Serendipitous discovery of a thin stellar stream near the Galactic bulge in the Pan-STARRS1 \(3\pi\) Survey

Edouard J. Bernard,\(^1\) Annette M. N. Ferguson,\(^1\) Edward F. Schlafly,\(^2\) Mohamad Abbas,\(^3\) Eric F. Bell,\(^4\) Niall R. Deacon,\(^2\) Nicolas F. Martin,\(^2,5\) Hans-Walter Rix,\(^2\) Branimir Sesar,\(^2\) Colin T. Slater,\(^4\) Jorge Peñarrubia,\(^1\) Rosemary F. G. Wyse,\(^6\) William S. Burgett,\(^7\) Kenneth C. Chambers,\(^7\) Peter W. Draper,\(^8\) Klaus W. Hodapp,\(^7\) Nicholas Kaiser,\(^7\) Rolf-Peter Kudritzki,\(^7\) Eugene A. Magnier,\(^7\) Nigel Metcalfe,\(^8\) Jeffrey S. Morgan,\(^7\) Paul A. Price,\(^7\) John L. Tonry,\(^7\) Richard J. Wainscoat\(^7\) and Christopher Waters\(^7\)

\(^1\)SUPA, Institute for Astronomy, University of Edinburgh, Royal Observatory, Blackford Hill, Edinburgh EH9 3HJ, UK
\(^2\)Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany
\(^3\)Astronomisches Rechen-Institut, Zentrum für Astronomie der Universität Heidelberg, Mönchhofstr. 12–14, D-69120 Heidelberg, Germany
\(^4\)Department of Astronomy, University of Michigan, 500 Church St, Ann Arbor, MI 48109, USA
\(^5\)Observatoire Astronomique de Strasbourg, Université de Strasbourg, CNRS, UMR 7550, 11 rue de l’Université, F-67000 Strasbourg, France
\(^6\)Department of Physics and Astronomy, The Johns Hopkins University, 3400 North Charles Street, Baltimore, MD 21218, USA
\(^7\)Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA
\(^8\)Department of Physics, Durham University, South Road, Durham DH1 3LE, UK

Accepted 2014 June 6. Received 2014 June 5; in original form 2014 May 26

ABSTRACT

We report the discovery of a thin stellar stream found in Pan-STARRS1 photometry near the Galactic bulge in the constellation of Ophiuchus. It appears as a coherent structure in the colour-selected stellar density maps produced to search for tidal debris around nearby globular clusters. The stream is exceptionally short and narrow; it is about 2.5 long and 6 arcmin wide in projection. The colour–magnitude diagram of this object, which harbours a blue horizontal-branch, is consistent with an old and relatively metal-poor population ([Fe/H] \(\sim -1.3\)) located 9.5 \(\pm 0.9\) kpc away at \((l, b) \sim (5^\circ, +32^\circ)\), and 5.0 \(\pm 1.0\) kpc from the Galactic centre. These properties argue for a globular cluster as progenitor. The finding of such a prominent, nearby stream suggests that many streams could await discovery in the more densely populated regions of our Galaxy.

Key words: surveys – globular clusters: general – Galaxy: halo – Galaxy: structure.

1 INTRODUCTION

Ever since the discovery of massive tidal tails emerging from the globular cluster Palomar 5 (Odenkirchen et al. 2001), the Milky Way field population has been the subject of intense scrutiny to find more of these cold stellar streams. The main incentives for finding streams are their potential use to constrain the shape and mass of the Milky Way dark matter halo (e.g. Johnston, Law & Majewski 2005; Koposov, Rix & Hogg 2010; Varghese, Ibata & Lewis 2011; Peñarrubia, Koposov & Walker 2012; Lux et al. 2013; Vera-Ciro & Helmi 2013; Deg & Widrow 2014), their ability to constrain the existence of mini-haloes (Ibata et al. 2002; Siegal-Gaskins & Valluri 2008; Yoon, Johnston & Hogg 2011; Carlberg 2012), as well as the fossil record they provide of the mass assembly history of the Milky Way.

