Evolution of star clusters in arbitrary tidal fields

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ABSTRACT

We present a novel and flexible tensor approach to computing the effect of a time-dependent tidal field acting on a stellar system. The tidal forces are recovered from the tensor by polynomial interpolation in time. The method has been implemented in a direct-summation stellar dynamics integrator (NBODY6) and test-proved through a set of reference calculations: heating, dissolution time and structural evolution of model star clusters are all recovered accurately. The tensor method is applicable to arbitrary configurations, including the important situation where the background potential is a strong function of time. This opens up new perspectives in stellar population studies reaching to the formation epoch of the host galaxy or galaxy cluster, as well as for starburst events taking place during the merger of large galaxies. A pilot application to a star cluster in the merging galaxies NGC 4038/39 (the Antennae) is presented.

Key words: methods: analytical – methods: numerical – globular clusters: general – open clusters and associations: general – galaxies: star clusters: general.

1 INTRODUCTION

The haloes of nearly all galaxies are populated by old globular clusters that presumably formed in gaseous discs at high redshift (z \sim 3–5; Kravtsov & Gnedin 2005). Young ‘populous clusters’ or ‘super-star clusters’ are found in the Large Magellanic Cloud (Hodge 1961), starburst galaxies (e.g. van den Bergh 1971), interacting galaxies and merger remnants (e.g. Whitmore & Schweizer 1995; Miller et al. 1997) and also in quiescent spirals (Larsen 2002; Figer 2008). This suggests that globular cluster formation is not unique to the early Universe and that the formation of these dense stellar systems is a common phenomenon in star formation (Elmegreen & Efremov 1997; Portegies Zwart, McMillan & Gieles 2010).

There is increasing evidence from studies of the Milky Way and the Andromeda galaxy that some globular clusters have only recently (past \sim 5 Gyr) been brought in through satellite accretion (e.g. Marín-Franch et al. 2009; Mackey et al. 2010). In order to understand the relation between the young massive clusters and the old globular clusters, i.e. their life cycle, we need to place their evolution in a cosmological context. During the formation process of galaxies through repeated accretion phases, substructures such as star clusters or dwarf satellites evolve along complex orbits in a non-static external potential. This makes their evolution difficult (if not impossible) to describe analytically.

To approach this problem numerically is also challenging because of the large range of evolutionary time-scales involved, ranging from several days for tight binaries in the cores of clusters to a Hubble time for the host galaxy (see the recent review by Dehnen & Read 2011). To be able to self-consistently model the evolution of star clusters in ‘live’ galaxies one needs to rely on a direct-tree hybrid approach (e.g. Fuji et al. 2007).

This is why most studies of the (dynamical) evolution of star clusters simplify the effect of the external tidal field by assuming a static background potential (e.g. Chernoff & Weinberg 1990; Vesperini & Heggie 1997; Fall & Zhang 2001; Baumgardt & Makino 2003; Dehnen et al. 2004; Gieles & Baumgardt 2008; Hurley et al. 2008; Peñarrubia, Walker & Gilmore 2009; Zonoozi et al. 2011). Although this is probably an adequate approximation for many purposes, it does not suffice for more complicated orbits such as those of clusters in mergers of massive galaxies or satellite galaxies that are accreted.

In light of the last point, several recent studies have adopted a (semi)analytical approach to star cluster evolution in more realistic external tides, such as the hierarchical build-up of galaxies (e.g. Prieto & Gnedin 2008) and galaxy mergers (e.g. Kruisjes et al. 2011). In here the effect of mass-loss because of stellar evolution, evaporation of stars over the tidal radius and the shock-enhanced escape of stars because of rapidly varying tidal fields (i.e. disc crossings and bulge shocks) are applied analytically. In almost all

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cases these processes are assumed to be independent of each other such that the individual resulting mass-loss rates are simply added.

An important, and often dominant, mass-loss mechanism is the relaxation-driven escape of stars, so-called evaporation. Some models assume that a constant fraction of the stars evaporate per half-mass relaxation time, independent of the orbit (e.g. Gnedin & Ostriker 1997; Prieto & Gnedin 2008). Others consider that the escape fraction depends on the galactocentric radius, often assumed to be the pericentre distance (King 1962; Innanen, Harris & Webbink 1983; Fall & Zhang 2001). The final lifetime of clusters can be quite different, depending on the details of the assumptions that are made.

Another critical disruptive agent is mass-loss due to external tidal perturbations or ‘shocks’. The related lifetime scales linearly with the density of the cluster (Spitzer 1958; Ostriker et al. 1972) and it is, therefore, that the results of semi-analytical models rely critically on what is assumed for the relation between the cluster mass and the half-mass radius; a relation which is not only affected by evaporation (i.e. because of the reduction of the mass in time), but also because of relaxation-driven expansion (Hénon 1965; Goodman 1984; Baumgardt, Hut & Heggie 2002; Gieles et al. 2010), which as far as we are aware is neglected in all semi-analytic approaches. For clusters on circular orbits in static potentials that are well within the tidal limit initially (the half-mass radius being less than a few percent of the tidal radius), it was found that the expansion phase dominates the evolution in the first half of the cluster’s lifetime, while evaporation dominates in the second half (Gieles, Heggie & Zhao 2011). In the latter stage, the cluster density adjusts to the mean (galactic) density along the orbit (Küpper et al. 2010). It is, therefore, necessary to consider both the internal evolution (relaxation) and the external effects (tides) simultaneously if one considers the entire life cycle of clusters.

Current direct N-body codes are capable of solving the N-body problem numerically for $N \sim 10^5$, together with the effects of mass-loss of the individual stars, binary interactions and tidal fields (Portegies Zwart et al. 2001; Aarseth 2003). Thanks to increasing computational power, it is now possible to combine the relaxation-driven evolution of star clusters in cosmologically motivated external conditions. In this paper, we present a new method that integrates the effect of any tidal field to the evolution of galaxy substructures that we specifically apply to star clusters. We do this by extracting the tidal tensor that contains all information about the tidal field at the location of the substructure, from galaxy simulations and subsequently ‘feed’ this to a stellar dynamics code. However, we have not yet implemented the other half of the scale coupling, as the feedback from the small scales (e.g. metal enrichment, stellar winds and supernova explosions) is not retroceded to the ambient galactic medium.

