Black hole clustering in cosmological hydrodynamic simulations: evidence for mergers

Colin Degraf,1,* Tiziana Di Matteo1 and Volker Springel2

1McWilliams Center for Cosmology, Carnegie Mellon University, 5000 Forbes Avenue, Pittsburgh, PA 15213, USA
2Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Straße 1, 85740 Garching bei München, Germany

Accepted 2010 December 16. Received 2010 December 7; in original form 2010 September 10

ABSTRACT
We examine the clustering properties of a population of quasars drawn from fully hydrodynamic cosmological simulations that directly follow black hole (BH) growth. We find that the BH correlation function is best described by two distinct components: contributions from BH pairs occupying the same dark matter halo ('1-halo term', $\xi_{BH,1h}$) which dominate at scales below $\sim 300$ kpc $h^{-1}$, and contributions from BHs occupying separate haloes ('2-halo term', $\xi_{BH,2h}$) which dominate at larger scales. From the 2-halo BH term, we find a typical host halo mass for faint-end quasars (those probed in our simulation volumes) ranging from $M \sim 10^{11}$ to a few $10^{12}$ M$_\odot$ from $z = 5$ to 1 respectively. The BH correlation function shows a luminosity dependence as a function of redshift, though weak enough to be consistent with observational constraints. At small scales, the high resolution of our simulations allows us to probe the 1-halo clustering in detail, finding that $\xi_{BH,1h}$ follows an approximate power law, lacking the characteristic decrease in slope at small scales found in 1-halo terms for galaxies and dark matter. We show that this difference is a direct result of a boost in the small-scale quasar bias caused by galaxies hosting multiple quasars (1-subhalo term) following a merger event, typically between a large central subgroup and a smaller, satellite subgroup hosting a relatively small BH. We show that our predicted small-scale excess caused by such mergers is in good agreement with both the slope and amplitude indicated by recent small-scale measurements. Finally, we note the excess to be a strong function of halo mass, such that the observed excess is well matched by the multiple BHs of intermediate mass ($10^7$–$10^8$ M$_\odot$) found in hosts of $M \sim 4$–$8 \times 10^{11}$ M$_\odot$, a range well probed by our simulations.

Key words: black hole physics – methods: numerical – galaxies: active – galaxies: evolution – galaxies: nuclei – quasars: general.

1 INTRODUCTION
With supermassive black holes (BHs) being found at the centre of most galaxies (Kormendy & Richstone 1995), interest in quasars has increased significantly, with substantial investigation into fundamental relations between BH masses and their host galaxies’ properties (Magorrian et al. 1998; Ferrarese & Merritt 2000; Gebhardt et al. 2000; Tremaine et al. 2002; Graham & Driver 2007). In addition to these relations, statistical studies of the spatial clustering of quasars provide the potential to better understand the relation between quasars, their hosts and the underlying dark matter (DM) distribution, as well as estimate quasar lifetimes (see e.g. Haiman & Hui 2001; Martini & Weinberg 2001) across a relatively large range of redshift. For example, strong clustering would suggest quasars should reside in massive groups. If so, they should be rare and in order to reproduce the quasar luminosity density, they must have long lifetimes. Conversely, low correlation would suggest more common quasars, and thus shorter quasar lifetimes.

Early studies of quasar clustering produced varying results for the clustering amplitude, with no clear agreement on overall evolution with redshift, some suggesting minimal or decreasing clustering evolution (Mo & Fang 1993; Croom & Shanks 1996), while others found an increase in clustering with redshift (Kundic 1997; La Franca, Andreani & Cristiani 1998). These findings were generally poorly constrained due to the small sizes of available quasar samples. With the emergence of large-scale surveys such as Sloan Digital Sky Survey (SDSS; York et al. 2000) and the Two-degree Field (2dF) QSO Redshift Survey (Lewis et al. 2002), substantially larger catalogues have been compiled, permitting more detailed investigation into the clustering properties of quasars, and many recent studies have been made into this area (e.g. La Franca, Andreani & Cristiani 1998; Porciani, Magliocchetti & Norberg 2004; Croom et al. 2005; Shen et al. 2007; Myers et al. 2007a; da Angela et al.

*E-mail: cdegraf@andrew.cmu.edu

© 2011 The Authors
In addition to large-scale evolution, the luminosity dependence (if any) of large-scale clustering can provide significant insight into what quasar populations dominate different luminosity ranges. For example, the model of Hopkins et al. (2005a,b,c,d, 2006) suggests that bright and faint quasars are similar objects which are observed at different phases of their lifetimes, rather than being fundamentally different populations of quasars (as simpler, ‘on-off’ models assume). This model would suggest that both bright and faint quasars should populate similar haloes. Thus, while there may be some correlation between peak luminosity and halo mass, clustering dependence on instantaneous luminosity should be relatively weak, particularly when compared to more traditional ‘on-off’ models of quasar luminosity (Lidz et al. 2006). Recent observational studies have generally found a lack of luminosity dependence in the correlation function (see, e.g. Croom et al. 2005; Myers et al. 2007a; da Ángela et al. 2008), though Shen et al. (2009) found evidence for some, though weak, luminosity dependence. Several semi-analytic models have also been used, finding differing luminosity dependences, such as a significant dependence for sufficient luminosity ranges, but limited when considering only luminosities probed by observation (Bonoli et al. 2009), or weak dependence at low redshift ($z < 1$), but stronger at higher redshift (Croton 2009).

