this, from a heterogeneous sample of Palomar galaxies, Satyapal et al. (2007) and Satyapal et al. (2008) have also concluded that optically unidentified AGNs exist in some late-type spiral galaxies. With the inclusion of these new optically unidentified AGNs, it is natural to ask, what are the masses of local active SMBHs, what are their Eddington ratios, and hence, how rapidly are active SMBHs growing in the local Universe?

In this paper, we investigate the growth rates and space density of actively accreting SMBHs using the 17 AGNs identified in the volume-limited survey of GA09. Whilst the source statistics considered here are significantly smaller than those studies using the SDSS, this paper compliments that of H04 by including a relatively large number (given the considered small volume) of optically unidentified AGNs (10) which would not be reliably identified or characterized in the SDSS survey. Furthermore, by including a significant population of bolometrically luminous (but dust-obscured) late-type spiral galaxies (Sc–Sd) we are able to extend the SMBH density function to $M_{\text{BH}} < 3 \times 10^6 \, M_\odot$. As many of the late-type spiral galaxies host small galactic bulges, and hence lower mass SMBHs, particular attention is paid to obtaining accurate mass estimates for these SMBHs. Given their proximity, many of the sources in GA09 are well studied and have multiple estimates of SMBH mass ($M_{\text{BH}}$) from a variety of methods (i.e. reverberation mapping techniques, mapping of water maser spots, gas kinematical estimates, the $M - \sigma$ relation and correlation of $M_{\text{BH}}$ with the luminosity of the galactic bulge); in the following we discuss the relative accuracy of each SMBH mass estimate technique. Furthermore, to determine the relative mass accretion rates and hence average growth times of the SMBHs in our sample we require the best available estimates of the AGN bolometric luminosity ($L_{\text{bol,AGN}}$). Here we use two approaches: (1) for the AGNs with currently published data, we use high-quality well-constrained sensitive hard X-ray ($2 - 10$ keV) luminosities to directly measure $L_{\text{bol,AGN}}$ and (2) we accurately infer $L_{\text{bol,AGN}}$ using a well-constrained hard X-ray to high-ionization mid-IR emission-line relation.

In Section 2 we outline the construction and basic reduction analysis of the AGN sample derived from GA09. In Section 3 we present the SMBH mass estimates. For a minority of objects (three out of 17 AGNs) without published $M_{\text{BH}}$ estimates we outline the use of a bulge/disc decomposition method with Two Micron All Sky Survey (2MASS) $K$-band images, and following Marconi & Hunt (2003), we use the $M_{\text{BH}} - L_{\text{IR}}$ relation to estimate their SMBH masses. In Section 4 we use hard ($2 - 10$ keV) X-ray measurements and high-ionization mid-IR emission to estimate the intrinsic luminosity of the AGNs considered in our sample. Using our well-defined estimates for SMBH mass and AGN bolometric luminosity, we investigate the relative mass accretion rates of our sample of active SMBHs in Section 5. We use these estimates to provide new constraints on the volume-averaged SMBH growth rates in the local Universe. We further compare these results to the previous works of H04 and Greene & Ho (2007) by producing a local AGN population density function. Finally, in Section 6 we present our conclusions.
sample does not include galaxies from local overdensities such as the Virgo cluster at $D \sim 16$ Mpc (see Section A1 for a detailed analysis and validation of the considered space volume in this survey).

The *Spitzer*-IRS spectroscopic data presented in GA09 were reduced using a custom *irsa* pipeline which utilizes the *Spitzer* Science Center data processing packages *spice*, *irsclean* and *cubism*. For further detailed information on the reduction processes and spectral analyses see sections 2.1 and 2.2 of GA09 and references therein.

### 2.2 [Ne v] as an unambiguous AGN indicator

Due to the very high-ionization potential of [Ne v] (97.1 eV) we consider its detection coincident with the galactic nucleus in mid-IR spectroscopy to be an almost unambiguous identifier of AGN activity. Theoretically, Schauer & Stasinska (1999) have predicted that extremely hot O and B stars, in particular dense populations of Wolf–Rayet stars, may produce ionization spectra capable of exciting lines such as [O iv] (54.9 eV) and [Ne v]. However, observationally, GA09 found that even in extreme Wolf–Rayet galaxies, whilst [O iv] is clearly detected in these types of systems, [Ne v] emission remains absent to the detection limits of this survey. Complimentary to this, Hao et al. (2009) find from a *Spitzer*-IRS study of 12 blue compact dwarf galaxies that the mid-IR spectroscopy for eight of their sample contain [O iv] emission; however, none appears to be producing [Ne v].

Similarly, extreme starburst-driven shocks have also been predicted to excite some high-ionization lines such as [Ne v] (Allen et al. 2008); however, these require exceptionally high velocities, and based on the [Ne v] $\lambda 14.32$ [$\mu$m]–[Ne ii] $\lambda 12.81$ [$\mu$m] and [Ne ii] $\lambda 15.51$ [$\mu$m]–[Ne ii] $\lambda 12.81$ [$\mu$m] emission-line ratios presented in GA09, the AGNs in this paper are not consistent with shock models. Furthermore, AGNs which contain strong star formation contributions to their bolometric luminosity are often found to have relatively low mid-IR [Ne v]–[Ne ii] and [Ne iii]–[Ne ii] ratios (Armus et al. 2006; Satyapal et al. 2008; Dale et al. 2009; GA09). However, these ratios are not necessarily strong tracers of the intrinsic power of the AGN. Indeed, many Seyfert 2s are found to be hosted in galaxies whereby star formation dominates the IR spectral energy distribution (Weedman et al. 2005; Buchanan et al. 2006; Deo et al. 2007), thus yielding a low [Ne v]–[Ne ii] ratio (log-average $\approx 0.02$); however, the central source may still be extremely luminous at other energies, e.g. NGC 4945 is the most luminous local AGN at $E > 20$ keV (Done, Madejski & Smith 1996) and by contrast has a [Ne v]–[Ne ii] ratio of $\approx 0.01$.

The mid-IR spectra of 17 of the 64 galaxies ($\approx 27\%$ per cent) presented in GA09 were found to contain the [Ne v] $\lambda 14.32$ [$\mu$m] emission line, and hence, host AGN activity. These 17 sources are the main focus of the current paper (see Table 1).}

### 3 BLACK HOLE MASS DETERMINATION

#### 3.1 Archival data

To accurately determine the relative mass accretion rates and space density of active SMBHs in the local Universe requires reliable SMBH mass ($M_{\text{BH}}$) estimates for the 17 AGNs within our volume-limited sample. Here we outline the construction of the heterogeneous data base of the most reliable available SMBH masses for these AGNs, derived from a variety of archival sources and $M_{\text{BH}}$ estimation methods.

Many of the sources in the sample are late-type galaxies hosting relatively small bulges, and hence low-mass SMBHs. In such systems, SMBH mass estimates are often challenging to determine as (1) characteristically low velocity dispersions can be difficult to measure as they are often at the resolution limit of published observations, (2) modelling of the contamination from composite stellar populations can often lead to inconsistencies between published measurements and (3) SMBH mass relations are poorly constrained at $M_{\text{BH}} \sim 10^6 M_{\odot}$.

Given the varying degrees of accuracy associated with $M_{\text{BH}}$ measurements estimated from differing methodologies, we have chosen to prioritize the archival data (which we further expand on here) based upon two broad categories: (i) physically motivated AGN modelling techniques (i.e. reverberation mapping, water maser mapping and gas kinematics of the central engine) and (ii) indirect estimations from observational relations (the $M-\sigma_*$ and $M_{\text{BH}}$–$L_{\text{bol}}$ relations). The adopted $M_{\text{BH}}$ for the sources and their associated measurement methodologies are given in columns 12 and 13 of Table 1. The data base contains four SMBH measurements determined from physically motivated AGN modelling and 13 from indirect methods.

### 3.1.1 Direct SMBH mass constraints from reverberation mapping, maser mapping and gas kinematics

Under the assumption that the gas in the broad-line region (BLR) is virialized by the SMBH and the orbital motion of the gas is Keplerian, $M_{\text{BH}}$ estimations are possible through reverberation mapping techniques (Blandford & McKee 1982). The time-lag between changes in the AGN continuum flux and the response of the BLR is used to directly infer the size of the virial radius and hence the mass of the SMBH (for an in-depth review see Peterson 2001). To date, reverberation mapping is widely accepted to be the most reliable of $M_{\text{BH}}$ estimation methodologies (e.g. Peterson & Wandel 1999; Wandel, Peterson & Malkan 1999; Onken et al. 2003; Bentz et al. 2009b). Only a minority of the galaxies in the sample (NGC 1068, 4051 and 5033) are known to have detected BLRs. However, the BLRs of NGC 1068 and 5033 (both are Seyfert 1.9) are extremely weak and are therefore likely to suffer from optical extinction. Thus, $M_{\text{BH}}$ estimates using the BLR are not possible for these two AGNs. Hence, only one object (NGC 4051; Wandel 1999) in our sample has archival reverberation mapping data giving $M_{\text{BH}} \approx 1.4_{-0.8}^{+5} \times 10^6 M_{\odot}$.