The paper is organized as follows: first, we set up the framework for the computation of the tidal acceleration by means of tensors (Section 2). The expressions found are then applied to the special cases of circular orbits: well-known expressions are retrieved using the new formalism in Section 3. The role of the galactic profile on the evolution of clusters through the escape of stars is specially explored in Section 3.4. The limitations of the analytical approach are explained in Section 4, while Section 5 presents the numerical implementation of the method to compute the tidal forces in a stellar dynamics code. A comparison with previous results obtained for idealized configurations is carried out. Applications to innovative cases are presented as the first practical illustrations of the method. Finally, the limitations and some possible future developments of our approach are discussed.

2 ANALYTICAL DESCRIPTION OF ARBITRARY TIDAL FIELDS

The main goal of the paper is to provide a general framework within which to follow the evolution of self-gravitating stellar associations in arbitrary and time-dependent tidal fields. For concreteness, in the remainder of the paper, we focus on star clusters orbiting within a galaxy in equilibrium, but the formalism can be exported to many other situations (such as dwarfs galaxies, galaxy mergers, galaxy clusters, etc.).

2.1 Tidal and effective tensors

It is convenient to work in coordinates centred on the star cluster being embedded in the background gravitational potential, as opposed to the global system’s barycentre. The tides derive from gradients in the external gravitational acceleration across the cluster. Subtracting the acceleration of the cluster’s centre of mass by the host galaxy, the relative acceleration of a member star at the position $r'$ in this frame reads

$$\frac{d^2 r'}{dt^2} = -\nabla \phi_c(r') - \nabla \phi_H(r') + \nabla \phi_H(0),$$

where $\phi_c$ and $\phi_H$ are the gravitational potentials of the star cluster and the host galaxy, respectively. The standard treatment of the background gravity in the tidal limit (e.g. Binney & Tremaine 2008) consists in regrouping the last two terms in equation (1) and performing a linear expansion by considering that $r' \ll R_G$, with $R_G$ being the distance between the cluster and the galaxy’s barycentre. It is important to recall that the linear expansion will hold when and if the Laplacian of the galaxy’s potential is small at the location of the cluster, regardless of the ratio $r'/R_G$. Bearing this in mind, the linearized equations of motion are expressed in a general form through the tidal tensor $T_i^j$ of components

$$T_i^j(r') = \left( -\frac{\partial^2 \phi_H}{\partial x^i \partial x^j} \right)_r.$$

To first order in $r'$ we have

$$\nabla \phi_H(r') = \nabla \phi_H(0) - T_i^j(r') \cdot r' + O(r'^2).$$

Substituting in equation (1):

$$\frac{d^2 r'}{dt^2} = -\nabla \phi_c(r') + T_i^j(r') \cdot r'.$$

The symmetry $T_i^j = T_j^i$ allows us to express $T_i^j$ in a diagonal form in the base of its eigenvectors $v_i$, ($i = 1–3$): the amplitude of the eigenvalues $\lambda_i$ is a measure of the strength of the tidal field along the corresponding eigenvector. When the proper base of the tensor is used to express the accelerations, the reference frame becomes non-inertial. Even so, only a rotational component at the angular frequency $\Omega$ appears because the translational component is absorbed in equation (1). The net acceleration now includes non-inertial terms from fictitious forces:

$$\frac{d^2 r}{dr^2} = -\nabla \phi_c(r) + T_i^j(r') \cdot r' - \Omega \times (\Omega \times r) - \frac{d\Omega}{dr} \times r - 2\Omega \times \frac{dr}{dr}. $$

\[ \text{circular forces} \]

\[ \text{Coriolis} \quad \text{Euler} \]

\[ \text{fictitious} \]

\[ \text{gravitational} \]

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where $r$ is the position vector in the non-inertial frame. The centrifugal acceleration can be derived from the gradient of a scalar potential

$$
\phi_\ell(r) = \frac{1}{2} (r \cdot \Omega)^2 - \frac{1}{2} \Omega^2 r^2,
$$

(6)
defined up to an arbitrary additive constant. This in turn leads to an effective tidal potential $\phi_e$ and the associated effective tidal tensor $T_e$ of components

$$
T_e \equiv T_e^i_j(r) = T_e^i_j(r) + \left( - \frac{\partial^2 \phi_e}{\partial x^i \partial x^j} \right)_r = \left( - \frac{\partial^2 \phi_e}{\partial x^i \partial x^j} \right)_r.
$$

(7)
The total acceleration becomes

$$
d^2 r \over dt^2 = -\nabla \phi_\ell(r) + T_e(r) \cdot r - \frac{d\Omega}{dt} \times r - 2 \omega \times \frac{dr}{dt}.
$$

(8)
In the diagonal form, we write the effective tensor $T_e$ as

$$
T_e(r) = \begin{pmatrix} \lambda_{e,1} & 0 & 0 \\ 0 & \lambda_{e,2} & 0 \\ 0 & 0 & \lambda_{e,3} \end{pmatrix},
$$

(9)
with the convention $\lambda_{e,1} \geq \lambda_{e,2} \geq \lambda_{e,3}$. In the rest of the paper we will refer to the three $\lambda_{e,i}$ as the effective eigenvalues.

Equation (8) cannot be simplified for a non-zero Euler acceleration. Hence, the following analytical Sections 2.2, 2.3 and 3 focus on cases where $\Omega$ is constant in time. More general configurations will be explored numerically in Section 5.

2.2 Tidal radius

The positions where the internal gravitational acceleration of the cluster is exactly balanced by all the other accelerations are called the Lagrange points $L_i$ (with $i = 1$ to 5). By convention, $L_1$ and $L_2$ fall down the galaxy–cluster axis ($L_1$ being between the two objects). The distance between the centre of the cluster and $L_1$ is referred to as the tidal radius $r_t$.