In addition to large-scale behaviour, the possibility of excess quasar clustering on very small scales has arisen in several recent studies. While some observed quasar pairs are believed to be the result of gravitationally lensed quasars, it has been proposed that others may be physically distinct quasar binaries, which would suggest quasars cluster much more strongly on small scales than extrapolation of large-scale clustering would imply (Djorgovski 1991; Hewett et al. 1998; Kochanek, Falco & Muñoz 1999; Mortlock, Webster & Francis 1999), suggesting a connection between galaxy mergers and quasar activity (see e.g. Kochanek et al. 1999). However, investigating the smallest scale clustering has typically been problematic due to observational limitations (such as fibre collisions preventing small-separation pairs from being distinguished as distinct objects) and sample sizes insufficient for probing the smallest scales, where quasar pairs are rare. There have been several studies probing clustering at sub-Mpc scales, generally finding no excess clustering relative to an extrapolation of the large-scale clustering behaviour (see e.g. Shen et al. 2010; Padmanabhan et al. 2009). However, these studies have been limited to scales above 100 kpc $h^{-1}$, while several recent studies have managed to probe even smaller scales, where they do indeed find a significant excess (Hennawi et al. 2006; Myers et al. 2007b, 2008).

In particular, Hennawi et al. (2006) studied binary quasars from SDSS and 2dF Quasar Survey to compute quasar clustering for scales as small as 20 kpc $h^{-1}$ (comoving), and found significant excess clustering relative to the power-law extrapolation of larger scale observations (by an order of magnitude at comoving scales below 100 kpc $h^{-1}$, and growing stronger with decreasing scale). This excess implies that the quasars are more strongly clustered than galaxies at these small scales, supporting the theory that quasar activity is triggered by galaxy interactions. Using the quasar sample from Myers et al. (2007a,b) found only a slight excess in small-scale clustering, and put an upper limit for the excess at a factor of $4.3 \pm 1.3$ for physical scales of $\sim 28$ kpc $h^{-1}$. They suggest that the significantly larger excess of Hennawi et al. (2006) is a result of a selection effect, possibly due to studies tending to target tracers of the Ly$\alpha$ forest, causing a bias towards $z > 2$, which may be more highly clustered. Myers et al. (2008) used a complete spectroscopic sample of quasars over physical scales of 23.7–29.9 kpc $h^{-1}$ from SDSS to find an excess clustering factor of $\sim 4$, consistent with the upper limit of Myers et al. (2007b), which, while $2\sigma$ below the excess found by Hennawi et al. (2006), nonetheless supports the general finding of a clustering excess which may be a result of galaxy interactions.

In this paper, we use cosmological hydrodynamic simulations which directly model the growth, accretion and feedback processes of BHs to investigate the properties and underlying causes of BH clustering. Although the simulation volume limits our analysis to BH luminosities and host group masses below those typically studied, the self-consistent modelling of BHs allows us to study the clustering behaviour without post-processing models. Additionally, the high resolution allows us to investigate clustering behaviour at extremely small scales, well below those studied with semi-analytic models, thereby providing a means of using simulations to investigate the observed small-scale excess for the first time, and provide a physical explanation for the underlying cause.

In Section 2, we describe the numerical modelling for the BHs formation and accretion (Section 2.1), the simulation parameters used (Section 2.2), the details of the subgroup finder (Section 2.3) and our method of calculating correlation functions (Section 2.4). In Section 3, we investigate the quasar clustering properties at both large and small scales, and we summarize our results in Section 4.

2 METHOD

2.1 Numerical simulation

In this study, we analyse the set of simulations published in Di Matteo et al. (2008). The code we use is the massively parallel cosmological TreePM–SPH code gadget2 (Springel 2005), with the addition of a multiphase modelling of the interstellar medium, which allows treatment of star formation (Springel & Hernquist 2003), and BH accretion and associated feedback processes (Springel, Di Matteo & Hernquist 2005; Di Matteo, Springel & Hernquist 2005). Detailed studies of the prescription for accretion and associated feedback from massive BHs and associated predictions have been presented in Sijacki et al. (2007), Di Matteo et al. (2008), Colberg & Di Matteo (2008), Croft et al. (2009) and Degraf, Di Matteo & Springel (2010). Important for our discussion is that the model has been shown to reproduce remarkably well both the observed $M_{\text{BH}}$–$\sigma$ relation and total BH mass density $\rho_{\text{BH}}$ as well as the quasar luminosity functions and their evolution in optical, soft and hard X-ray band (where our region of overlap spans 1–2 dex in luminosity). So the model, within its intrinsic limitations, appears to serve as a fair standard for representing growth, activity and evolution of massive BHs in numerical simulations. This needs to be viewed in the context of cosmological growth of BHs and not detailed accretion physics (the detailed treatment of which is completely infeasible in cosmological simulations). We note also that at least two independent teams (Booth & Schaye 2009; Johansson, Naab & Burkert 2008) have now also adopted the same modelling for BH accretion, feedback and BH mergers in the context of hydrodynamic simulations. These independent works, and in particular, the cosmological simulations by Booth & Schaye (2009) [part of the OverWhelmingly Large Simulations (OWLS) programme] have allowed to fully and independently explore the parameter space of the reference model of Di Matteo et al. (2008), as well as variations.
to some prescriptions. This large body of already existing work and investigations make this particular model somewhat of a good choice for further pursuing studies of quasar clustering.

For completeness, we here briefly review the basics of our BH accretion and feedback model and refer to Di Matteo et al. (2008) as well the references above for all details. Technically, we represent BHs with collisionless particles that are created in newly emerging and resolved groups/galaxies. To find these groups, a friends-of-friends (FoF) group finder is called at regular intervals on the fly (the time intervals are equally spaced in log a, with 2 log a = log 1.25).