Complimentary to reverberation mapping, and also assuming Keplerian motion, mapping of water maser spots and observations of gas kinematics within the gravitational sphere of influence of the SMBH are thought to yield relatively precise measurements of $M_{\text{BH}}$. Again, due to the nature of the observations and the requirement of a suitable gas disc (i.e. a relatively face-on inclination to the observer), few $M_{\text{BH}}$ estimations using these methods exist in the current literature. Indeed, in our sample there are currently only two AGNs (NGC 1068 and 4945) with $M_{\text{BH}}$ estimations from the mapping of water maser spots ($M_{\text{BH}} \approx 1.6 \times 10^7$ and $1.1 \times 10^6 M_{\odot}$, respectively) and one AGN (NGC 5128) with a spatially resolved gas dynamical $M_{\text{BH}}$ estimate ($M_{\text{BH}} \approx 2.4 \times 10^5 M_{\odot}$).

### 3.1.2 Indirect SMBH mass constraints from the $M-\sigma_*$ relation

Since the seminal discovery that the mass of the stellar spheroid is closely related to $M_{\text{BH}}$ (Magorrian et al. 1998), indirect $M_{\text{BH}}$ estimation methods have become ubiquitous in the current literature (e.g. Ferrarese & Merritt 2000; Gebhardt et al. 2000; Nelson et al. 2004; Onken et al. 2004; Greene & Ho 2006). Comparisons of spatially resolved stellar dynamics with reverberation-based $M_{\text{BH}}$
Specifically, using the Echellette spectrograph on Keck-II, E121–G006, N2010 The Authors. Journal compilation 2010 RAS, MNRAS measurements for NGC 1448 (Scd) is estimated from the relation. Specifically, using the Echellette spectrograph on Keck-II, Barth et al. (2009) measure an accurate line-of-sight stellar velocity dispersion of the Ca II triplet lines (CaII, λ8498, 8542, 8662) observed in the central nuclear star cluster. They find a velocity dispersion consistent with a SMBH mass of $M_{\text{BH}} \approx 3 \times 10^6 M_\odot$. We find that all of the AGNs in our sample host SMBHs with $M_{\text{BH}} \approx 10^8 M_\odot$ and thus are unlikely to suffer significant systematic uncertainties arising from the use of the $M-\sigma_*$ relation even in the most late-type galaxies. 10 of the 13 AGNs in our sample without direct $M_{\text{BH}}$ measurements have published $M_{\text{BH}}$ estimates using the $M-\sigma_*$ relation. For consistency purposes (and where possible) we have used the central stellar velocity dispersions given in the recently published catalogue of Ho et al. (2009). They measure the central $\sigma_*$ for the 486 galaxies within the Palomar survey (Ho, Filippenko & Sargent 1997a,b) using the averaged values derived from the measurement. A. D. Goulding et al. 600–611

Table 1. Catalogue of $D < 15$ Mpc mid-IR identified AGNs and derived quantities.

<table>
<thead>
<tr>
<th>Common name</th>
<th>$D$ (Mpc)</th>
<th>Hubble type</th>
<th>AGN ID</th>
<th>$\log(L_{\text{bol}})$ (erg s$^{-1}$)</th>
<th>$\log(L_X)$ (erg s$^{-1}$)</th>
<th>Ref.</th>
<th>$K_{\text{bol}}$ (mag)</th>
<th>$K_{\text{bul}}$ (mag)</th>
<th>$\log(M_{\text{BH}})$ (M$_\odot$)</th>
<th>Method</th>
<th>Ref.</th>
<th>$M_{\text{BH}}$ (M$_\odot$)</th>
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<tr>
<td>E121–G006</td>
<td>14.5</td>
<td>Sc</td>
<td>IR</td>
<td>39.04</td>
<td>41.82</td>
<td>8.98</td>
<td>10.91</td>
<td>6.10</td>
<td>$14$</td>
<td></td>
<td></td>
<td>L$_{\text{bol}}$</td>
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<td>IR</td>
<td>39.38</td>
<td>42.26</td>
<td>7.03</td>
<td>7.34 $^{+0.08}_{-0.15}$</td>
<td>M–$\sigma_*$</td>
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<td>Sa</td>
<td>IR</td>
<td>39.71</td>
<td>42.69</td>
<td>7.34</td>
<td>7.34 $^{+0.08}_{-0.16}$</td>
<td>M–$\sigma_*$</td>
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<td>IR,O,X</td>
<td>41.66</td>
<td>45.26</td>
<td>5.79</td>
<td>7.20 $^{+0.12}_{-0.12}$</td>
<td>M</td>
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<td>IR</td>
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<td>42.28</td>
<td>7.66</td>
<td>10.64</td>
<td>5.99 $^{+0.11}_{-0.12}$</td>
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<tr>
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<td>IR</td>
<td>38.26</td>
<td>40.49</td>
<td>7.01</td>
<td>9.08</td>
<td>6.83 $^{+0.12}_{-0.12}$</td>
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<td>M–$\sigma_*$</td>
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<td>S0</td>
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<td>39.38</td>
<td>42.26</td>
<td>3.94</td>
<td>8.36 $^{+0.20}_{-0.26}$</td>
<td>G</td>
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<td>10.2</td>
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<td>IR</td>
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<td>40.30</td>
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<td>7.31 $^{+0.07}_{-0.13}$</td>
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<td>Sbc</td>
<td>IR,O</td>
<td>39.78</td>
<td>42.79</td>
<td>6.93</td>
<td>6.80 $^{+0.02}_{-0.24}$</td>
<td>M–$\sigma_*$</td>
<td>24</td>
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</table>

Note: (1) Common galaxy name. (2) Distance to source in Mpc from the Revised Bright Galaxy Survey (RBGS; Sanders et al. 2003). (3) Morphological type from RC3 (de Vaucouleurs et al. 1991). (4) Waveband of AGN identification; IR: mid-IR spectroscopy (GA09); O: optical spectroscopy (references presented in GA09); X: X-ray spectroscopy (2–10 keV; see column 7 for references); R: radio observations. (5) Logarithm of $[O IV]$ λ25.89 µm luminosity in erg s$^{-1}$ calculated using [O IV] flux presented in GA09; mean uncertainty is approximately 10 per cent. (6) Logarithm of absorption corrected hard X-ray luminosity (2–10 keV) in erg s$^{-1}$ which have been converted to the distances given in column 2. (7) Reference for X-ray data. (8) Logarithm of bolometric luminosity of the AGN estimated from $L_{\text{bol}}$ using equation (4). (9) Logarithm of bolometric luminosity of the AGN estimated from $L_{\text{bol}}$ using the bolometric corrections described in Marconi et al. (2004). (10) Total K-band magnitude from 2MASS Large Galaxy Atlas (Jarrett et al. 2003). (11) K-band magnitude of bulge produced using GALFIT (Peng et al. 2002; see Section 3.2.1). (12) Logarithm of estimated black hole mass and associated 1σ errors in solar masses. (13) Method of $M_{\text{BH}}$ measurement; M: maser mapping; G: gas kinematics; R: reverberation mapping; M–$\sigma_*$: mass–velocity dispersion correlation; $L_{\text{bol}}$: K-band luminosity–bulge correlation. (14) Reference for $M_{\text{BH}}$ measurement.

References: (1) Matt et al. (1997); (2) Dadina (2007); (3) Pounds et al. (2004); (4) Tueller et al. (2008); (5) Winter et al. (2009); (6) Guainazzi et al. (2004); (7) Itoh et al. (2008); (8) Cappi et al. (2006); (9) Bird et al. (2007); (10) Fukazawa et al. (2001); (11) Maiolino et al. (1998); (12) Guainazzi et al. (2004); (13) Matsumoto et al. (2004); (14) this paper; (15) HyperLeda; (16) Greenhill et al. (1996); (17) Greenhill et al. (1997); (18) Marconi et al. (2001); (19) Wandel (1999); (20) Ho et al. (2009); (21) Barth, Ho & Sargent (2002); (22) Barth et al. (2009); (23) Whittle (1992); (24) Garcia-Rissmann et al. (2005).

For a review see Israel (1998).
fitting of stellar absorption templates (Valdes et al. 2004) to the blue (4230–5110 Å) and red (6210–6860 Å) spectral ranges (i.e. the published spectroscopy does not include standard velocity dispersion measurement features; e.g. CaT). Where available, Ho et al. (2009) compare \( \sigma_* \) values derived from the modelling of the stellar absorption features to previously published measurements from CaT lines which are available in the HyperLeda data base. Measured errors are compared between the Palomar \( \sigma_* \) measurements and the weighted average adopted by HyperLeda for the available published \( \sigma_* \) measurements. The final adopted measurement of \( \sigma_* \) by Ho et al. (2009) is that with the smallest overall error. Where the values assumed by Ho et al. (2009) are previously published or are from HyperLeda we quote these references in column 14 of Table 1 (six objects). For the AGNs in our sample which are not part of the Palomar survey, values of \( \sigma_* \) derived from direct fitting analyses of the CaT lines are adopted from other published sources (see column 14 of Table 1; four objects). All final adopted \( \sigma_* \) measurements are converted to \( M_{\text{BH}} \) estimates using equation (1).