At $L_1$, it is reasonable to approximate the potential of the cluster with that of a point of mass $M_c$. Furthermore, the effective tidal acceleration¹ there is $\lambda_{e,1} r_t$. Finally, the Lagrange points are stationary in this reference frame, meaning that the Coriolis acceleration of $L_1$ is zero. With these considerations, equation (8) gives the expression of the tidal radius:

$$
r_t = \left( \frac{GM_c}{\lambda_{e,1}} \right)^{1/3}.
$$

(10)
Note that this definition applies to all galactic potentials.

The sphere of radius $r_t$ can be seen as an approximation of the physical boundary of the cluster. A more precise three-dimensional definition, called the Jacobi surface, can also be used.

2.3 Jacobi surface

The effective tidal potential derives from the linearization of equation (7):

$$
\phi_e(r) = -\frac{1}{2} r^T T_e(r) \cdot r
$$

(11)

¹ The eigenvector related to the largest eigenvalue $\lambda_{e,1}$ points towards the galaxy.

The three-dimensional surface of equipotential passing in $L_1$ is called the Jacobi surface. From equation (12), we find that the corresponding potential energy (also called critical energy) is

$$
E_\ell = \frac{3}{2} \frac{GM_c}{r_t}.
$$

(13)
The equality of equations (12) and (13) defines the equation of the Jacobi surface:

$$
0 = 2r_t^3 + \left( \frac{\lambda_{e,1}}{\lambda_{e,2}} \right)^2 x^2 + \left( \frac{\lambda_{e,1}}{\lambda_{e,3}} \right)^2 z^2 \left( x^2 + \frac{\lambda_{e,2}}{\lambda_{e,1}} y^2 + \frac{\lambda_{e,3}}{\lambda_{e,1}} z^2 \right). \quad (14)
$$

A star whose energy is exactly $E_\ell$ cannot pass through this surface, and thus can only escape through the points $L_1$ or $L_2$, where the surface is ‘opened’ (see an example in Fig. 1). At energies higher than $E_\ell$, the apertures of the surface are larger. This plays a non-trivial role in the escape rate, as discussed in Section 3.4.

3 APPLICATION TO CIRCULAR ORBITS

In this section, we apply the formulae obtained in Section 2 to the special case of circular orbits around various galaxies sitting in $(-R_0, 0, 0)$. The centrifugal term is derived from the rotation speed of the co-rotating reference frame which is, by definition, the orbital angular velocity $\Omega = (0, 0, \Omega)$:

$$
\Omega = \sqrt{\frac{GM_0(R_0)}{R_0^3}},
$$

(15)
where $M_0(R_0)$ is the mass of the galaxy enclosed within the orbital radius $R_0$, and $M_0$ (with no argument) is the total mass of the galaxy, and $M_c \ll M_0$.
The centrifugal acceleration is strictly opposed to the tidal contribution along the $y$- and $z$-axes:

$$\Omega^2 r = -\left( \frac{\partial^2 \phi_G}{\partial y^2} \right)_{R_0} r = -\left( \frac{\partial^2 \phi_G}{\partial z^2} \right)_{R_0} r,$$

so that the effective tidal eigenvalues are

$$\lambda_{e,1} = \left( \frac{\partial^2 \phi_G}{\partial x^2} \right)_{R_0}, \quad \lambda_{e,2} = \left( \frac{\partial^2 \phi_G}{\partial y^2} \right)_{R_0}, \quad \lambda_{e,3} = \left( \frac{\partial^2 \phi_G}{\partial z^2} \right)_{R_0},$$

for all circular orbits.

### 3.1 Point-mass galaxy

We first focus on the academic case of a cluster in the circular orbit around a point-mass galaxy. Using equation (17), we find that the triplet of effective eigenvalues reads

$$\{\lambda_{e,1}, \lambda_{e,2}, \lambda_{e,3}\} = \frac{GM}{R_0^3} \{3, 0, -1\}.$$

Note that when replacing this value of $\lambda_{e,1}$ in equation (10), we recover the well-known expression of the tidal radius:

$$r_t = R_G \left( \frac{M}{3M_G} \right)^{1/3} = \left( \frac{GM}{3\Omega^2} \right)^{1/3},$$

(see King 1962; Fukushige & Heggie 2000, or Binney & Tremaine 2008). The corresponding Jacobi surface is displayed in Fig. 1, and the one-dimensional projections of the total potential (solid black) are compared to those of a cluster in isolation (dashed green) in Fig. 2. The potential yields a saddle shape in $L_1$ (and in $L_2$, by symmetry).

### 3.2 Power-law galaxy

The formalism of the tidal tensor allows us to evaluate the tidal acceleration from any galactic potential. As an illustration, we now focus on power-law galactic profiles of index $\alpha < 3$ whose density is

$$\rho_0(x, y, z) = \rho_0 \left[ (x + R_G)^2 + y^2 + z^2 \right]^{-\alpha/2},$$

with $\rho_0$ being a constant. At the position of the cluster $(0, 0, 0)$, the effective eigenvalues are

$$\{\lambda_{e,1}, \lambda_{e,2}, \lambda_{e,3}\} = \frac{4\pi \rho_0}{(3 - \alpha)R_0^3} \{\alpha, 0, -1\},$$

which is comparable to the point-mass case discussed above. This sets the value of the tidal radius at

$$r_t = R_G^{\alpha/3} \left( \frac{M_\alpha(3 - \alpha)}{4\pi \rho_0 \alpha} \right)^{1/3} = \left( \frac{GM_\alpha}{\alpha \Omega^2} \right)^{1/3}.$$

The projections of the potential are plotted in Fig. 2 for $\alpha = 2.5$ and 2.0. When normalized to the tidal radius, only the $z$-component differs from the point-mass case. The impact of this difference is further explored in Section 3.4.