Each group with a mass above \(5 \times 10^{10} h^{-1} M_\odot\) that does not already contain a BH is provided with one by turning its densest particle into a sink particle with a seed mass of \(M = 5 \times 10^7 h^{-1} M_\odot\). This seeding prescription was selected to reasonably match the expected formation of supermassive BHs either by collapse of a supermassive star to a BH with \(M_{\text{BH}} \sim M_{\text{seed}}\) (e.g. Bromm & Loeb 2003; Begelman, Volonteri & Rees 2006) or by Population III stars collapsing into \(\sim 10^9 M_\odot\) BHs at \(z \sim 30\) followed by exponential growth (Bromm & Larson 2004; Yoshida et al. 2006), reaching \(\sim M_{\text{seed}}\) by the time the group reaches \(\sim 10^{10} M_\odot\). After insertion, the BH particle grows in mass via accretion of surrounding gas according to \(\dot{M}_{\text{BH}} = 4\pi \alpha \eta G \rho v_c^2\) (Hoyle & Lyttleton 1939; Bondi & Hoyle 1944; Bondi 1952), and by merging with other BHs. Note that to allow for the initial rapid BH growth necessary to produce supermassive BHs of \(\sim 10^9 M_\odot\) at early time (\(z \sim 6\)), we allow for mildly super-Eddington accretion (consistent with models of e.g. Volonteri & Rees 2006; Begelman et al. 2006), but limit these phases to a maximum of \(3 \times M_{\text{seed}}\) to prevent artificially high values.

The accretion rate of each BH is used to compute the bolometric luminosity, \(L = \eta \dot{M}_{\text{BH}} c^2\) (Shakura & Sunyaev 1973). Here \(\eta\) is the radiative efficiency, and it is fixed at 0.1 throughout the simulation and this analysis. Some coupling between the liberated luminosity and the surrounding gas is expected: in the simulation 5 per cent of the luminosity is (isotropically) deposited as thermal energy in the local BH kernel, acting as a form of feedback energy. This parameter is fixed at 5 per cent based on earlier galaxy merger simulations such that the normalization \(M_{\text{BH}} - \sigma\) relation is reproduced (Di Matteo et al. 2005). This model of active galactic nucleus (AGN) feedback as isotropic thermal coupling to the surrounding gas, albeit simple, is likely a good approximation to any physical mechanism which leads to a shock front which isotropizes and becomes well mixed over physical scales smaller than those relevant in our simulations and on time-scales smaller than the dynamical time-scales of the haloes (see e.g. Hopkins & Hernquist 2006; Di Matteo et al. 2008, for a detailed discussion).

When DM haloes merge into a single halo, the BHs typically fall towards the centre of the new halo, eventually merging with one another. For these cosmological volumes, it is not possible to directly calculate the details of the infalling BHs at the smallest scales, so a simple merger prescription is used. Since the merging BHs are typically found in a gaseous environment at the centre of a galaxy, we assume that the final coalescence will be rapid (Makino & Funato 2004; Escala et al. 2004; Mayer et al. 2007), except in instances of BHs rapidly passing each other early in the merging of the galaxies. Thus our merger criterion is to merge BHs when they come within the spatial resolution of the simulation, but only if their relative velocity is small (comparable to the local sound speed).

In addition to having been used in galaxy merger simulations to investigate the regulation of BH growth and correlation with host galaxies (Di Matteo et al. 2005; Robertson et al. 2006; Hopkins et al. 2007a), Di Matteo et al. (2008) previously investigated the validity of this method of modelling BHs in these cosmological simulations, finding that in addition to producing an \(M_{\text{BH}} - \sigma\) relation that matches observations, the BH mass density matches values inferred from the integrated X-ray luminosity function (Shankar et al. 2004; Marconi et al. 2004) and the accretion rate density is consistent with the constraints of Hopkins, Richards & Hernquist (2007b, and Degraf et al. 2010) found that the luminosity function agrees well with observations in the range of luminosities common to both observations and our simulations.

### 2.2 Simulation parameters

Two simulation runs are analysed in this paper to allow for different volume size and resolution. The main parameters are listed in Table 1. Both runs were of moderate volume, with box sizes of side length \(33.75 h^{-1} \text{Mpc}\) (D6 simulation) and \(50 h^{-1} \text{Mpc}\) (E6). For both simulations \(N_p = 2 \times 486^3\) particles were used. The moderate box sizes prevent the simulations from being run below \(z \sim 1\) to keep the fundamental mode linear, but provide a large enough scale to produce statistically significant quasar populations. The limitation on the box sizes is necessary to allow for appropriate resolution to carry out the subgrid physics in a converged regime [for further details on the simulation methods, parameters and convergence studies, see Di Matteo et al. (2008)].

<table>
<thead>
<tr>
<th>Run</th>
<th>Box size ((h^{-1} \text{Mpc}))</th>
<th>(N_p) ((h^{-1} M_\odot))</th>
<th>(m_{DM}) ((h^{-1} M_\odot))</th>
<th>(m_{gas}) ((h^{-1} M_\odot))</th>
<th>(\epsilon) ((h^{-1} \text{kpc}))</th>
</tr>
</thead>
<tbody>
<tr>
<td>D6</td>
<td>33.75</td>
<td>(2 \times 486^3)</td>
<td>(2.75 \times 10^7)</td>
<td>(4.24 \times 10^6)</td>
<td>2.73</td>
</tr>
<tr>
<td>E6</td>
<td>50</td>
<td>(2 \times 486^3)</td>
<td>(7.85 \times 10^7)</td>
<td>(1.21 \times 10^7)</td>
<td>4.12</td>
</tr>
</tbody>
</table>

\(N_p\): total number of particles.
\(m_{DM}\): mass of dark matter particles.
\(m_{gas}\): initial mass of gas particles.
\(\epsilon\): comoving gravitational softening length.
3 RESULTS

To illustrate the distribution of quasars (as a function of their luminosity) with respect to the underlying matter distribution, in Fig. 1 we plot a slice through the D6 simulation at \( z = 1, 2, 3, 4 \), with BH positions indicated by coloured dots for four luminosity range bins:

- \( L < 10^8 \, L_\odot \) – orange;
- \( 10^8 \, L_\odot < L < 10^9 \, L_\odot \) – pink;
- \( 10^9 \, L_\odot < L < 10^{10} \, L_\odot \) – blue;
- \( L > 10^{10} \, L_\odot \) – green.