3.2 Galaxy decompositions using GALFIT

3.2.1 Indirect SMBH mass constraints from the \( M_{\text{BH}}–L_K \) relation

Three of the AGNs within our sample (ESO121−G006, NGC 1448 and 1792) currently lack archival direct or indirect SMBH mass constraints. Hence, for these three objects, we follow the formalism of Marconi & Hunt (2003, hereafter MH03) and use 2MASS \( K \)-band imaging and GALFIT, the 2D imaging analysis software of Peng et al. (2002), to constrain the bulge luminosities and therefore, \( M_{\text{BH}} \) for these three AGNs.

Near-IR (0.9–4.8 \( \mu \)m) emission is a strong tracer of stellar mass and is less susceptible to the effects of dust/gas extinction than optical emission. As a result of this, the \( K \) band (2.2 \( \mu \)m) is shown to provide the strongest correlation of all near-IR bands between the luminosity of the bulge and \( M_{\text{BH}} \) (MH03).

For the bulge–disc decomposition image analysis, we have obtained archival \( K \)-band imaging for ESO121−G006, NGC 1448 and 1792. These images were retrieved from the 2MASS extended source catalogue and consist of pre-mosaicked (1 arcsec pixel\(^{-1}\) resolution) all-sky atlas images. The \( K \)-band images of the three galaxies were modelled with a central point spread function (PSF) and a constant sky background contribution, whilst the bulge and host galaxy components were modelled using variations of the Sérsic profile:

\[
\Sigma(r) = \Sigma_0 e^{-\left(\frac{r}{r_e}\right)^{1/n}} \left(\frac{r}{r_e}\right)^{r_e}, \tag{2}
\]

where \( r_e \) is the effective radius of the profile, \( \Sigma_0 \) is the surface brightness at the effective radius, \( n \) is the power-law (Sérsic) index, and \( \kappa \) is coupled to \( n \) such that half of the total flux of the object is within the effective radius. We employ two special forms of the Sérsic profile in our GALFIT modelling, the exponential \((n = 1)\) and the de Vaucouleurs \((n = 4)\) profiles, which are classically used to model galactic discs and bulges, respectively.

Häußler et al. (2007) have shown that the reliability of the fitting parameters produced by GALFIT are strongly dependent on its initial setup (e.g. background estimates). By extension, we find that GALFIT will, in general, fail to find the overall global chi-squared (\( \chi^2 \)) minimum to the fit if the initial parameters estimates are poorly constrained. Thus, to reduce these systematic effects, and aid the fitting routine, we use a simplified 1D fit to produce initial estimates of the fitting parameters. A 1D surface brightness slice of the \( K \)-band image was taken across the major axis of each of the galaxies. A surface brightness profile extending from the nucleus was produced by averaging the two semimajor axes from the slice, and removing the measured background flux. Few spiral galaxies are found to host bulges with true de Vaucouleurs profiles (e.g. Graham & Worley, 2008), thus a global \( \chi^2 \) reduction process was used to simultaneously fit a generalized Sérsic profile and a fixed \((n = 1)\) exponential disc to the 1D surface brightness profile. From these, we calculate Sérsic and disc radii, as well as the Sérsic index of the bulge. Combining the 1D parameter estimates with the total \( K \)-band magnitude from the 2MASS Large Galaxy Atlas (Jarrett et al. 2003), we generate an appropriate set of constraints and initial parameters to be input to GALFIT.

Using the derived parameter estimates, GALFIT is used to fit a generalized 2D Sérsic profile with an exponential disc to the \( K \)-band image. To aid the GALFIT reduction analysis, particular attention is paid to simulating accurate PSFs for the 2MASS images using known standard stars (J. R. Lucey, private communication). Results of this bulge/disc reduction for the three objects are presented in Fig. 1 and column 11 of Table 1.

We have directly tested our robust GALFIT method using the late-type galaxies (i.e. S0 or later) presented in the data set of MH03 and find close agreement (\( \approx 0.1 \) dex). We do note however that we find a systematic offset of a factor of \( \approx 2 \) in bulge luminosity for the AGNs in MH03 that are hosted in low-inclination angle late-type galaxies, which is likely to be caused by GALFIT overestimating the contribution of the bulge to the total flux of the galaxy. Indeed, when directly comparing a sample of reverberation mapped X-ray-detected AGNs to \( M_{\text{BH}} \) estimations using the MH03 formalism, Vasudevan et al. (2009) find similar results. However, the three galaxies fitted in our sample are all moderately to highly inclined, and thus this systematic effect will be negligible.

In Fig. 1 we show the three GALFIT produced image cubes obtained following our bulge/disc fitting routines. Within each of the residual (observed – model) images it is clear that the bulge is well fitted by a Sérsic profile. The edge-on galaxy ESO121−G006 is well fitted by an exponential disc combined with a Sérsic profile, with no distinguishing residual features. The residual of NGC 1448 highlights the existence of its spiral arms and shows the presence of a truncated disc combined with a possible bar structure which our simplified modelling technique is incapable of fitting; however, the bulge fit does not appear to be compromised. Indeed, our derived \( M_{\text{BH}} \) estimations for ESO121−G006 and NGC 1448 are consistent with the \( M_{\text{BH}} \) upper limits obtained from stellar mass-to-light ratio analyses (Ratnam & Salucci 2000). The residual image of the moderately inclined galaxy, NGC 1792, contains strong spiral arms as well as a point-like nuclear source. We note that due to the

---

\(^{4}\) The HyperLeda data base is a continuously updated electronic catalogue of galactic measurements available at http://leda.univ-lyon1.fr/. Specifically, HyperLeda contains a consolidated list of archival velocity dispersions for many nearby galaxies.

\(^{5}\) We note that NGC 6300 currently has two measurements of \( \sigma_* \) from fitting of the CaT lines (Garcia-Rissmann et al. 2005), which were obtained through direct-fitting and cross-correlation analyses. For consistency with other measurements, we adopt the direct-fitting value of \( \sigma_* \). We find that the derived \( M_{\text{BH}} \) from the cross-correlation method is a factor of 2 larger; however, using this larger \( M_{\text{BH}} \) measurement will have little impact on our overall results.

\(^{6}\) A detailed discussion of GALFIT problems caused by poor PSF modelling can be found in Bentz et al. (2009a).
4.1 Hard X-ray luminosity as a tracer of the bolometric luminosity of an AGN

High-quality hard X-ray spectral analyses arguably provide the most unambiguous method for measuring the intrinsic luminosity of an AGN since (1) X-rays are relatively unaffected by dust extinction; (2) intrinsic absorption can be directly constrained from high signal-to-noise ratio (S/N) data and (3) star formation contamination is often found to be negligible. We therefore divide our sample into two categories based on the quality and energy range of their available published X-ray data: (1) AGNs with high-S/N X-ray spectra where $N_H$ has been accurately constrained and/or AGNs with $E > 10$ keV constraints where the observed X-ray emission will only be strongly absorbed for heavily Compton-thick ($N_H > 10^{25}$ cm$^{-2}$) sources (eight AGNs) and (2) those AGNs with no or low-S/N X-ray data, i.e. where there are insufficient counts to accurately determine $N_H$ and the X-ray flux could have large contributions from star formation (nine AGNs). Hence, we specifically do not estimate $L_{\text{bol,AGN}}$ for those AGNs with $L_X$ measurements using Chandra that also have no further hard X-ray spectral constraints ($E > 10$ keV) due to the limited bandpass of the instrument at $z \sim 0$ (0.5–8 keV).

For the AGNs in our sample currently with either high-S/N spectroscopy or $E > 10$ keV constraints, we estimate $L_{\text{bol,AGN}}$ using the Eddington ratio independent AGN bolometric corrections outlined in equation (21) of Marconi et al. (2004). We note that Vasudevan & Fabian (2009) have suggested that the bolometric correction factor ($\kappa_{2-10\text{keV}} = L_{\text{bol}}/L_{2-10\text{keV}}$) may be a function of the Eddington ratio of a considered source. Values of $\kappa_{2-10\text{keV}} \sim 10$–30 are considered to be relatively low bolometric corrections and are generally found in AGNs with $\eta < 0.1$ (e.g. Vasudevan & Fabian 2009; Vasudevan et al. 2010). For the sample of X-ray-detected AGNs considered here, we calculate similarly consistent values of $\kappa_{2-10\text{keV}} \sim 8$–30, and thus conclude that the Eddington ratio is unlikely to be dominating the bolometric corrections adopted here from Marconi et al. (2004).