### 3.3 Plummer galaxy

Consider now a Plummer (1911) potential of characteristic radius $r_0$, once again centred at $(-R_G, 0, 0)$,

$$\phi_G = \frac{-GM}{\left[ r_0^2 + (x + R_G)^2 + y^2 + z^2 \right]^{1/2}}.$$

We introduce $\xi = R_G/r_0$ and evaluate the effective eigenvalues at the position of the cluster:

$$\{\lambda_{e,1}, \lambda_{e,2}, \lambda_{e,3}\} = \frac{GM_\alpha}{r_0^2 (1 + \xi^2)^{3/2}} \left\{ \frac{3\xi^2}{(1 + \xi^2)}, 0, -1 \right\}.$$

First, we consider $\xi = 1$ so that the cluster lies in the tidally extensive regime of the Plummer sphere, i.e. where the tidal contribution to the first effective eigenvalue is positive (Renaud et al. 2008). The triplet of effective eigenvalues is

$$\{\lambda_{e,1}, \lambda_{e,2}, \lambda_{e,3}\} = \frac{GM_\alpha}{r_0^2} \left\{ \frac{3\xi^2}{8}, 0, -\frac{\sqrt{2}}{4} \right\}.$$

![Figure 2. Projections of the total potential (from equation 12, with $G = M_c = r_1 = 1$), along the $x$-, $y$- and $z$-axes, for the cluster in the circular orbit in the $xy$-plane around a point-mass galaxy, galaxies with a power-law density profile of slope 2.5 and 2.0, and a Plummer galaxy with the cluster in the tidally extensive ($\xi = 1$) and compressive regimes ($\xi = 0.66$, see the text). Along the $x$- and $y$-axes, all these projections are identical. A circle marks the position of the Lagrange point $L_1$. The potential of the cluster in isolation ($r^{-1}$, green) is also plotted, for comparison.](https://academic.oup.com/mnras/article-abstract/418/2/759/1068470/762-F.-Renaud-M.-Gieles-and-C.-M.-Boily)
which gives a tidal radius\(^2\) of

\[
r_t = r_0 \left( \frac{M_\star}{\frac{5}{3} \rho \lambda} \right)^{1/3} = \left( \frac{GM_\star}{\frac{7}{5} r_t^2} \right)^{1/3}.
\]

(26)

Secondly, by decreasing \(\xi\), we shift the cluster towards the centre of the Plummer potential. The tidal contribution to \(\lambda_{e,1}\) first tends towards zero; a value reached for \(\xi = 2^{-1/3}\). For smaller values of \(\xi\), the cluster lies in the core region of the galactic potential, and thus is in compressive tidal mode (see Renaud et al. 2009): the tidal acceleration acts in the same direction as the internal gravitational acceleration of the star cluster. Following Chandrasekhar (1942), Appendix A demonstrates that the centrifugal contribution always compensates the compressive tidal acceleration on circular orbits, so that the Lagrange points still exist and thus the tidal radius can be defined even in compressive tidal mode.

The expression of \(\lambda_{e,1}\) in equation (24) reveals that a given tidal radius can be obtained at two different galactic radii: one in the tidally extensive mode and another in the compressive mode. The compressive counterpart of our extensive example (equation 26) is obtained for \(\xi \approx 0.66\). Both cases are plotted in Fig. 2. Interestingly, the potential in the extensive case (\(\xi = 1\)) is strictly identical to that found with a power-law galaxy of index \(\alpha = 1.5\): the evolution of identical star clusters set in these two configurations would be impossible to distinguish.

The comparison between these five configurations emphasizes that the potential well of a star cluster not only depends on the tidal radius, but also on the three-dimensional shape of the effective tensor, which varies from galaxy to galaxy. In the next subsection, we examine the implication of this on the escape rate of stars from the cluster.

3.4 Escape rate

For circular orbits, the escape rate varies, to first order, as \(M_\star \propto \Omega\) (e.g. Lee & Ostriker 1987; Baumgardt & Makino 2003; Gieles & Baumgardt 2008). However, the constant of proportionality contains details of the shape of the density profile of the galaxy. Tanikawa & Fukushige (2010) recently demonstrated the importance of this secondary effect. They did this by considering clusters in galaxies with different power-law density profiles (as in our Section 3.2). The constant \(\rho_0\) was varied such that the tidal radius was the same in all cases. They found, somewhat counter-intuitively, that in this setup the clusters with the lowest orbital angular velocity \(\Omega\) had the highest escape rate. In the following, we confirm and generalize their result, using the tensor formalism.

To escape from the cluster, a star needs (1) to be able to fly outside of the Jacobi surface and (2) to exit in a way not to fall back in. The first condition implies that the total energy \(E = v^2/2 + \phi\) of the star must exceed the Jacobi energy: \(E > E_j\). The second condition tells us that a candidate escaper, which fulfills the first condition, can still be trapped in the potential well of the cluster for many crossing-times (Fukushige & Heggie 2000; Baumgardt 2001): the potential barrier keeps increasing with the distance in all directions except along the galaxy–cluster axis linking \(L_1\) and \(L_2\) (see Fig. 2). Therefore, the actual escapers are stars (1) with exceeding energy and (2) flying through the apertures in the corresponding equipotential surface

\[t_{\text{esc}}(E) = \frac{2}{\pi^{1/3}} C \left( \frac{GM_\star}{E - E_j} \right)^{1/2} \left( \frac{1 - \frac{\rho_0}{\lambda_{e,1}}} {\lambda_{e,1}} \right). \]

(27)

where \(C\) is a dimensionless constant which depends on the intrinsic properties of the cluster and which we compute by numerical integration (\(C \approx 0.4\), see Appendix B). Baumgardt (2001) wrote the (time-dependent) dissolution time-scale \(t_{\text{diss}}\) of the entire cluster as

\[t_{\text{diss}} \propto \frac{1}{t_{\text{h}}} \left( \frac{E_j}{E} \right)^{1/4} \left( E = 2E_j \right),\]

(28)

where \(t_{\text{h}}\) is the half-mass relaxation time. This relation holds for homologous clusters for which the half-mass radius scales linearly with the tidal radius. We can then write the dependence of the dissolution time-scale for circular orbits on the galactic parameters as

\[t_{\text{diss}} \propto \frac{1}{t_{\text{h}}} \left( \frac{1 - \frac{\rho_0}{\lambda_{e,1}}} {\lambda_{e,1}} \right)^{1/8} \left( E = 2E_j \right).\]

(29)

We observe not only a first-order dependence of the dissolution time-scale on the tidal radius, but also a second-order effect due to the shape of the galactic potential. This relation confirms and extends the conclusion of Tanikawa & Fukushige (2010) to any galaxy: for a given tidal radius, a highly negative ratio of the effective eigenvalues corresponds to a slow dissolution, as illustrated in Fig. 4.