As required by our BH model, since supermassive BHs are hosted by galaxies, the quasars (particularly the most luminous sources) are located in some of the densest regions, with low redshift tending to exhibit more BHs, though at generally fainter luminosities. To characterize the relation between BH and host halo mass more precisely, in Fig. 2 we show the relation between the group halo mass and the mass of its most massive (central) BH, with colour representing the respective (instantaneous) quasar luminosity. There is a correlation between halo mass and BH mass, and to a lesser extent between halo mass and BH luminosity, with large haloes tending to host more massive, more luminous BHs than smaller haloes, albeit with significant scatter. This is due to the light curve that a BH has in our simulations (regulated by the complex hydrodynamics, see e.g. Di Matteo et al. 2008; Degraf et al. 2010). We also find that as redshift decreases, the simulation is more densely populated with BHs, which tend to be more massive and less luminous than at earlier redshift.

To study the relation between BHs and other structures, we compute the two-point correlation function \( \xi(r) \):

\[
dP = \rho_0^2 [1 + \xi(r)] dV_1 dV_2 \tag{1}
\]

(Peacock 1999), where \( dP \) is the probability of finding one object in each volume element \( dV_1 \) and \( dV_2 \), separated by a distance \( r \), with an average number density of \( \rho_0 \). In Fig. 3 we show the correlation functions of BHs found in the D6 (solid black) and the E6 (solid pink) simulations for scales between \( 10 \, \text{kpc} \, h^{-1} \) and \( \sim 10 \, \text{Mpc} \, h^{-1} \) at \( z = 1, 3, 5 \), with Poisson error bars. Note that the results from the two simulations are very similar, with the higher resolution D6 simulation showing a small boost at small scales (below \( \sim 200 \, \text{kpc} \, h^{-1} \)). To confirm the validity of the clustering results for our limited box size, we also show the DM correlation function from the

![Figure 1. An example of the distribution of BHs in the simulations: the same slice (33.75 Mpc \( h^{-1} \) across, 2 Mpc \( h^{-1} \) thick) through the D6 simulation at \( z = 1, 2, 3, 4 \). The positions of BHs in different luminosities bins (\( L < 10^8 \, L_\odot \) – orange; \( 10^8 \, L_\odot < L < 10^9 \, L_\odot \) – pink; \( 10^9 \, L_\odot < L < 10^{10} \, L_\odot \) – blue; \( L > 10^{10} \, L_\odot \) – green) are plotted on top of the gas density distribution (shown in the grey scale).](https://academic.oup.com/mnras/article/413/2/1383/1069830)
D6 simulation (inset, solid) and compare with the predicted matter correlation function using the fit of Smith et al. (2003)\(^1\) (inset, dotted) for scales corresponding to \(k < 10 h \text{ Mpc}^{-1}\) where Smith et al. (2003) quotes a 3 per cent accuracy. We note that the agreement is typically very good, with minor discrepancies beginning at scales above a few \(h^{-1} \text{ Mpc}\), consistent with the findings of Bagla & Prasad (2006). We also note the small scales investigated here should be unaffected by the simulation softening, as we restrict ourselves to scales above \(\sim 10 \text{ kpc} h^{-1}\), well above the convergence condition of \(2\epsilon\) found by Bagla & Khandai (2009).

In general, we see \(\xi_{\text{BH}}\) typically takes the form of a power law (with some possible excess at small scales at \(z = 1\)). We also divide \(\xi_{\text{BH,D6}}\) into two terms: a 1-halo term (dashed blue) produced by BH pairs occupying the same host group and the 2-halo term (dashed green) produced by pairs occupying different groups. As expected, the 2-halo term dominates at large scales (above \(\sim 300 \text{ kpc} h^{-1}\)), while at smaller scales the 1-halo term dominates, indicating that our small-scale clustering is really measuring BH properties within the scales of the host haloes. A distinction between the 1-halo and 2-halo terms is expected (as BHs are hosted by galaxies within haloes) and is consistent with the theoretical expectations (see e.g. Cooray & Sheth 2002), as well as what has been found in galaxy correlation functions (see e.g. Magliocchetti & Porciani 2003; Zehavi et al. 2004).

### 3.1 Large-scale clustering

It may be expected that BHs will cluster similarly to their host galaxies (within their haloes). To investigate the relation between BH clustering and that of their host haloes, in the left-hand column of Fig. 4 we plot the 1-halo (blue) and 2-halo (red) contributions to the correlation function (at \(z = 1, 3\) and 5) for BHs (solid lines) and galaxies (as identified by the subgroup finder described in Section 2.3) populating haloes (i.e. groups) in the specified mass ranges (dashed lines). These mass ranges were chosen to reproduce the closest agreement between \(\xi_{\text{BH}}\) and \(\xi_{\text{subgroup}}\) in the 2-halo regime at each redshift, so as to be used as an indicator of the typical halo mass for BH hosts (at each redshift). The same is shown in the right-hand column of Fig. 4, where we only include some of the most luminous BHs in the simulations (\(10^9 \text{L}_\odot < L_{\text{BH}} < 10^{10} \text{L}_\odot\), a range which is probed by observations).

For the full BH population, the typical host mass increases slightly with decreasing redshift, from \(\sim 10^{11} \text{M}_\odot\) to somewhat below \(10^{12} \text{M}_\odot\) from \(z = 5\) to \(z = 1\) respectively. When limited to the luminosity range \(10^9 \text{L}_\odot < L < 10^{10} \text{L}_\odot\), we again find increasing host mass with decreasing redshift, but with a sharper increase up to masses a few times \(10^{11} \text{M}_\odot\) at \(z = 1\) (still in the faint end of the luminosity function, see Degraf et al. 2010).