For further consistency, all archival X-ray luminosities were adjusted to the distances adopted in column 2 of Table 1. Final adopted $L_X$ measurements, estimated $L_{\text{bol,AGN}}$ from $L_X$, and archival references for $L_X$ are given in columns 6, 7 and 9 of Table 1, respectively.

4.2 [O IV] luminosity as a tracer of the bolometric luminosity of an AGN

For those galaxies in the sample without good-quality hard X-ray constraints we require an additional approach to estimate the intrinsic luminosity of the AGN, and hence the relative mass accretion rate. Here we build upon an $L_{\text{bol,AGN}}$ estimation which relies on the AGN-produced [O IV] $\lambda 25.89\mu$m luminosities ($L_{\text{[O IV]}}$) for the AGNs in our sample (e.g. Dasyra et al. 2008; M08).

Based on the simplest unified model of AGN (Antonucci 1993), the hot dust within the predicted torus, close to the central engine, reprocesses absorbed UV, optical and X-ray emission into mid-IR emission. Hence, AGN emission detected at IR wavelengths is relatively accurate estimates of the bolometric luminosities of the AGNs ($L_{\text{bol,AGN}}$) in our sample. We use two methods to estimate $L_{\text{bol,AGN}}$: (1) a direct approach using the best available measured hard X-ray (2–10 keV) luminosities and (2) a well-constrained $L_{\text{bol,AGN}}$-[O IV] luminosity relation to infer the intrinsic luminosity of the AGN (e.g. Dasyra et al. 2008; Meléndez et al. 2008, hereafter M08).
likely to be isotropic and independent of viewing angle. As discussed in Section 2, the detection of high-ionization [Ne v] emission (97.1 eV) coincident with the nucleus of a galaxy is considered a robust indicator of AGN activity (e.g. Armus et al. 2006; GA09). Complementary to this, GA09 find that [Ne v] emission is also well correlated with [O iv] emission (54.9 eV) with an intrinsic scatter of only 0.24 dex. As [Ne v] emission, and thus [O iv] emission, do not suffer from significant star formation contamination and are both comparatively extinction-free, they may be used as relatively clean proxies for the bolometric luminosity of the AGNs presented here ($L_{\text{bol,AGN}}$). Indeed, for a sample of 35 well-studied optically unobscured AGNs, Dasyra et al. (2008) show that both [Ne v] and [O iv] emission are well correlated with the luminosity of the 5100 Å optical continuum, and hence $L_{\text{bol,AGN}}$ with an intrinsic scatter of 0.46 and 0.47 dex, respectively. However, the relation of Dasyra et al. (2008) is derived from AGNs with $L_{\text{[O iv]}} > 2 \times 10^{40} \text{erg s}^{-1}$ ($L_{\text{[Ne v]}} > 7 \times 10^{39} \text{erg s}^{-1}$), and hence we test whether it may be reliably extrapolated to the more modest luminosity AGNs considered here (log-average $L_{\text{[O iv]}} \approx 2 \times 10^{39} \text{erg s}^{-1}$).

We combine our robustly adopted $L_{\text{bol,AGN}}$ from 2–10 keV flux measurements with the X-ray catalogue of nearby ($z < 0.08$) Seyfert galaxies in the Swift-BAT survey which have published [O iv] luminosities in M08. The catalogue of sources in M08 contains 2–10 keV luminosities obtained primarily from ASCA data, 14–195 keV luminosities from the Swift-BAT survey and [O iv] luminosities from Spitzer-IRS spectroscopy. Combining the M08 sample with our eight AGNs with high-quality X-ray constraints, the range covered in $L_{\text{[O iv]}}$ is $0.7 \times 7000 \times 10^{39} \text{erg s}^{-1}$. For consistency, we convert the M08 2–10 keV luminosities to $L_{\text{bol,AGN}}$ using the same bolometric corrections adopted in Section 4.1. In Fig. 2 we plot $L_{\text{bol,AGN}}$ versus $L_{\text{[O iv]}}$ for the M08 sample and the eight AGNs in our GA09 sample with good-quality hard X-ray constraints (grey filled squares and blue filled squares, respectively).

As noted by M08, several of the AGNs in their sample have low 2–10 keV luminosities when compared to the 14–195 keV luminosities, most likely due to absorption of the 2–10 keV flux. These are highlighted with open circles in Fig. 2 and are shown in Table 2. Using high-quality X-ray spectral analyses, in the literature, five of the seven AGNs are identified as Compton-thick AGNs (see column 4 of Table 2 of this paper, and Table 4 of M08). For the purpose of our analyses, we use absorption-corrected measurements of $L_{\text{2–10 keV}}$ from the literature, which are in good agreement with the M08 14–195 keV luminosities (see columns 2, 3 and 5 of Table 2; dotted lines in Fig. 2). The other two AGNs (NGC 2992 and 4388) in M08 with inconsistent 2–10 and 14–195 keV fluxes are found to be highly variable (Elvis et al. 2004; Beckmann et al. 2007, respectively).

With the inclusion of the obscuration-corrected 2–10 keV luminosities to infer $L_{\text{bol,AGN}}$, we find a strong correlation between $L_{\text{[O iv]}}$ and $L_{\text{bol,AGN}}$ which is characterized by the equation

$$
\log\left(\frac{L_{\text{bol,AGN}}}{10^{44} \text{erg s}^{-1}}\right) = (0.38 \pm 0.09) \log\left(\frac{L_{\text{[O iv]}}}{10^{44} \text{erg s}^{-1}}\right) + (1.31 \pm 0.09) \log\left(\frac{L_{\text{[O iv]}}}{10^{44} \text{erg s}^{-1}}\right)
$$

(4)

We note that since star formation can also produce [O iv] emission, the [Ne v]–[O iv] relation may be unreliable for sources with exceedingly high star formation rates (e.g. ultraluminous IR galaxies with $L_{\text{IR}} > 10^{12} \text{erg s}^{-1}$; ULIRGs). However, there are no ULIRGs within the sample considered here.

Table 2. Catalogue of revised 2–10 keV luminosities for a subset of M08 sample.

<table>
<thead>
<tr>
<th>Common name</th>
<th>log($L_{\text{2–10 keV,M08}}$) (erg s$^{-1}$)</th>
<th>log($L_{\text{2–10 keV,alt.}}$) (erg s$^{-1}$)</th>
<th>Compton-thick AGN? (5)</th>
<th>Ref.</th>
</tr>
</thead>
<tbody>
<tr>
<td>Circinus</td>
<td>40.58</td>
<td>42.04</td>
<td>✔</td>
<td>1</td>
</tr>
<tr>
<td>Mrk 3</td>
<td>41.95</td>
<td>43.20</td>
<td>✔</td>
<td>2</td>
</tr>
<tr>
<td>NGC 1365</td>
<td>40.99</td>
<td>42.40</td>
<td>✔</td>
<td>3</td>
</tr>
<tr>
<td>NGC 2992</td>
<td>41.69</td>
<td>43.50</td>
<td>✗</td>
<td>4</td>
</tr>
<tr>
<td>NGC 3079</td>
<td>40.02</td>
<td>42.18</td>
<td>✗</td>
<td>5</td>
</tr>
<tr>
<td>NGC 4388</td>
<td>41.91</td>
<td>42.57</td>
<td>✗</td>
<td>6</td>
</tr>
<tr>
<td>NGC 6240</td>
<td>42.23</td>
<td>44.00</td>
<td>✔</td>
<td>7</td>
</tr>
</tbody>
</table>

Note: (1) Common galaxy name. (2) Logarithm of 2–10 keV luminosity adopted by M08. (3) Logarithm of absorption corrected 2–10 keV luminosity adopted from individual studies (2–10 keV luminosities were adjusted using our adopted distances). (4) Is the AGN a Compton-thick source? (5) Reference for adopted $L_{\text{2–10 keV}}$ measurement.

References: (1) Yang et al. (2009); (2) Awaki et al. (2008); (3) Risaliti et al. (2009); (4) Yaqoob et al. (2007); (5) Iyomoto et al. (2001); (6) Shirai et al. (2008); (7) Vignati et al. (1999).
we show that the Dasyra et al. (2008) relation will overestimate $L_{\text{bol, AGN}}$ by typically $\approx 0.5$ dex.