The proportionality factor in equation (29) depends on the properties of the cluster \((t_{\text{h}}, C, M_\star)\), which all evolve with time. Thus, \(t_{\text{diss}}\) is an instantaneous estimate of the dissolution time-scale and should not be mistaken with the actual lifetime of the cluster. Moreover, the evolution of these properties depends on the tidal field and is very involved (if not impossible) to estimate analytically.

\[\text{Figure 3. Projection of the surface of equipotential in the } xz\text{-plane, for an energy } E = 0.95E_j \text{ and for the five galaxies considered in Fig. 2. The green circle represents the tidal radius.} \]

\[\text{around } L_1 \text{ and } L_2. The size of these apertures depends on the excess of energy and on the shape of the tidal field (see Fig. 3).} \]

To evaluate the escape time, we seek the flux of stars through the apertures. We compute it by repeating the calculation applied by Fukushige & Heggie (2000) to point-mass galaxies, but here for any galactic profile by means of the effective eigenvalues. Details of the derivation are given in Appendix B. We find that the time-scale for escape for a star with an excess of energy \(E > E_j\) is

\[t_{\text{diss}} \propto \frac{1}{t_{\text{h}}} \left( \frac{E_j}{E} \right)^{1/4} \left( E = 2E_j \right),\]

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In Section 5.3, numerical experiments show indeed that the actual lifetime can be very different from the analytical value of $t_{\text{diss}}$ given by equation (29).

4 NON-CIRCULAR ORBITS

As stated earlier, writing the tidal tensor in its proper base makes the reference frame non-inertial: the fictitious accelerations must be added. In the cases of linear or circular orbits, the Euler force is null and the formalism of the effective tensor allows for a simple mathematical description of the net acceleration (see Section 3). However, for time-dependent rotations, the Euler effect must be added: the effective tensor does not suffice to describe the full impact of the galaxy, and the formalism loses its advantage of analytical simplicity.

However, nothing forbids to write the tidal tensor in the inertial frame, where it is non-diagonal, and to compute the tidal acceleration numerically. The major advantage of such a method is that the centrifugal, Euler and Coriolis accelerations, which are difficult to evaluate along complex orbits, are not required anymore. That is, the tidal tensor computed in the inertial reference frame fully represents the galactic acceleration on star clusters and allows for a numerical treatment in any galaxy and along any orbit. The next section explains how this method can be implemented in an $N$-body code.

5 N-BODY SIMULATIONS: NBODY6tt

When expressed in the inertial reference frame, the tidal tensor contains all the information on the effect of the galaxy on its cluster. Therefore, the equations of motion of all the stars of the cluster can be solved numerically, for any tidal field. In this section, we briefly present one implementation and a suite of tests, before applying the method to innovative cases.

5.1 Retrieving the external force

The simulations of the star clusters are done with the stellar dynamics code NBODY6 and its version for Graphics Processing Unit (GPU; Aarseth 2010). Several cases of galactic potentials already exist in NBODY6 and have been widely used in previous studies. For testing purposes, we use these features and refer to them as built-in methods. On top of these pre-existing tools, we have modified NBODY6 to include the tidal forces by means of the tidal tensor: this new version of the code is called NBODY6tt.

Before the simulation, the tidal tensor is computed in the inertial reference frame (equation (2)): its nine components are sampled along the orbit of the cluster within the galaxy, either analytically or by means of independent galactic simulations (see Renaud et al. 2009; Renaud 2010). The sampling frequency is chosen to ensure that the high-frequency features of the tidal acceleration, both in terms of intensity and orientation, are recovered. A table of sampled tensors is then passed to NBODY6tt. During the simulation of the cluster, the components of the tensor are quadratically interpolated whenever the gravitational force on a particle needs to be updated. For a star of mass $m$ at the position $\{x^i\}$ with respect to the centre of the cluster, the tidal force, whose $i$th component reads

$$F_i = m \sum_j T_{ij} x^j,$$

is added to the gravitational force due to the $N - 1$ other particles (see Aarseth 2003, for details on the solving of the $N$-body problem).

As soon as physically time-dependent processes (like stellar evolution) are not involved, the entire study remains scale-free. The units adopted below are the cluster's mass, velocity, and length, denoted $M$, $v$, and $L$, respectively. These units are then passed to NBODY6tt. During the simulation of the cluster, the tidal force, whose $i$th component reads

$$F_i = m \sum_j T_{ij} x^j,$$

is added to the gravitational force due to the $N - 1$ other particles (see Aarseth 2003, for details on the solving of the $N$-body problem).

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Table 1. One possible scaling of the simulations in physical units.

<table>
<thead>
<tr>
<th>Quantity</th>
<th>$N$-body units</th>
<th>Physical units</th>
</tr>
</thead>
<tbody>
<tr>
<td>Cluster scale</td>
<td></td>
<td></td>
</tr>
<tr>
<td>Mass ($M_c$)</td>
<td>1</td>
<td>$8 \times 10^3$ M$_\odot$</td>
</tr>
<tr>
<td>Mass of a particle ($m$)</td>
<td>1/8000</td>
<td>1 M$_\odot$</td>
</tr>
<tr>
<td>Virial radius ($r_v$)</td>
<td>1</td>
<td>1 pc</td>
</tr>
<tr>
<td>Characteristic radius ($r_{0,c}$)</td>
<td>$3\pi/16$</td>
<td>0.59 pc</td>
</tr>
<tr>
<td>Half-mass radius ($r_h$)</td>
<td>$r_{0,c}/\sqrt{2M/\pi}$</td>
<td>0.77 pc</td>
</tr>
<tr>
<td>Crossing time ($t_{\text{cr}}$)</td>
<td>$2\sqrt{2}$</td>
<td>0.47 Myr</td>
</tr>
</tbody>
</table>

| Galactic scale            |                |                 |
| Mass ($M_G$)              | $1.25 \times 10^6$ | $10^{10}$ M$_\odot$ |
| First effective eigenvalue | $[3.75, 0.13] \times 10^{-3}$ | $[135.0, 5.0] \times 10^{-3}$ Myr$^{-2}$ |
| Orbital radius ($R_G$)    | $[1, 3] \times 10^3$ | $[1, 3]$ kpc    |
| Orbital period ($t_{\text{orb}}$) | $[178.3, 927.7, 504.8]$ | $[29.6, 154.0, 83.8]$ Myr |

For the orbits A and B, respectively.