We compare the typical host mass found in this way to the mean (median) mass of the host haloes (see Table 2) for several luminosity ranges, and find that the 2-halo clustering as described above does indeed provide an estimator for the mean host halo mass at the corresponding redshift in the simulations. In addition, the table shows that for a given halo mass the luminosity of its typical BH decreases with time, particularly at low redshift (below \(z \sim 2\)–3), as seen more generally in Fig. 2. This is shown explicitly in the bottom of Table 2, where we calculate the mean and median BH luminosities found within groups of specified halo mass ranges. Note that the mean quasar luminosity actually peaks at \(z = 3\) for massive (\(M > 10^{12} \text{M}_\odot\)) groups as a result of a few highly luminous sources.

To better characterize the overall clustering strength, and in particular its luminosity dependence and evolution with redshift, we use the correlation length \(r_0\), defined as the scale at which \(\xi(r_0) = 1\) (which we calculate using a linear interpolation of \(\xi\)). In Fig. 5, we plot \(r_0\) versus \(z\) for BHs in several luminosity bins (solid coloured lines) and, for comparison, groups in several mass bins (dashed grey lines). In general, we find a weak evolution of the quasar clustering with redshift. This can be simply explained by the evolution of the

---

\(^1\)Computed using the \texttt{halofit} function of CAMB: http://lambda.gsfc.nasa.gov/toolbox/tb_camb_form.cfm (Lewis, Challinor & Lasenby 2000).

© 2011 The Authors, MNRAS 413, 1383–1394
bias of its underlying host halo masses. In particular, the correlation length for luminous ($L > 10^9 L_\odot$) BHs tends to decrease slightly as a function of decreasing redshift until $z \sim 3$, following closely the bias of the $10^{11}$–$10^{12} M_\odot$ groups (consistent with the constraints on the host masses of these BHs). At fixed mass, these groups are less biased as a function of decreasing redshift (Mo & White 2002; Bahcall et al. 2004). This is also in accord with our results from Fig. 4, that the typical host halo mass remains roughly constant for $z > 3$. For lower redshift (particularly $z < 2$), we instead see a significant upturn in $r_0$ versus $z$, corresponding to the increase in typical host halo mass, just as we found in Fig. 4 and Table 2. The lowest luminosity sources, however, show only minor change in $r_0$, corresponding to a host mass which changes only slightly with redshift (consistent with the median host masses found in Table 2).

This luminosity dependence is sufficiently weak (less than a factor of 2 increase in $r_0$ across several orders of magnitude in luminosity) to remain broadly consistent with the predictions from models that suggesting bright and faint quasars occupy similar haloes (e.g. Lidz et al. 2006; Bonoli et al. 2009). Indeed, our simulations produce complex light curves for our BHs, with luminosity varying rapidly across several orders of magnitude (see e.g. Di Matteo et al. 2008; Degraf et al. 2010). This produces significant scatter in the relation between BH luminosity and host mass, so general agreement with light-curve-based models is expected (which are indeed motivated by simulations similar to our own). None the less, as seen in Fig. 2, there remains some correlation between BH instantaneous luminosity and group mass, so a weak dependence on luminosity is expected even in this model.

### 3.2 Small-scale clustering

Although we find that the 2-halo terms for BHs and subgroups (galaxies) can be easily matched to provide a good estimator for typical host mass, there is significant discrepancy between their respective 1-halo terms (Fig. 4, blue lines). The 1-halo BH correlation function is different both in shape and amplitude to the 1-halo term of galaxies, suggesting that, unlike at large scales, BHs do not cluster like their host galaxies on small scales. Or in other words, the distribution of BHs within haloes does not follow closely that of their galaxies and hence does not trace the underlying matter distribution.

In terms of amplitude, $\xi_{BH,1h}$ can be adjusted by only considering the BHs in those groups that match the mass range constrained by the 2-halo term, thereby minimizing the suppression of $\xi_{BH,1h}$ from the numerous BHs in groups too small to contribute to the 1-halo term (due to hosting only a single BH). As expected, in this case, (Fig. 4, green line) the amplitude increases and is more in agreement with the 1-halo term of the subgroups ($\xi_{subgroup,1h}$, at least at $z = 1–3$ where the statistics are good enough).

It is, however, hard to account for the substantial difference in shape: $\xi_{BH,1h}$ follows an approximate power law, lacking the decrease in slope at small scales (below $\sim 200–300$ kpc $h^{-1}$) observed in $\xi_{subgroup,1h}$ and expected from the 1-halo clustering produced by a general Navarro–Frenk–White (NFW) profile (Navarro, Frenk & White 1996; Cooray & Sheth 2002; Zehavi et al. 2004). Thus the BHs are distributed significantly differently than an NFW profile, showing a significant boost at small scales.

We investigate the reason for this difference in the shape of the BH 1-halo term in terms of multiple BHs coexisting in a given subgroup. These BHs end up in a given subgroup as a result of mergers between their host galaxies, so that multiple BHs are expected to coexist in a remnant (until dynamical friction is able to bring them close enough together to eventually merge).

To understand the effect this has on the small-scale clustering of BHs, we calculate the contributions to $\xi_{BH,1h}$ from pairs of BHs occupying the same galaxy (we will call this the ‘1-subhalo’ term) and from pairs of BHs occupying different galaxies within the same group (‘2-subhalo’ term), in analogy with dividing the overall correlation function into its 1-halo and 2-halo terms. We note that the existence of multiple BHs within a single subgroup necessarily indicates a previous merger event, since BH particles are not inserted into galaxies which already contain a BH particle, and thus any 1-subhalo contribution is inherently a result of previous galaxy mergers.