In Fig. 2, we highlight two AGNs from the $D < 15$ Mpc sample (NGC 4945 and 5643) which appear to be significant outliers of the observed correlation. NGC 5643 possibly harbours a variable central source. From the detection of strong Fe Kα emission, Maiolino et al. (1998) suggest from using BeppoSAX data that NGC 5643 is possibly Compton thick ($\tau_{\text{N}} \approx 10^{23} \text{cm}^{-2}$); however, Guainazzi et al. (2004) find using XMM–Newton data that it may be Compton thin with $\tau_{\text{N}} \approx 6 \times 10^{23} \text{cm}^{-2}$. Thus, from current available data, the true intrinsic luminosity of the AGN is highly uncertain. Here we conservatively adopt the Compton-thin $L_X$ value of Guainazzi et al. (2004); however, we note that if we use the value of Maiolino et al. (1998), then NGC 5643 would lie on our derived relationship. NGC 4945, by contrast, has observations using Ginga, ASCA, OSSE, the Rossi X-ray Timing Explorer, BeppoSAX, Suzaku and most recently Swift-BAT (Iwasawa et al. 1993; Tanaka, Inoue & Holt 1994; Done, Madejski & Smith 1996; Madejski et al. 2000; Guainazzi et al. 2000; Itoh et al. 2008; Tueller et al. 2008, respectively) all of which provide excellent and consistent spectral constraints over a wide X-ray band ($1–200$ keV) suggesting that the central source is Compton thick ($\tau_{\text{N}} \approx 4 \times 10^{23} \text{cm}^{-2}$), with $L_X \approx 3 \times 10^{42} \text{erg s}^{-1}$ and an observed intrinsic variability of a factor of $\approx 2$. Here we adopt the luminosity from the most recent observation by Swift, $L_X \approx 3 \times 10^{42} \text{erg s}^{-1}$. On the basis of our mid-IR constraints, it would therefore appear that NGC 4945 is underluminous in [O IV] flux by a factor of $\approx 50$ (see Table 1 and Fig. 2). We suggest this deficit in observed [O IV] flux is unlikely to be due to host galaxy extinction; the required absorption to account for a factor of 50 flux difference is $A_V \approx 240$ mag [$\tau_{\text{N}} \approx 5 \times 10^{23} \text{cm}^{-2}$; assuming typical dust-to-gas ratios; using $A_V/E(B-V) \approx 3.1$]. Another possible explanation is a temporary decoupling of the X-ray-emitting and narrow-line (NL) regions (i.e. the highly luminous state of NGC 4945 may be a somewhat recent event). Given the spatial difference (and hence, the light traveltime) between the two emission regions: $10^4$ and $1–10$ pc, respectively, the photoionization of the NL region, and thus the observed [O IV] emission, may take $\approx 100$ yr to respond to the changes in the X-ray-emitting region. We therefore suggest that the intrinsic scatter in the observed [O IV]--$L_{\text{bol, AGN}}$ relation may be significantly reduced if it was possible to account for variability in the central region of all of the AGNs. Indeed, we find the average dispersion decreases to $\approx 0.3$ dex if we remove NGC 4945 from our analysis.

Using equation (4) and $L_{\text{OIV}}$ from GA09 (column 5 of Table 1), we estimate $L_{\text{bol, AGN}}$ (column 7 of Table 1) for those AGNs in our sample currently without good hard X-ray measurements. We use these $L_{\text{bol, AGN}}$ estimates to assess the relative mass accretion rates ($\approx L_{\text{bol, AGN}}/L_{\text{edd}}$) of the SMBHs in our $D < 15$ Mpc sample.

### 5 RESULTS AND DISCUSSION

From a volume-limited sample of 64 bolometrically luminous galaxies to $D < 15$ Mpc ($\approx 94$ per cent complete) GA09 unambiguously identified 17 ($\approx 27\%$ per cent) sources to be hosting AGN activity in galaxies with $L_B \geq 3 \times 10^{10} L_\odot$, using [Ne v] $\lambda 14.32 \mu \text{m}$ emission as a robust AGN indicator. Using the SMBH mass and AGN bolometric luminosity estimates derived here, we discuss the relative mass accretion rates of these 17 AGNs and use them to derive the average present-day growth times of SMBHs in the very nearby Universe. Furthermore, we evaluate the unique contribution that our new optically unidentified AGNs make to the space density of active SMBHs in the local Universe, which have until now been previously derived from large-scale optical surveys (e.g. H04; Greene & Ho 2007).

#### 5.1 Derived AGN properties and relative mass accretion rates

In Fig. 3, we plot $L_{\text{bol, AGN}}$ against our adopted $M_{\text{BH}}$ estimates (the associated 1σ errors for $M_{\text{BH}}$ measurements are described in Section 3) for the 17 AGNs in our volume-limited sample. $L_{\text{bol, AGN}}$ is inferred from either accurate intrinsic high-quality hard X-ray ($2–10$ keV) constraints (where available) or AGN-produced [O IV] $\lambda 25.89 \mu \text{m}$ emission (see Sections 4.1 and 4.2). The $L_{\text{bol, AGN}}$ 1σ errors for the sources with hard X-ray constraints are the result of combining the uncertainty in the $L_X$ measurement with that of the mean spread in the bolometric correction factor employed from Marconi et al. (2004). For those AGNs with $L_{\text{bol, AGN}}$ derived from [O IV], the error is derived from the uncertainty in $L_{\text{bol, AGN}}$ as quoted in GA09 combined in quadrature with the intrinsic scatter of the empirical [O IV]--$L_{\text{bol, AGN}}$ relation (eq. 4).

Seven objects within the sample have been classified as optical AGNs from previous surveys (see GA09 and Table 1) using optical emission-line diagnostics (e.g. the Baldwin–Phillips–Terlevich diagnostic diagrams; Baldwin, Phillips & Terlevich 1981); however, all have detected [Ne v] $\lambda 14.32 \mu \text{m}$ emission, and thus are unambiguously identified to host AGNs at mid-IR wavelengths (GA09). We find that with the exception of NGC 5128 (Centaurus A), our sample is dominated by AGNs with SMBHs in the mass range $M_{\text{BH}} \approx (0.1–5) \times 10^8 M_\odot$ (median of $M_{\text{BH}} \approx 7 \times 10^7 M_\odot$). Due to the irregular structure of one of the galaxies in the sample (NGC 5195), $M_{\text{BH}}$ is poorly determined; in Fig. 3 we plot $M_{\text{BH}}$ estimates from both the $M-\sigma$, and $L_{\text{BH}}$--$L_{K,BH}$ relations (connected blue dashed line).

We find the AGNs in our sample are spread over a wide range of bolometric luminosities, $L_{\text{bol, AGN}} \approx 10^{43}–10^{46} \text{erg s}^{-1}$. To assess the relative mass accretion rates of the sample ($L_{\text{bol, AGN}}/L_{\text{edd}} \sim \eta$), we overplot lines of constant Eddington ratios ($\eta \approx 10^{-3}$, $10^{-1}$, $1$; derived following Rees 1984) and their associated mass-doubling times ($t \approx 30$, $0.3$, $0.03$ Gyr, respectively). Given the large range in bolometric luminosities, it is not surprising that the AGNs in the sample are found to be accreting at rates covering over five orders of magnitude ($\eta \approx 10^{-5}$ to $1$). With the exception of a few AGNs, the observed range in Eddington ratios is found to be roughly consistent with those found by H04 for active galaxies (solid contours in Fig. 3).

As our work is not limited by the spectral resolution of the SDSS (i.e. with a limit of $M_{\text{BH}} \geq 3 \times 10^8 M_\odot$, we show in Fig. 3 that significant accretion, $\eta > 10^{-3}$ (i.e. radiatively efficient accretion systems; e.g. thin discs) occurs on to SMBHs with $M_{\text{BH}} \approx (1–3) \times 10^9 M_\odot$. The majority of these low-mass, rapidly accreting SMBHs are hosted in late-type, disc-dominated spiral galaxies (Sc–Sd). By contrast, it is generally assumed that gas-rich late-type spirals are preferentially inactive galaxies and that a large bulge may be a necessary component for the existence of a SMBH, and thus a luminous AGN. Furthermore, of the four AGNs within the sample

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8 We note that NGC 3627 is ambiguously classified as T2/S2 from the optical spectroscopy presented in Ho et al. (1997a), and thus may host an active central source. However, Roberts, Schurch & Warwick (2001) suggest from its optical emission-line ratios that NGC 3627 is most likely a LINER/H II composite.
The incidence of growing black holes

Figure 3. AGN bolometric luminosity ($L_{\text{Bol,AGN}}$, in erg s$^{-1}$) is plotted against SMBH mass ($M_{\text{BH}}$) for the $D < 15$ Mpc mid-IR identified AGNs presented in GA09. Associated 1σ error bars for $M_{\text{BH}}$ and $L_{\text{Bol,AGN}}$ estimations are shown (see Sections 3 and 5.1, respectively, for details of their derivations). AGNs which are previously identified in optical surveys are highlighted with open circles. AGNs with $M_{\text{BH}}$ estimates from reverberation mapping (downward triangles), gas dynamics (hour glass), maser mapping (diamond), the $M-\sigma_*$ relation (squares) and the $M_{\text{BH}}-L_{K,\text{bul}}$ relation (upward triangles) are plotted. NGC 5195 is represented with both an upward triangle and square (with a dashed-line connector) as both of the $M_{\text{BH}}$ estimates for this galaxy are highly uncertain given its irregular morphology. Constant ratios of Eddington luminosity and their implied SMBH mass-doubling times are illustrated for $\eta = 10^{-3}, 10^{-1}, 1$ (30, 300 and 30 Myr; solid, short-dashed and long-dashed lines, respectively). Contours are shown for the active galaxies in the SDSS optical survey of H04. In general, we probe lower SMBH masses and AGN luminosities than those of H04, and we find that the majority of these AGNs would not be detected using optical SDSS data alone.

with SMBHs consistent with $M_{\text{BH}} \approx 10^6 M_{\odot}$, we find that three sources are not identified as AGNs in sensitive optical surveys. This indicates that significant SMBH accretion may be missed by statistically large optical surveys such as H04 even if the spectral resolution was sufficient to identify SMBHs down to $M_{\text{BH}} \approx 10^6 M_{\odot}$.