For the orbits A, B and C, respectively.

Figure 5. Top-right panel: orbits of the clusters in the galactocentric frame. Other panels: evolution of the 10, 50 and 90 per cent Lagrange radii of the cluster on the three orbits, computed by the built-in method (black) and by NBODY6tt (colour). On the top of each panel, crosses mark the passages of the cluster at its left-most position in the top-right panel (i.e. at pericentre for orbit C). The core-collapse phase is well visible at $t \sim 1700–2000$.

The evolution of some Lagrange radii of the cluster is plotted in Fig. 5: before core collapse, the relative differences ($1 - r_{\text{NBODY6}}/r_{\text{NBODY6tt}}$) remain below 5 per cent, in all cases. The differences increase after core collapse, but not systematically, probably because of the formation of binaries which is sensible to numerical and $N$-body noises. As a complement, Fig. 6 plots the evolution of the number $N$ of stars in the cluster, normalized to its initial value. In both measurements ($r_{\text{Lagr}}$ and $N$), the agreement between the two approaches is very good for all the orbits. In particular, along the elliptical orbit C, the expansion of the outermost layers of the cluster and the increased mass-loss near the pericentre passages is well reproduced by the new method, both in terms of time (epoch, delay and duration) and amplitude. These tests demonstrate that the interpolation scheme used to evaluate the tidal tensor at any time and the computation of the force done by NBODY6tt allow us to retrieve the results obtained with well-tested methods, at a high level of accuracy.
F. Renaud, M. Gieles and C. M. Boily

par-C. M. ≈ ξ with shows that the actual lifetime of
and 32
4 < N < 2011 The Authors, MNRAS ν 10 d ≃ NBODY6 λ 418,
ρ 6 0.66 is ·) and density (see equation 28 with the usual scal-
statistics of our
and the lifetimes measured numerically for our five
run is
is shown on the logarithmic scale in Fig. 8(b). The

Figure 6. Number of stars in the cluster along the orbits A, B and C around a
point-mass galaxy, normalized to its initial value. The black curves represent
the solution using the built-in method of NBODY6, while the coloured ones
are associated with the use of the tidal tensor. The symbols mark the passage
of the cluster at the positions marked in the top-right panel of Fig. 5.

The evolution of \( N \) is presented in Fig. 7. The lifetimes of the clusters are ordered
as predicted in Section 3.4. We note however
that, for example, the lifetime of the cluster in the Plummer galaxy
at \( \xi = 0.66 \) is \( \approx 2.2 \) times longer than that around the point-mass
galaxy, while equation (29) predicts a value of \( \approx 1.57 \). The discrepancy
is due to a different evolution of the internal properties of the
center. C and \( M_c \), which are enclosed in the constant of pro-
portionality of equation (29). Furthermore, because of the different
shape of the potential (i.e. the second-order effect in equation 29),
the escape of stars does not occur at the same rate in all galaxies,
which modifies \( M_c \) and the tidal radius itself (i.e. the first-order
effect) differently from galaxy to galaxy.

5.4 Fully arbitrary tidal field

As a last step towards generality, this subsection presents the results
obtained for a complex and highly time-dependent tidal field. The
orbit chosen is extracted from a simulation of the Antennae galaxies
(NGC 4038/39), a prototypical major merger. The galactic run,
the parameters and the orbit are described in Renaud et al. (2009,
fig. 8, orbit B, see also their fig. 3): the cluster starts orbiting in the
disc of NGC 4038 at \( \sim -76 \) kpc (on average) from the galactic
centre; then it is ejected by the first galactic pericentre passage into
the intergalactic bridges before falling back into the central region
and remaining there for the rest of the merger. With the Antennae
being a real, observed object, the units are now scaled according to
the galactic simulation (based on the observed spatial extension of
the tidal tails and the peak radial velocity). The NBODY6tt run is
arbitrarily started 100 Myr before the first pericentre passage of the
two galaxies. The orbit of the cluster has also been integrated within
its host galaxy (NGC 4038) in isolation, as a reference simulation.
In both cases, the cluster is set up identically to those of the previous
sections (see Table 1). Our physical scaling makes it comparable in mass
(8000 \( M_\odot \)) and density (\( \sim 2000 M_\odot \) pc\(^{-3}\) within the half-
sphere-radius) to Westerlund 1 or NGC 3603 (Portegies Zwart et al.
2010).

The maximum eigenvalue \( \lambda_1 \) of the tidal tensor is plotted in
Fig. 8(a), in the case of the merger (red) and of the isolated galaxy
(blue). The cluster is in compressive mode for
\( \lambda_1 < 0 \). The centrifugal term\(^5 \) \( \Omega^2 \) is shown on the logarithmic scale in Fig. 8(b). The
peaks denote the velocity kicks the cluster receives when it is grav-
itionally slingshot. Fig. 8(c) displays the ratio of the density of the
cluster and the local density of the galaxy.\(^6 \) Finally, the evolution
of the number of stars in the cluster is displayed in Fig. 8(d), and
compared to that of the cluster in isolation (green).

At \( t \sim 180 \) Myr, the cluster is in the extensive regime of the tidal
bridges, marked by a slow increase of \( \lambda_1 \); the mass-loss accelerates

\[ \sum_i \lambda_i = -\nabla^2 \phi(r) = -4\pi G \rho(r). \]

\(^5\) To evaluate the centrifugal term \( \Omega^2 \), we use the

\[ \Omega(t) \approx \frac{\text{acos}[\psi(t) - \psi(t - dt)]}{dt}. \]

\(^6\) The local density of the galaxy is given by the trace of the tidal tensor,
through Poisson’s law:

\[ \sum_i \lambda_i = -\nabla^2 \phi(r) = -4\pi G \rho(r). \]
with respect to that of the cluster in the isolated galaxy. During
the second galactic pericentre passage ($t \simeq 285$ Myr), the tidal
field is mostly extensive, which once again, enhances the escape of
stars. Later, the galaxies have merged and the cluster is orbiting the
remnant. Its orbital period ($\sim 140$ Myr) is clearly visible not only
in both $\lambda_1$ and $\Omega^2$, but also in the mass-loss which is accelerated
at pericentre, in a comparable fashion to the elliptical test case
of Fig. 6. The rapid variations of $\lambda_1$ near the pericentre passages
exist because the cluster flies at high speed in a highly asymmetric
potential, which can be momentarily compressive at the location
of the cluster.