In Fig. 6 we plot the 1-subhalo (solid blue) and 2-subhalo (solid red) components of $\xi_{BH,1h}$ together with $\xi_{subgroup,1h}$ for subgroups in groups within the mass ranges listed in Fig. 4 (dotted line). The 1-subhalo term does indeed have a steeper slope than the 2-subhalo term, and is most significant at small scales. The 1-subhalo term is most significant at low redshift, and by $z = 1$ it dominates the entire 1-halo term. This is a result of having increasingly large groups at low redshift which have also undergone a relatively large number of mergers. These indeed contain multiply occupied subgroups (see Fig. 7). Conversely, the large scales ($> 100$ kpc $h^{-1}$) tend to be
Figure 4. Correlation functions for the D6 simulation BHs (solid) and subgroups within a specified mass range (dashed), at $z = 1, 3, 5$, with 1-halo and 2-halo terms plotted separately (blue and red, respectively). The BH correlation function is plotted using all BHs (left) and using only those with $10^8 \, L_{\odot} < L_{BH} < 10^{10} \, L_{\odot}$ (right). The mass range for $\xi_{\text{group}}$ is chosen so as to find the closest agreement between $\xi_{BH, 2h}$ and $\xi_{\text{group}}$. We also plot $\xi_{BH, 1h}$ using only BHs found in host groups in this fitted mass range (dot–dashed green line).

Table 2. Mean (median) halo mass of parent group and mean (median) luminosity of daughter BHs in D6 simulation.

<table>
<thead>
<tr>
<th>BH luminosity</th>
<th>$z = 1$</th>
<th>$z = 2$</th>
<th>$z = 3$</th>
<th>$z = 4$</th>
<th>$z = 5$</th>
</tr>
</thead>
<tbody>
<tr>
<td>$L_{BH} = 10^8 , L_{\odot}$</td>
<td>39.9 (10.1)</td>
<td>24.2 (9.07)</td>
<td>19.3 (9.51)</td>
<td>14.8 (8.34)</td>
<td>12.6 (8.34)</td>
</tr>
<tr>
<td>$L_{BH} = 10^9 , L_{\odot}$</td>
<td>61.5 (18.7)</td>
<td>27.2 (9.66)</td>
<td>18.9 (8.80)</td>
<td>13.1 (7.53)</td>
<td>11.5 (7.12)</td>
</tr>
<tr>
<td>$L_{BH} = 10^{10} , L_{\odot}$</td>
<td>252 (94.1)</td>
<td>54.9 (20.4)</td>
<td>38.2 (21.1)</td>
<td>25.5 (14.9)</td>
<td>15.9 (10.3)</td>
</tr>
<tr>
<td>Group mass</td>
<td></td>
<td>Log(mean(median) BH luminosity)</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>$M_{\text{group}} = 10^{13} , M_{\odot}$</td>
<td>7.88 (7.69)</td>
<td>8.96 (8.55)</td>
<td>8.99 (8.83)</td>
<td>9.39 (9.25)</td>
<td>9.36 (9.34)</td>
</tr>
<tr>
<td>$10^{11} , M_{\odot} &lt; M_{\text{group}} &lt; 10^{13} , M_{\odot}$</td>
<td>8.66 (7.86)</td>
<td>9.09 (8.76)</td>
<td>9.39 (9.07)</td>
<td>9.54 (9.31)</td>
<td>9.71 (9.48)</td>
</tr>
<tr>
<td>$10^{12} , M_{\odot} &lt; M_{\text{group}} &lt; 10^{15} , M_{\odot}$</td>
<td>9.09 (8.23)</td>
<td>9.64 (9.19)</td>
<td>9.90 (9.42)</td>
<td>10.19 (9.64)</td>
<td>10.65 (9.88)</td>
</tr>
<tr>
<td>$10^{13} , M_{\odot} &lt; M_{\text{group}} &lt; 10^{15} , M_{\odot}$</td>
<td>9.45 (8.85)</td>
<td>10.19 (9.81)</td>
<td>11.16 (10.34)</td>
<td>10.82 (10.66)</td>
<td>10.92 (10.68)</td>
</tr>
<tr>
<td>$10^{15} , M_{\odot} &lt; M_{\text{group}}$</td>
<td>10.49 (9.51)</td>
<td>10.33 (10.20)</td>
<td>12.40 (11.49)</td>
<td>11.48 (11.54)</td>
<td>N/A</td>
</tr>
</tbody>
</table>
Solid lines: BH correlation length as a function of redshift for several luminosity bins (coloured lines). Grey dashed lines: group correlation length as a function of redshift for group mass ranges (from top to bottom) $>10^{12} M_\odot$, $10^{11} - 10^{12} M_\odot$, $5 \times 10^{10} - 5 \times 10^{11} M_\odot$, $10^{10} - 10^{11} M_\odot$.

dominated by the 2-subhalo term, since it is only at low redshift ($z \sim 1$) that there are subgroups large enough to contain a pair of BHs at such high separation. We also find that, if restricted to BHs within the same host mass range as the subgroups, the 2-subhalo term of $\xi_{BH,1h}$ matches $\xi_{subgroup,1h}$ quite closely. Thus we find that, within sufficiently large groups (such that the simulation contains enough groups hosting multiple BHs to produce a well-defined 1-halo term), $\xi_{BH,1h}$ has two distinct components: one due to BH pairs which occupy separate galaxies, exhibiting good agreement with $\xi_{subgroup,1h}$; and a steeper one caused by BH pairs which co-occupy individual galaxies as a result of previous galaxy mergers, causing a boost in the small-scale $\xi_{BH,1h}$, particularly evident at low redshift, where typical groups are largest and have undergone significant merging.

In Fig. 7 we plot the relative mass of each multiply occupied subgroup and its host group, with circles indicating central subgroups and triangles showing satellite subgroups. We clearly see that these multiply occupied subgroups tend to be the primary (central) subgroup within a given group, typically containing $\sim 65$–$70$ per cent of the total group’s mass. We also colour-code the data points to show the number of BHs within each subgroup, and see that the central subgroup of larger groups tends to contain more BHs.