For the subset of our AGN sample which host SMBHs with $M_{\text{BH}} \gtrsim 3 \times 10^6 M_{\odot}$, we find that many of the optically unidentified AGNs are accreting at relatively low Eddington ratios ($\eta \gtrsim 10^{-3}$), and are unlikely to make a significant additional contribution to the present-day growth of SMBHs. However, these same AGNs may form part of a separate, underlying population of radiatively inefficient accretion systems such as advection dominated accretion flows (e.g. Narayan & Yi 1994) or those which contain optically thick slim discs. Further spectral analysis of the X-ray data may distinguish between these particular accretion systems, but is beyond the scope of these analyses (see Goulding et al., in preparation).

5.2 The present-day growth of SMBHs

Using the relative mass accretion rates estimated for our sample (Fig. 3), we can infer the volume-averaged growth time of SMBHs in the local Universe. Assuming a mean Kerr spin parameter ($a$) for our sample of $a \approx 0.67$ (e.g. Treister & Urry 2006; Hopkins, Richards & Hernquist 2007), i.e. an accretion efficiency ($\epsilon$) of $\approx 0.1$, the characteristic mass-doubling time ($t_{2M}$) of a SMBH accreting matter at the Eddington limit is $t_{2M} \approx 30$ Myr (Rees 1984). Under the further assumption that $a$, and hence $\epsilon$, does not vary significantly for changes in $M_{\text{BH}}$ (King, Pringle & Hofmann 2008), we assess the present-day growth rate of SMBHs.$^9$

$^9$ We note that the spin variation and spin directionality of SMBHs in AGNs is currently an ongoing area of research, and a consensus between groups has yet to be reached for an average value of the Kerr spin parameter; for example see Brenneman & Reynolds (2006), King et al. (2008) and Fabian et al. (2009).
and extend to lower masses ($M_{\text{BH}} < 3 \times 10^6 M_\odot$) the integrated growth of SMBHs. Growth time errors are calculated from the lognormal standard deviations of the sample. We note here that we also include the optically unidentified AGNs which would not be detected in the SDSS.

In Fig. 4, we find that the mean growth time for low-mass SMBHs ($M_{\text{BH}} \approx 10^6 M_\odot$) is $\approx 6.6^{+1.3}_{-1.1}$ Gyr, which is consistent with these AGNs growing on time-scales similar to that of the age of the Universe. Our results are found to be broadly consistent with a simple extrapolation of the growth times calculated by H04 to $M_{\text{BH}} \approx 10^7 M_\odot$ (dashed line in Fig. 4). Thus, the AGNs hosting SMBHs in the mass range $M_{\text{BH}} \approx 10^6$–$10^7 M_\odot$, which are dominated by optically unidentified AGNs (see Fig. 3), are acquiring a significant proportion of their mass in the present day, and are amongst the most rapidly growing SMBHs in the local Universe. Furthermore, we find our derived growth times of SMBHs with $M_{\text{BH}} \approx 3 \times 10^6 M_\odot$ are in good agreement with those presented in H04, with mean growth times of $t_{2M} \approx 47^{+23}_{-18}$ and $\approx 198^{+190}_{-196}$ Gyr for AGNs with $M_{\text{BH}} \approx (0.3–3)$ and (3–30) $\times 10^5 M_\odot$, respectively.\footnote{We note that our sample contains only one galaxy with $M_{\text{BH}} \approx 10^5 M_\odot$ (NGC 5128) and thus may not be representative for high $M_{\text{BH}}$ systems.}

5.3 Space density of AGNs in the local Universe

An accurate active SMBH mass function, especially for lower mass SMBHs ($M_{\text{BH}} \approx 10^6 M_\odot$), is crucial for extending our understanding of the role played by accretion in the growth of SMBHs across cosmic time. In this section we calculate the space density of active SMBHs for our sample and compare it to complimentary optical studies of local NL (H04) and BL (Greene & Ho 2007) AGNs, and the total mass function of local SMBHs by Marconi et al. (2004).

Following Greene & Ho (2007), in the top panel of Fig. 5 we plot the volume-weighted space density, $\Phi$ against $M_{\text{BH}}$ in mass bins of 0.5 dex. The volume, $V$, encompassed by the GA09 sample to $D < 15$ Mpc is $V \approx 1.3 \times 10^6$ $\text{Mpc}^3$\footnote{The RBGS includes all IR-bright galaxies detected by IRAS with $f_{60\mu\text{m}} > 5.24$ Jy at $|b| > 5^\circ$.}. As we note in Section 2, given the luminosity limit imposed in our volume-limited survey ($L_{\text{IR}} \approx 3 \times 10^8 L_\odot$), we do not include the dwarf like systems which are likely to host the very smallest SMBHs ($M_{\text{BH}} < 10^5 M_\odot$).

With an adjustment for the distance model adopted in this paper, an examination of the Palomar survey (Ho et al. 1997a) shows there are possibly three optical Seyferts (NGC 185, 1058 and 4395) with $M_{\text{BH}} < 10^5 M_\odot$ which are not included in our sample due to our lower luminosity limit. However, we also note that NGC 185 has since been reclassified as an H II galaxy (Ho & Ulvestad 2001).

The derived volume-weighted space density for our active SMBHs (filled squares), which is dominated by NL AGNs, is found to be significantly greater (a factor of $\approx 100$) than the SMBH density of BL AGNs (filled circles) presented in Greene & Ho (2007) in the mass region $M_{\text{BH}} \approx (0.9–90) \times 10^5 M_\odot$. The significant increase in active SMBH density when compared to the BL AGN density of Greene & Ho (2007) is partially to be expected due to the greater relative sensitivity of our Spitzer-IRS observations coupled with the greater abundance of observed Seyfert 2 to Seyfert 1 galaxies identified in the local Universe. However, this still may not be a good indicator of the intrinsic Seyfert 1:Seyfert 2 ratio. Tommasin et al. (2010) find in a large sample of 81 Seyfert galaxies that only six sources do not contain significant [Ne v] $\lambda 14.32 \mu$m emission in their mid-IR spectroscopy, the majority of which are Seyfert 1s. It is likely that the identification of low equivalent width emission lines (such as [Ne v] or [O iv]) in BL AGNs is further complicated by a strong IR continuum emission which dominates the mid-IR regime. Hence, it is possible that by requiring the detection of [Ne v] $\lambda 14.32 \mu$m to infer AGN status, we may be rejecting BL objects, and thus finding a lower Seyfert 1:Seyfert 2 ratio than is representative in the local Universe.

In comparison to the active SMBH mass function containing the optically identified NL AGNs of H04 (dotted line), we also find a significantly larger space density of SMBHs. We find that the space density of active SMBHs identified in the mid-IR is roughly constant in the mass region $M_{\text{BH}} \approx (0.9–90) \times 10^5 M_\odot$ with a value of $\Phi \approx 6.3 \times 10^{-4}$ $\text{Mpc}^{-3}$ $\log M_{\text{BH}}$. This space density of AGNs is a factor of $\approx 10$ greater than that estimated by H04 over the same $M_{\text{BH}}$ range. Since we find only two of the 17 AGNs in our sample are sufficiently luminous/unobscured to be detected in the SDSS survey, we determine that this is consistent with our results. We further suggest that the space density derived here may still be a lower limit for the number of NL AGNs in the local Universe. A further examination of the (distance model adjusted) Palomar survey suggests that at least four further NL AGNs are not included in our volume-limited survey. Of these, two (NGC 3486 and 4565) lack high-resolution Spitzer-IRS spectroscopy of the central region (as noted in table 2 of GA09), one (NGC 3031; $L_{\text{IR}} \approx 2.8 \times 10^8 L_\odot$) lies fractionally below our luminosity limit for this survey, and NGC 4258 is not included in the RBGS due to its extremely large angular size.
The incidence of growing black holes

Figure 5. Upper panel: comparison of volume-weighted space densities of active SMBHs in the local Universe, \( \Phi \) in units of number \( \text{Mpc}^{-3} \log M_{\text{BH}}^{-1} \). Mid-IR active SMBH function (filled squares; GA09) is compared to the optically identified NL AGN function (dotted curve) of H04, the BL AGN function (filled circle) of Greene & Ho (2007), and total SMBH function (active + inactive galaxies; solid curve; Marconi et al. 2004). Sample selection bias is analysed using a robust Monte Carlo simulation (shaded region; see Appendix A). Lower panel: ratio of mid-IR active SMBHs to the total local SMBH mass function. The total SMBH mass function is extrapolated by 0.3 dex to \( M_{\text{BH}} < 10^6 M_\odot \). For comparison the volume-weighted AGN fraction of H04 is also shown (dashed line). We estimate a mean volume-weighted local AGN fraction of \( \approx 25^{+29}_{-14} \) per cent over the range \( M_{\text{BH}} \approx (0.5–500) \times 10^6 M_\odot \).