Along this particular orbit, the local galactic density is, on average,
about three orders of magnitude smaller than that of the cluster,
but the ratio of the two is peaked at several epochs. Therefore, the
tidal field only affects the cluster during precise and short periods of
time: the cluster is dense, robust enough to remain mildly disturbed
on the long time-scale. In other words, the lifetime of this cluster is
almost independent of the merger; only its ‘life style’ differs from
the case of the isolated galaxy. However, the analysis of this single
case is not statistically relevant to conclude that the merger has no
secular impact on its clusters. The properties of the cluster population,
in particular the cluster age function, rely on many parameters
(structural and orbital), and we leave their study to a forthcoming
paper.

6 SUMMARY, LIMITATIONS AND CONCLUSIONS

In this contribution, we propose a new formalism to describe the
tidal field in $N$-body simulations by the means of tensors. Although
the analytical approach rapidly becomes too involved for accelerated
motions, we have derived the expressions of quantities representing
the effect of the tides on stellar systems for circular orbits,
with no restriction on the shape of the external potential. The main
results of this study are as follows.

(i) The use of the tidal and effective tensors allows us to simplify
the representation of the problem, without loss of information (Sec-
tion 2.1). Useful quantities like the tidal radius (equation 10), the
energy of the Jacobi surface (equation 13) or the escape time-scale
(equation 27) can be easily computed for a given cluster.

(ii) The tidal radius is a first-order approximation of the effect of a
galaxy on a star cluster. The three-dimensional shape of the galactic
potential has a second-order effect (Section 3.4 and equation 29).

(iii) For a given tidal radius, a point mass is the most efficient
galactic profile in dissolving a cluster. Shallower potentials lead to
longer dissolution times (Fig. 4).

(iv) A cluster in compressive tidal mode loses its stars more
rapidly than in isolation, but significantly more slowly than if it was
in the extensive mode (Section 3.4).

(v) The knowledge of the evolution of the cluster parameters
(half-mass radius, mass and energy distribution) is key to estimate
the cluster lifetime (Section 3.4 and Appendix B).

(vi) For non-circular orbits, the analytical approach becomes
very involved so that general and simple formulae do not exist
(Section 4).

To overcome this issue, we have developed and implemented a
numerical method called $\texttt{NBody6tt}$ which computes the evolution of
$N$-body models of star clusters in any tidal field, by means of
the pre-calculation of inertial tidal tensors. This method has been
successfully tested and applied to innovative cases. Our main con-
cclusions are as follows.

(i) The tidal force felt by the stars of a star cluster can be accu-
rately computed by evaluating the tidal tensor at all time by means
of quadratic interpolation (Section 5.2).

(ii) The dependence of the dissolution time on the three-
dimensional shape of the galactic potential highlighted in Section
3.4 is confirmed by our numerical experiments. The numerical
method does not suffer from the lack of knowledge of the evolution
of the cluster properties that one has to face in a (semi)analytical
approach (Section 5.3).

(iii) Our implementation has been applied to the complex case
of a cluster in the major merger of the Antennae galaxies (Sec-
tion 5.4). Its mass-loss reflects the nature of the time-dependent
tidal field experienced along the orbit. In particular, the alternation
of extensive and compressive tidal modes strongly affects the
instantaneous dissolution rate, over a time-scale of several $\times 10^3$ yr.
This preliminary study shows the way to a very wide range of possible applications. However, one should keep in mind that our study is limited in several aspects. On the one hand, our cluster simulations are gas-free. Therefore, we cannot address the important point of the early life of the star cluster, prior to gas expulsion (first \( \sim 10^6 \) yr of the cluster’s life), and we limit our study to initially relaxed systems. On the other hand, the second half of the galaxy–cluster coupling is not taken into account by our method: the feedback from stellar evolution, the escape of stars in the field and the formation of long tidal tails or streams are not implemented. Although they have a limited impact on the evolution of the cluster itself, it would be important to monitor their effects at larger scale (e.g. the metal enrichment of the interstellar medium). Furthermore, the mass-loss experienced by a cluster could affect its orbit within its host galaxy and thus change the tidal field. In our approach, the scale decoupling does not allow us to account for such effect.

To conclude, we have shown that a star cluster plunged in a time-dependent tidal field leads to a complex evolution, out-of-reach of (semi-)analytical approaches. A numerical method like the one proposed by NODY6t will provides a framework for future explorations of the role of the tides on star clusters. Among others, the use of a stellar mass function, primordial binaries and high-order stars, stellar evolution and stellar mass-loss are as many lines of investigation to be pursued on top of the evolution in a background, time-dependent, galactic potential.

In a forthcoming paper, we will use more detailed galactic simulations including a prescription on the star formation, to compute the tidal tensors of an entire cluster population. NODY6t will help us to derive the cluster- mass and age functions and their evolution, in various type of galaxies, in particular in mergers.

ACKNOWLEDGMENTS

We thank Douglas Heggie for the interesting discussions that greatly helped to improve our understanding of the topic, and Jorge Peñarrubia, Dominique Aubert and the referee for their input. Sverre Aarseth is warmly acknowledged for making NODY6 available, as well as Keigo Nitadori for the development of the GPU part of the code. FR thanks a visiting grant of the Institute of Astronomy of Cambridge where part of this work was done, as well as support from the EC through grant ERC-StG-257720; MG acknowledges the Royal Society for financial support.