To investigate the masses of BHs which populate these multiply occupied subgroups, in Fig. 8 we plot the mass of the largest (primary) BH within a given subgroup relative to the masses of the other BHs in the same subgroup, colour-coded to show the number of BHs within the subgroup. In only a few rare cases do we have more than one massive BH, while in the majority of cases we have, at most, a single massive BH with one or more smaller BHs, generally within an order of magnitude of the seed mass. This suggests that the majority of BHs in multiply occupied subgroups come from relatively small satellite subgroups (hosting correspondingly small BHs) which have fallen in and merged with the large, central subgroup, but do not grow substantially, instead remaining much less massive than the primary BH in the galaxy. Additionally, we observe that over time the fraction of BHs in the simulation located within these multiply occupied subgroups increases from 2 per cent at $z = 5$ to 15 per cent at $z = 1$, as typical groups get larger and have had more opportunity for satellite subgroups to merge with the central subgroup. This increase in typical group mass causes an increase in both the number of multiply occupied subgroups as well as an increase in the typical number of BHs found within them (as seen in Fig. 8), which produces the increased importance of the 1-subhalo term with decreasing redshift seen in Fig. 6.

We will compare the small-scale clustering from the simulations to observations in Section 3.4.

### 3.3 Quasar bias

To further characterize the clustering properties of BHs, we now consider the quasar bias as a function of scale and redshift. The bias is obtained by taking the square root of the ratio between $\xi_{BH}$ and the DM correlation function (shown as dotted lines in Fig. 4). Based on our results of the small-scale clustering, we expect the quasars to be strongly biased with respect to the DM distribution at small scales, particularly at high redshift. This general trend is clearly seen in Fig. 4, where $\xi_{DM}$ (dotted lines) increases with time (due to gravitational collapse), while $\xi_{BH}$ tends to decrease slightly (seen more clearly in Fig. 5). More importantly, we see that the BH clustering bias relative to that of DM is strongly scale-dependent, with $\xi_{BH}$ exhibiting a significant increase in clustering at small redshift.
scales (below $\sim 300 \, \text{kpc} \, h^{-1}$) due to the strong 1-halo term, whereas $\xi_{\text{DM}}$ shows only a slight increase at these small scales.

In the top of Fig. 9, we plot the scale-dependent BH bias and subgroup bias (defined as $\sqrt{\xi_{\text{BH}}/\xi_{\text{DM}}}$, $\sqrt{\xi_{\text{subgroup}}/\xi_{\text{DM}}}$, respectively) found within the hosts of the best-fitting mass ranges found in Fig. 4. Here we see that the subgroup bias levels off (as did $\xi_{\text{subgroup}}$ in Figs 4 and 6), but the 1-subhalo term causes the BH bias to continue increasing to the smallest scales probed in our simulation. To show this more clearly, the middle panel of Fig. 9 shows the bias of BHs relative to the subgroups ($\sqrt{\xi_{\text{BH}}/\xi_{\text{subgroup}}}$) for $z = 1$–5. Within a given host mass range, the BHs cluster very similarly to the subgroups (galaxies), except at the smallest scales (below $\sim 100 \, \text{kpc} \, h^{-1}$), where we again see the increased clustering caused by the multiply occupied subgroups remaining from merger events, as discussed earlier. Although we note that this small-scale bias appears to be redshift-dependent, it is actually a result of the evolution of the host mass being considered. At higher redshifts, the typical host mass is smaller, and thus fewer will have undergone subgroup mergers producing multiply occupied subgroups (as confirmed in Fig. 7), thereby making the small-scale boost less apparent. When considering behaviour for a fixed group mass (as shown in the bottom panel of Fig. 9), we see that the bias between BH and subgroup clustering is redshift-independent, and consistently exhibits a strong small-scale boost from past subgroup mergers.

### 3.4 Comparison with observations: projected correlation function

In order to compare with the observational constraints on the small-scale clustering (see Hennawi et al. 2006; Myers et al. 2008), we compute the projected correlation function $W_P(R_{\text{min}}, R_{\text{max}})$. This projected correlation function is computed using the same estimator described in Section 2.4, but the separation between points is the projected separation on to the $xy$, $xz$ or $yz$ plane, rather than the separation in three space. Although these three projections provide comparable results, we average across the three directions to avoid any potential directional bias. We then average across redshifts 1–3 (to match the observational data redshift range), weighted by the number of BHs at each redshift, and plot the result in Fig. 10, together with the data from Hennawi et al. (2006) (asterisks), Myers et al. (2008) (triangle) and the extension of the best-fitting power law...
Figure 9. Top: solid lines: BH bias defined as $\sqrt{\xi_{BH}/\xi_{DM}}$, using only BHs occupying haloes in the best-fitting mass ranges specified in Fig. 4. Dotted lines: subgroup bias defined as $\sqrt{\xi_{subgroup}/\xi_{DM}}$, using only subgroups occupying haloes in the best-fitting mass ranges specified in Fig. 4. Middle: bias of BHs relative to subgroups ($\sqrt{\xi_{BH}/\xi_{subgroup}}$) occupying haloes in the typical mass ranges found in Fig. 4. Bottom: bias of BHs relative to subgroups ($\sqrt{\xi_{BH}/\xi_{subgroup}}$) occupying haloes of mass $2\times10^{11}M_\odot$.

for the large-scale clustering found by Porciani et al. (2004) (step function). We have also plotted the projected correlation function for subhaloes found in our simulations for several host mass ranges (dotted lines).

Fig. 10 shows a remarkable agreement between the small-scale clustering of BHs from the simulations with the observations. In particular, when considering BHs within groups in the mass range of $4-8\times10^{11}M_\odot$, the small-scale boost (magenta line) matches the observed clustering very well. For completeness we also show (dashed black line) the signal expected from BHs in the hosts of the typical mass ranges found in Fig. 4. For our simulations, this is consistent with the complex light curves we derive from our direct modelling in which quasar luminosities vary over relative short time-scales for a given host (as regulated by hydrodynamical processes). This is also consistent with the models of Lidz et al. (2006).