Whilst the main focus of this paper is to compare SMBH statistics derived from mid-IR and optical detection techniques, it is prudent to note that the majority of AGN space densities calculated at higher redshifts are typically derived from sources detected in wide-field X-ray surveys. Hence, we now establish whether our mid-IR active space density may be missing a significant fraction of X-ray-detected AGNs. Recently, a comparison between high-resolution Spitzer-IRS spectroscopy and X-ray-detected AGNs was made by Dudik, Satyapal & Marcu (2009) for a large sample of optically classified LINERS. Dudik et al. (2009) reported inconsistencies between [Ne v] non-detections and the presence of hard X-ray nuclear emission in a subset of their sample. However, they conclude that the limited sensitivity of their mid-IR observations may be driving their observed result. Using the relations between high-ionization mid-IR emission and hard X-ray luminosity (M08; GA09; Goulding et al., in preparation), we suggest that a detected [Ne v] luminosity of \( L_{\text{[Ne v]}} \approx 10^{48} \text{erg s}^{-1} \) in an AGN (i.e. the limiting luminosity in Dudik et al. 2009) would be equivalent to a hard X-ray luminosity of \( L_{\text{X,2–10keV}} \approx 5 \times 10^{40} \text{erg s}^{-1} \); indeed, almost all of the X-ray-detected AGNs which lack significant [Ne v] emission in Dudik et al. (2009) are below this threshold. We thus conclude that with sufficiently sensitive high-resolution mid-IR spectroscopy, there is currently no conclusive evidence to suggest that X-ray-detected type II AGNs lack significant [Ne v] emission in the mid-IR. Hence, with the exception of some Seyfert 1 galaxies (as discussed previously), it is unlikely that our derived mid-IR space density lacks significant numbers of X-ray-detected AGNs with [Ne v] emission below our sensitivity limit.

Given our comparatively small volume to that considered by using the SDSS (e.g. Greene & Ho 2007; H04), we further validate our derived space density of active SMBHs by robustly testing our results to find if (1) our sample is overdense, and thus strongly subject to cosmic variance or (2) given the modest errors associated with our \( M_{\text{BH}} \) estimations, the derived space density is strongly subject to scattering of objects in our defined binning structure. We discuss these analyses in Appendix A. Briefly, we find that our sample is broadly representative of galaxies to \( z \sim 0.3 \), and further show that even in our most pessimistic case, we find an increase in our derived space density of at least a factor of \( \approx 2 \) (maximum increase by a factor of \( \approx 11 \)) at \( M_{\text{BH}} \approx 3 \times 10^6 M_\odot \) over the optical NL space density of H04.

5.4 The volume-weighted local AGN fraction

The ratio of the space densities of the active SMBH to total SMBH mass function (i.e. the volume-weighted local active SMBH fraction) is shown in the lower panel of Fig. 5. We calculate an overall active SMBH fraction of \( \approx 25\%^{+14}_{-16} \) per cent for SMBHs of \( M_{\text{BH}} \approx (0.5-500) \times 10^6 \text{M}_\odot \) down to our [Ne v] completeness limit (\( L_{\text{Ne v}} \gtrsim 10^{58} \text{erg s}^{-1} \)). We find that this fraction is consistent with being constant throughout this \( M_{\text{BH}} \) range. However, given our detection sensitivity limit, we are unable to probe lower Eddington ratios for AGNs hosting smaller SMBHs. Instead, we consider the effect of the AGN fraction for a fixed value of Eddington ratio (e.g. \( \eta > 10^{-3} \); i.e. thin-disc accretion systems) and find tentative evidence that the AGN fraction (\( \approx 16^{+4}_{-6} \) and \( \approx 8^{+10}_{-3} \) per cent) may increase with decreasing SMBH mass (\( M_{\text{BH}} \approx 10^6-10^7 \text{M}_\odot \) and \( M_{\text{BH}} \approx 10^7-10^{11} \text{M}_\odot \) bins, respectively).

For the lowest mass SMBHs (\( M_{\text{BH}} \approx (5-30) \times 10^6 \text{M}_\odot \)), we estimate an overall non-negligible volume-weighted AGN fraction of \( 18^{+12}_{-7} \) per cent, potentially showing that a considerable proportion of small-bulge (and pseudo-bulge) galaxies (i.e. late-type spiral galaxies; Sc–Sd) host AGN activity. It has been previously suggested by Ho4 and Greene & Ho (2007) that the AGN fraction may peak at \( M_{\text{BH}} \approx (0.7-2) \times 10^6 \text{M}_\odot \). However, with the inclusion of the additional low-mass optically unidentified AGNs (see Fig. 3), we find that the AGN fractions are consistent with remaining constant or even increasing for \( M_{\text{BH}} < 10^7 \text{M}_\odot \).

As noted in Section 5.3, the space density of AGNs, and hence the local AGN fraction derived in this paper, may only be a lower limit given the nature of our volume-limited survey which by definition does not include IR-faint systems. To improve upon these current source statistics, a larger sample of late-type spiral galaxies would be required to investigate our findings further. With the greater sensitivity and resolving power proposed for the next generation of space-based mid-IR spectrographs, for example the Space Infrared Telescope for Cosmology and Astrophysics (SPICA)\(^{12}\) and the mid-IR instrument (MIRI) on board the James Webb Space Telescope (JWST),\(^{13}\) surveys such as these can be continued and extended to study more distant (i.e. greater volumes) and heavily obscured AGNs.

6 CONCLUSIONS

We have presented the mean growth times and volume-weighted space density of active SMBHs in the local Universe. Our sample of 17 AGNs was derived from a sensitive volume-limited mid-IR spectral survey of all IR-bright galaxies to \( D < 15 \text{Mpc} \) carried out using the NASA Spitzer Space Telescope (see GA09 for further details on the sample selection). The most accurate SMBH masses available for the objects are compiled from a variety of sources. For the three AGNs without published \( M_{\text{BH}} \) estimates, we use a bulge/disc decomposition method to determine the bulge luminosity and hence a SMBH mass (see Section 3). Our main findings are the following.

1. Using combined mid-IR emission line and high-quality hard X-ray constraints, we have derived accurate measurements of the intrinsic luminosities of our sample of AGNs (see Section 4). In conjunction with the well-established SMBH measurements from previous studies and our own estimates from our aforementioned bulge/disc decomposition method, we have assessed the relative mass accretion rates of our sample. Due to our high sensitivity and the ability to probe low SMBH masses, we find that significant mass accretion (\( \eta > 10^{-3} \)) occurs on to SMBHs with \( M_{\text{BH}} \approx 10^6 \text{M}_\odot \), the majority of which would not be detected in even the most sensitive optical surveys (see Section 5.1).

2. Using our derived relative mass accretion rates for the sample and assuming a typical accretion efficiency of \( \epsilon \approx 0.1 \), we assessed the characteristic mean mass-doubling times (\( \tau_{\text{MD}} \)) for AGNs in the very nearby Universe. For AGNs hosting SMBHs with \( M_{\text{BH}} \approx (0.5-50) \times 10^6 \text{M}_\odot \) we find consistent growth times (\( \tau_{\text{MD}} \approx 47-198 \text{Gyr} \)) with those of the NL AGNs identified in the SDSS (H04). However, we also find that SMBHs with \( M_{\text{BH}} < 5 \times 10^6 \text{M}_\odot \) (i.e. below the spectral resolution limit of the SDSS) are amongst the most rapidly growing SMBHs in the local Universe, with present-day growth times consistent with (and possibly less than) the current age of the Universe (\( \tau_{\text{MD}} \approx 6\times 10^9 \text{Gyr} \)) (see Section 5.2).

3. To assess the incidence of this population of low-mass, rapidly growing SMBHs, we constructed a local space density function of active SMBHs. We find that active SMBHs may be at least a factor of \( \approx 2 \) more common than previously identified in NL AGN surveys using SDSS data. Furthermore, we estimate a non-negligible space density for low-mass SMBHs (\( M_{\text{BH}} \approx 10^6 \text{M}_\odot \)) of \( \Phi \approx 6 \times 10^{-4} \text{Mpc}^{-3} \log M_{\text{BH}} \), which is consistent with the space density of more massive active SMBHs (\( M_{\text{BH}} \approx 10^7 \text{M}_\odot \); i.e. those previously determined to be the most rapidly accreting population of SMBHs) (see Section 5.3).