While a few streams have been revealed by kinematics alone (e.g. Helmi et al. 1999; Newberg, Yanny & Willett 2009; Williams et al. 2011), the vast majority were found by searching for coherent stellar overdensities in the homogeneous, wide-field photometric catalogue provided by the Sloan Digital Sky Survey (SDSS; York et al. 2000) (e.g. Belokurov et al. 2006, 2007; Grillmair 2006a,b, 2009, 2012; Grillmair & Dionatos 2006; Grillmair & Johnson 2006; Bonaca, Geha & Kallivayalil 2012; Grillmair et al. 2013), and more recently in VST ATLAS (Koposov et al. 2014). As a result of SDSS’s predominant high-latitude, Northern hemisphere coverage, the known streams are located far from the Galactic disc and bulge that were mostly avoided by these surveys.

\*E-mail: ejb@roe.ac.uk
The Pan-STARRS1 (PS1; Kaiser et al. 2010) 3π Survey, observing the whole sky visible from Hawaii, has the advantage of providing some of the first deep imaging of the dense inner regions of our Galaxy. With a sky coverage spanning twice that of the SDSS footprint, it offers the possibility to survey these less well-studied areas to seek new tidal streams as well as further extensions of already known ones.

In this Letter, we report the discovery of a very thin stellar stream located in the inner Galaxy close to the Galactic bulge. It was found serendipitously when analysing PS1 data of wide areas around nearby globular clusters for the presence of tidal debris. We briefly describe the PS1 survey in Section 2, and present the new stream and measurements of its properties in Section 3. A summary is given in Section 4.

2 THE PS1 3π SURVEY

The PS1 3π Survey (Kaiser et al. 2010; Chambers et al., in preparation) is being carried out with the 1.8 m optical telescope installed on the peak of Haleakala in Hawaii. Thanks to the 1.4-gigapixel imager (Onaka et al. 2008; Tonry & Onaka 2009) covering a 7 deg² field of view (∼3.3 diameter), it is observing the whole sky north of δ ≳ −30° in five optical to near-infrared bands (g_p1,r_p1,i_p1,z_p1,y_p1; Tonry et al. 2012b) up to four times per year. The exposure time ranges from 30 to 45 s, leading to median 3σ limiting AB magnitudes of 21.9, 21.8, 21.5, 20.7 and 19.7 for individual exposures in the g_p1,r_p1,i_p1,z_p1,y_p1 bands, respectively (Morganson et al. 2012). At the end of the survey, the 12 or so images per band will be stacked, increasing the depth of the final photometry by ∼1.2 mag (Metcalfe et al. 2013).

The individual frames are automatically processed with the Image Processing Pipeline (Magnier 2006) to produce a photometrically and astrometrically calibrated catalogue. A detailed description of the general PS1 data processing is given in Tonry et al. (2012a). The analysis presented in this Letter is based on the photometric catalogue obtained by averaging the magnitudes of objects detected in individual exposures (Schlafly et al. 2012). At the end of the survey, the point source catalogue will be based on detections in the stacked images, leading to a significantly deeper photometry and better constraints on the tidal streams.

The catalogue used here was first corrected for foreground reddening by interpolating the extinction at the position of each source using the Schlafly et al. (2014) dust maps with the extinction coefficients of Schlafly & Finkbeiner (2011). In this part of the sky, E(B − V) ranges from 0.17 to ∼1.2 (see Fig. 1). We then cleaned the catalogue by rejecting non-stellar objects using the difference between point spread function and aperture magnitudes, as well as poorly measured stars by keeping only objects with a signal-to-noise ratio of 10 or higher.

3 A NEW STREAM IN OPHIUCHUS

The stream appears as a coherent structure in the maps showing the stellar density of objects with colours and magnitudes corresponding to the old, metal-poor main-sequence turn-off (MSTO) of nearby globular clusters. The colour and magnitude cuts were then refined based on the distribution of stars in the colour–magnitude diagram (CMD) of the stream region (see below). The resulting map is shown in Fig. 1 in both equatorial and Galactic coordinates, along with the corresponding reddening maps from Schlafly et al.