In addition to the large-scale clustering (the 2-halo regime), our simulations allow us to study the small-scale clustering (the 1-halo term) of $\xi_{BH}$. We found that $\xi_{BH,1h}$ follows a power-law behaviour all the way to the smallest scales. The clustering of BHs at small scale is unlike that of galaxies (or DM). We showed that the 1-halo BH term can be subdivided into two components: 1-subhalo and 2-subhalo. The 1-subhalo term, $\xi_{subgroup,1h}$, represents the clustering of BHs within a galaxy and 2-subhalo that of BHs occupying different galaxies. We have shown that the 1-subhalo is the one that provides the power-law behaviour, indicating that galaxies do contain multiple BHs as a result of mergers. These galaxies tend to be the central galaxy within relatively large groups (for our simulation), generally hosting at most a single massive BH and one or more smaller BHs, likely as a result of smaller satellite galaxies merging with the large, central galaxy within the group. In the absence of these multiply occupied galaxies, $\xi_{BH,1h}$ and $\xi_{subgroup,1h}$ exhibit very close agreement, but the inclusion of these merger remnants causes a significant boost in the small-scale BH clustering. This merger-based boost is most significant at low redshift, where typical group size is largest, though we find it in sufficiently massive groups at all redshifts.
Though observational limitations make observing these scales difficult, several recent studies have found a small-scale excess at scales below \(\sim 100 \, \text{kpc} \, h^{-1}\) (Hennawi et al. 2006; Myers et al. 2008). The observed excess is in remarkable agreement to the one predicted by our simulations coming from groups approaching \(10^{12} \, M_{\odot}\), which host mostly intermediate size BHs. This suggests that multiple BHs co-occupying a subgroup at low redshifts are likely faint(ish) AGNs hosted in Milky Way size haloes that have recently undergone merging. We also note that galaxies hosting multiple AGN (Komossa et al. 2003; Gerke et al. 2007; Barth et al. 2008; Comerford et al. 2009b) or inspiralling supermassive BHs (Comerford et al. 2009a) have been found in recent studies, further supporting our conclusion of multiply occupied subgroups. Although we leave more detailed investigation of the small-scale BH pairs in our simulations (particularly with regard to the luminosities of inspiralling BHs) for a future work, we note that our finding that multiply occupied galaxies tend to host a single massive BH with one or more small BHs appears to be in keeping with the observation that most of the inspiralling BH pairs power only a single AGN (Comerford et al. 2009a). Given that, our agreement in small-scale merger-induced boost certainly reinforces the importance of galaxy mergers on the evolution of supermassive BHs. We also note this small-scale excess’ sensitivity to the host mass suggests that future small-scale studies may provide a means to constrain the typical mass of merger events between galaxies hosting BHs, with current observational data combined with our simulations suggesting groups with typical masses comparable to those probed in our simulations (from a few \(10^{11} \, M_{\odot}\) to \(10^{12} \, M_{\odot}\)) produce the multiply occupied galaxies underlying the observed small-scale excess.

We would like to point out however that there are several aspect of our modelling approach, including numerical issues, in the simulations that potentially affect our results on the small-scale clustering. We have a very simplistic prescription to determine how BHs merge with one another (imposed by the limits on the resolution that can be achieved in these cosmological boxes). The current prescription has a BH pair merge when BHs are separated by less than their smoothing length and if the BHs relative velocity is small compared to the local sound speed. Changes to this prescription could accelerate (postpone) BH mergers, which would result in a suppression (increase) of our small-scale clustering signals. It would be desirable to compare our results with other simulations which implement different prescriptions, or in the future to include more direct physical modelling of this region in higher resolution simulations. However, neither of these are currently possible. A numerical issue that may affect the results of our 1-halo term is that BHs need to be fixed to potential minima (calculated among the neighbouring particles within the smoothing length used for the accretion model) in order to avoid them leaving their subhalo due to numerical noise (and the fact that dynamical friction is hard to calculate for sink particles). However, in some instances, this may cause a BH particle in a small subhalo in orbit in a bigger group to ‘hop’ to the potential minimum of the larger group. This effect may be exacerbated in situations where the small subhalo may be stripped of gas by infalling into a larger one. These effects could artificially increase the number of BHs within large, central haloes, thereby boosting small-scale clustering. However, when we measure what fraction of BHs appear to ‘hop’ into the centre of groups experiencing an unexpected jump in their position, we find that it is only \(\sim 1–2\) per cent. Future simulations and comparison amongst different approaches (once they become available) should of course attempt to characterize these effects more specifically. We further note, however, that observational studies have indeed found cases of galaxies hosting multiple BHs (Comerford et al. 2009a), so the existence of a 1-subhalo term is expected. Additionally, as seen in Fig. 10, the projected clustering of subgroups has a fundamentally different form than the observed quasar clustering. Thus the BHs cannot simply trace their host subgroups/galaxies and still produce the observed small-scale excess, but rather a significant 1-subhalo term is expected. Additionally, as seen in Fig. 10, the projected clustering of subgroups has a fundamentally different form than the observed quasar clustering. Thus the BHs cannot simply trace their host subgroups/galaxies and still produce the observed small-scale excess, but rather a significant 1-subhalo term is required to produce the small-scale power-law behaviour.

In future work, we also plan to simulate larger volumes (which are starting to be feasible with the most advanced technology) to allow us to study clustering of AGN at larger (mass and length).
REFERENCES

Di Matteo T., Springel V., Hernquist L., 2005, Nat, 433, 604
Heavens S. et al., 2006, AJ, 131, 1

This paper has been typeset from a TeX file prepared by the author.