4. Using a local total SMBH mass function (Marconi et al. 2004), we estimate a mean volume-weighted local AGN fraction of \( \approx 23^{+12}_{-17} \) per cent, which remains relatively constant in the mass range \( M_{\text{BH}} \approx (1-10) \times 10^6 \text{M}_\odot \). However, when only considering the SMBHs with \( \eta > 10^{-3} \) (i.e. radiatively efficient accretion systems), we find tentative evidence for an increasing AGN fraction (\( \approx 16^{+10}_{-6} \) and \( \approx 8^{+10}_{-3} \) per cent) with decreasing SMBH mass (for \( M_{\text{BH}} \approx 10^6-10^7 \text{M}_\odot \) and \( M_{\text{BH}} \approx 10^7-10^{10} \text{M}_\odot \), respectively) (see Section 5.4).

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\(^{12}\) See http://www.ir.isas.jaxa.jp/SPICA/.

\(^{13}\) See http://www.roe.ac.uk/uktac/consortium/miri/.
APPENDIX A: VALIDATION OF DERIVED SPACE DENSITY OF ACTIVE SMBHS

As the space density of active SMBHs derived in this paper is significantly larger than that found in previous studies, in this section we attempt to validate our results by discussing possible limitations and additional sources of error which may exist in these analyses: (1) whilst the sensitivity of the data used in this survey is high, the volume considered is relatively small compared to that of the SDSS, thus our results may be subject to cosmic variance and (2) given the modest errors associated with the \( M_{\text{BH}} \) estimates, the adopted \( M_{\text{BH}} \) binning structure is likely to be subjective and thus degenerate towards objects scattering between the defined bins. Under these assumptions, in Section A1 we investigate whether the sample is indeed representative of the local Universe and in Section A2 we discuss the construction of a Monte Carlo simulation to assess the effect of our adopted \( M_{\text{BH}} \) binning.

A1 Is our sample representative of the local Universe?

Given the large incidence of AGNs within our sample it is possible that the volume considered in our sample is overdense compared to other regions in the local Universe. The construction of the original sample of 64 bolometrically luminous galaxies in GA09 was designed to be complete down to the flux limit of the Revised Bright Galaxy Sample. This proposed a distance constraint of \( D < 15 \) Mpc (see fig. 1 of GA09), and hence did not include the Virgo cluster at \( D \approx 16 \) Mpc, and thus our sample does not incorporate known local overdensities; however, this volume may not be representative of the Universe at large.

To robustly test our considered volume (\( V \approx 1.3 \times 10^6 \) Mpc\(^3\)), we constructed a total SMBH space density function for all galaxies to \( D < 15 \) Mpc and compared this to the local total (active + inactive) SMBH mass function of Marconi et al. (2004) derived from the luminosity function of local galaxies. Given the large comoving volume (\( V_c \approx 1000 \) Gpc\(^3\)) considered in Marconi et al. (2004), their derived SMBH mass function is unlikely to suffer from significant cosmic variance.

Our total SMBH mass function was formulated using all galaxies identified in the NED to \( D < 15 \) Mpc with a total \( K \)-band luminosity of \( L_{K, \text{gal}} \gtrsim 1.5 \times 10^9 \) L\(_{\odot}\). The luminosity threshold is designed to include all galaxies which could potentially host a SMBH with \( M_{\text{BH}} \gtrsim 10^6 \) M\(_{\odot}\) using the \( M_{\text{BH}} = L_{K, \text{gal}} \) relation. This conservative lower limit assumes that \( L_{K, \text{gal}} \approx L_{K, \text{bul}} \) (i.e. that all galaxies in the sample have an early-type galaxy classification). In reality, the majority of the sources identified in NED to \( D < 15 \) Mpc are late-type galaxies (i.e. \( L_{K, \text{gal}} \gg L_{K, \text{bul}} \)), and therefore the \( M_{\text{BH}} \) limit is likely to be \( M_{\text{BH}} \ll 10^6 \) M\(_{\odot}\). To \( D < 15 \) Mpc, we identify 105 galaxies which potentially host a SMBH with \( M_{\text{BH}} \gtrsim 10^6 \) M\(_{\odot}\). To estimate \( M_{\text{BH}} \) in each of these galaxies, we relate the associated Hubble-type from the Third Reference Catalogue of Bright Galaxies (de Vaucouleurs et al. 1991) to a mean bulge/disc ratio (e.g. Benson et al. 2007) and establish an individual bulge luminosity based on galaxy-type and \( L_{K, \text{bul}} \). We convert the estimated bulge luminosity to \( M_{\text{BH}} \) using the \( M_{\text{BH}} = L_{K, \text{bul}} \) relation and construct a total SMBH mass function. Whilst this rather crude estimation carries large associated errors, we still find very good agreement (a mean variance of 0.1 dex) with the SMBH mass function of Marconi et al. (2004) throughout the mass range \( M_{\text{BH}} \approx 10^6 - 10^8 \) M\(_{\odot}\). Thus, to first order, our volume-limited sample does not appear to be overdense and/or subject to strong cosmic variance, and hence, is broadly representative of the typical field-galaxy population in the local Universe.

We further validate this conclusion by estimating the active SMBH mass function for those AGNs which are sufficiently optically bright to be included in H04. Under this assumption, the derived space densities of active SMBHs for this sample and that of H04 should be comparable if our sample is indeed representative of a field-galaxy population. We find that using the H04 detection limit inferred from Fig. 3, our new estimated space density is consistent with H04 (\( \Phi \approx 10^{-4} \)) throughout the mass range considered here \( [M_{\text{BH}} \approx (0.05-30) \times 10^6 \text{M}_{\odot}] \), although there are still considerable uncertainties given the small-number statistics inherent with our sample.

A2 Monte Carlo analysis of \( M_{\text{BH}} \) binning

We have established that our sample and the estimate of the SMBH mass density (both from active and inactive galaxies) do not appear to be subject to overdensities caused by cosmic variance. We therefore now investigate the effect of small-number statistics on our results which are inherent in relatively small samples such as this.

We show in Section 3 that accurate estimates of \( M_{\text{BH}} \) for our sample of AGNs from a homogeneous method is difficult to achieve. Thus, our adopted values of \( M_{\text{BH}} \) are derived from a variety of methodologies. In Fig. 5, the AGNs in the sample are placed into equal bins of \( M_{\text{BH}} \) with width \( 0.5 \) dex. However, this binning process does not allow for the error inherent to each individual \( M_{\text{BH}} \) measurement. Hence, some objects may scatter out of one defined bin and into another. The bin width we employ in our relatively modest sample (64 objects) may therefore be subjective and requires testing. Here we use a Monte Carlo analysis to assess the effect of the scattering of AGNs into different \( M_{\text{BH}} \) bins on the derived space density of active SMBHs in the local Universe (\( \Phi \)).

Our Monte Carlo analysis calculates \( \Phi \) by selecting a random set of SMBH masses from Gaussian probability distributions constructed using our adopted \( M_{\text{BH}} \) masses and their associated 1σ errors given in column 12 of Table 1. The binning structure was designed to incorporate at least one object in each bin from 250,000 realizations of the simulation across the considered mass range of the sample \( [M_{\text{BH}} \approx (0.5-500) \times 10^6 \text{M}_{\odot}] \). We impose the upper mass limit as we have only one high-mass AGN (NGC 5128; \( M_{\text{BH}} \approx 2.4 \times 10^9 \text{M}_{\odot} \)) in the sample. Thus, we determine that given our sample distribution the maximum number of equal-width bins to be nine (i.e. equal bin sizes of 0.33 dex with \( \gtrsim 1 \) object). Hence, we use our simulation to conservatively assess the maximum error on our calculated space density of active SMBHs.

The error shown in Fig. 5 (shaded region) is the standard deviation of the 250,000 simulations in our volume-limited sample combined in quadrature with the Poisson error determined from the number counts in our real sample. We found 250,000 realizations to be sufficient, since at this level the maximum variation in the standard deviation over multiple runs was less than \( 10^{-6} \).
We find that the mean spread in the derived value of $\Phi$ from our volume-limited sample is $\approx 0.85$ dex, and as predicted appears to be subject to some scattering of $M_{\text{BH}}$. However, using our conservative error analyses, we show that the space density of active SMBHs found from our sample is consistently greater than that found for optical NL AGNs (H04) with $M_{\text{BH}} < 10^8 M_\odot$. Thus, to first order, we find a significant increase (of a minimum factor of $\approx 2$ and a maximum of $\approx 80$) in the space density of active SMBHs in the local Universe in the mass range $M_{\text{BH}} \approx (0.5-100) \times 10^6 M_\odot$. For $M_{\text{BH}} > 10^8 M_\odot$ our errors increase significantly due to very limited source statistics (see Section 4.1). We thus conclude that while this survey is subject to the scattering of objects through the $M_{\text{BH}}$ bins, even in our most pessimistic case, we still find a significant increase in the space density of active SMBHs in the local Universe compared to that found in large-scale optical NL AGN studies.

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