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APPENDIX A: EXISTENCE OF THE LAGRANGE POINTS

The Lagrange points \( L_1 \) and \( L_2 \), which define the tidal radius \( r_t \), exist when the internal gravitational acceleration of the cluster is balanced by the effective tidal acceleration:

\[
\frac{GM_c}{r_t^2} = \lambda_\text{e,1} r_t.
\]  

(A1)

Along a circular orbit of radius \( R_G \), one gets

\[
\frac{GM_c}{r_t^2} = \left(-\frac{\partial^2 \Phi}{\partial r^2}\right)_{R_G} + \Omega^2 r_t, 
\]

which can be re-written by introducing the epicycle frequency \( \kappa \):

\[
\frac{GM_c}{r_t^2} = - (\kappa^2 - 3\Omega^2) + \Omega^2 r_t. \]

(A3)

This tells us that the tidal radius can be defined for \( \kappa^2/\Omega^2 < 4 \), which is always the case since the maximum value \( \kappa^2/\Omega^2 = 4 \) is

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reached for homogeneous mass distributions. To conclude, on a circular orbit, the centrifugal acceleration always compensates the tidal component, even in compressive mode, so that the Lagrange points \( L_1 \) and \( L_2 \) always exist.

**APPENDIX B: COMPUTATION OF THE ESCAPE TIME**

In this appendix, we generalize the expression of the escape time derived by Fujishige & Heggie (2000) for a cluster in a circular orbit around a point-mass galaxy to the case of any galactic potential. First, the origin of the coordinates is shifted to \( L_1 \) and the total potential is expanded to second order, so that

\[
\phi - E_I = \frac{\lambda_{e,1}}{2} \left[ -3\lambda^2 + y^2 + z^2 \left( 1 - \frac{\lambda_{e,1}}{\lambda_{e,1}} \right) \right].
\]  

(B1)

The flux of phase volume across the new \( x = 0 \) is expressed as

\[
\mathcal{F}(E) = \int_{x>0} \delta \left( \phi + \frac{v^2}{2} - E \right) \dot{x} \, dy \, dz \, dy \, dz,
\]  

(B2)

where the dot indicates derivation with respect to time and \( \delta \) is the Dirac function. We change variables to \( w : \dot{x} \mapsto \phi + v^2/2 - E \) so that \( dw = \dot{x} \, dx \) and integrate the Dirac function over \( w \) to get

\[
\mathcal{F}(E) = \int_{x>0} \delta \, dy \, dz \, dy \, dz,
\]  

(B3)

with integration boundaries satisfying

\[
2(E - E_I) - \lambda_{e,1} \left[ y^2 + z^2 \left( 1 - \frac{\lambda_{e,1}}{\lambda_{e,1}} \right) \right] > 0.
\]  

(B4)

We change the remaining four variables into the hyper-spherical coordinates \( \{ R, \theta, \tau, \psi \} \), i.e.

\[
\begin{align*}
y &= \lambda_{e,1}^{1/2} R \cos \theta \\
z &= \left( \frac{\lambda_{e,1}}{\lambda_{e,1}} \left( 1 - \frac{\lambda_{e,1}}{\lambda_{e,1}} \right) \right)^{-1/2} \frac{R \sin \theta \cos \tau}{R > 0} \\
\hat{y} &= R \sin \theta \sin \tau \cos \psi \\
\hat{z} &= R \sin \theta \sin \tau \sin \psi
\end{align*}
\]

(B5)

so that the condition equation (B4) becomes

\[
2(E - E_I) - R^2 > 0.
\]  

(B6)

The determinant of the Jacobian matrix of the transformation gives the hyper-volume element:

\[
dy \, dz \, dy \, dz = \frac{R^2 \sin^2 \theta \sin \tau}{\lambda_{e,1} \sqrt{1 - \frac{\lambda_{e,1}}{\lambda_{e,1}}}} \, dR \, d\theta \, d\tau \, d\psi.
\]  

(B7)

The flux is finally

\[
\mathcal{F}(E) = \frac{2 \tau^2 (E - E_I)^2}{\lambda_{e,1} \sqrt{1 - \frac{\lambda_{e,1}}{\lambda_{e,1}}}}.
\]  

(B8)

The total flux is \( 2 \mathcal{F} \) because stars can escape through apertures around two Lagrange points.

Similarly, the phase-space volume can be written as

\[
\mathcal{V} = \int \delta \left( \phi + \frac{v^2}{2} - E \right) \, d^3r \, d^3v.
\]  

(B9)

We first take out the angular part of the velocity and integrate the Dirac function over \( v \):

\[
\mathcal{V} = 4\pi \int \sqrt{2(E - \phi)} \, d^3r.
\]  

(B10)

Defining the dimensionless quantities

\[
\left\{ \begin{align*}
\Psi^* &= (E - \phi) \, r_i/(GM_c) \\
\gamma^* &= r_i/r_c
\end{align*} \right.
\]  

and substituting them in equation (B10) yields

\[
\mathcal{V} = 4\pi \sqrt{2} (GM_c)^{1/2} r_i^{3/2} C,
\]  

(B12)

where

\[
C = \int \sqrt{\Psi^*} \, r^2 \, dr^*.
\]  

(B13)

is a dimensionless quantity which describes the intrinsic properties of the cluster. Instead of integrating over the entire solid angle, we note that the flux \( \mathcal{V} \) is non-zero only at the vicinity of the Lagrange points \( L_1 \) and \( L_2 \). In a first approximation, we may consider a non-zero flux only at the exact position of \( L_1 \) and \( L_2 \). In that case, we replace the angular dependence of the previous integral with Dirac functions so that only two angular directions remain. That is, \( C \) becomes

\[
C = 2 \int \sqrt{\Psi^*} \, r^2 \, dr^*.
\]  

(B14)

For King profiles with the usual parameter \( \Psi_\sigma/\sigma^2 \) ranging from 3 to 12, we found (by means of numerical integrations) that \( C \) takes values ranging from 0.38 to 0.39 with a maximum reached for \( \Psi_\sigma/\sigma^2 \approx 8.5 \).

By using the general definitions given in Section 2, we find

\[
\mathcal{V} = 4\pi C \sqrt{2} (GM_c)^{3/2} \lambda_{e,1}^{-5/6}.
\]  

(B15)

It follows that the time-scale for escape with the energy \( E \) is

\[
\tau_{esc}(E) = \frac{\mathcal{V}}{2 \mathcal{F}(E)},
\]  

(B16)

which gives equation (27).

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