![Image](https://academic.oup.com/mnrasl/article-abstract/443/1/L84/1055239/1055239)

Figure 1. Top: density of stars with dereddened colours and magnitudes consistent with the main-sequence turn-off of an old and metal-poor population at heliocentric distances of 8–12 kpc (see selection box in Fig. 2), shown in equatorial (left) and Galactic (right) coordinates. Darker areas indicate higher stellar density. Bottom: reddening maps of the corresponding fields, derived from Pan-STARRS1 stellar photometry (see Schlafly et al. 2014). The colour scale is logarithmic; white (black) corresponds to E(B − V) = 0.17 (0.58). The stellar density and reddening maps have been smoothed with a Gaussian kernel with full-width at half-maximum of 12 and 6 arcmin, respectively. The stream is located close to the centre of each panel. The thick lines in the right-hand panels trace the best-fitting great circle containing the stream.
Assuming \((m - M)_0 = 14.44 \pm 0.02\) for NGC 5904 (Coppola et al. 2011) yields a true distance modulus of \((m - M)_0 = 14.9 \pm 0.2\) (i.e. 9.5 \pm 0.9 kpc) for the Ophiuchus stream. Finally, if we assume that the Sun is located 8 kpc from the Galactic Centre (GC), we find a Sun–GC–stream angle of \(\sim 89\) deg, placing the stream almost directly above the Galactic bulge at a Galactocentric distance of 5.0 \pm 1.0 kpc.

To estimate the total luminosity of the stream, we used IAC-STAR (Aparicio & Gallart 2004) to generate the CMD of an old population (11.5–12.5 Gyr) containing \(10^5\) stars with a metallicity comparable to that of NGC 5904 and with the same depth as our observations – stars fainter than our detection limit have a negligible contribution to the total magnitude. In Fig. 2, we find that there are about 500 more stars in panel (a) containing the stream than in the comparison field (panel b). Summing the luminosity of a sample of 500 stars extracted randomly from the synthetic CMD gives the total flux. We repeated this step \(10^4\) times, extracting between 300 and 700 stars each time, and found a total magnitude and luminosity of \(M_V = -3.0 \pm 0.5\) and \(L_V = 1.4 \pm 0.6 \times 10^4 L_\odot\), respectively. These are comparable to some of the fainter halo clusters (e.g. Whiting 1, Terzan 9, Palomar 1, Palomar 13: Harris 1996, 2010 Edition), although a significant fraction of the stars may have already been stripped off and lie further out along the stream orbit.

The extent of the stream on the sky was calculated by reprojecting the stellar maps to a new spherical coordinate system \((\Lambda, B)\) with the pole located at \((\alpha, \delta) = (184:32, +77:25)\) \([(l, b) = (125:37, +39:72)]\). In this system the stream approximately lies along the equator – shown by the red lines in Fig. 1 – making it easier to measure its width and length. In Fig. 3, we show the distribution of M斯特罗stars across the \(\Lambda\) and \(B\) dimensions. The left-hand panel shows the stellar density across the stream, where all the M斯特罗stars in the range \(-1^\circ < \Lambda < 1^\circ\) have been used. The stream is detected as a significant overdensity at \(B \sim 0^\circ\). The profile is best fitted by a Gaussian with \(\sigma = 2.99\) arcmin \pm 0.33 arcmin, corresponding to a full-width at half maximum (FWHM) of 7.0 arcmin \pm 0.8 arcmin, i.e. \(19 \pm 2\) pc at the distance of the stream. However, we note that in Fig. 1, the stream appears slightly more curved than the great circle shown, implying that the intrinsic width may be even smaller. The right-hand panel shows the profile along the length of the stream, for which we used the M斯特罗stars in a narrow strip (12 arcmin) centred on the stream. To estimate the background, we selected M斯特罗stars in six identical strips at higher and lower \(B\); their average and dispersion are shown as the blue line and the shaded area. The on-stream histogram shows a significant overdensity within the range \(-1^\circ - 0^\circ < \Lambda < 1^\circ\), and is therefore about 2.5 long (i.e. \(\sim 400\) pc in projection).

Table 1 summarizes the estimated properties of the stream: we list the approximate coordinates of its centre, the heliocentric and Galactocentric distances, the extent on the sky, and the estimated luminosity.

4 DISCUSSION AND CONCLUSIONS

We have identified a new stellar stream in the constellation of Ophiuchus, a part of the sky that has rarely been searched for streams because of the high stellar density and significant differential and foreground reddening. Both the morphology of the M斯特羅 and the presence of a blue HB are typical of an old and metal-poor population \(\gtrsim 10\) Gyr old, [Fe/H] \(\lesssim -1.0\). These properties, along with the small width of the stream and absolute magnitude suggest a globular cluster as progenitor.
Table 1. Summary of the stream properties.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>RA (J2000.0)</td>
<td>16:07:12</td>
</tr>
<tr>
<td>Dec. (J2000.0)</td>
<td>−06:55:30</td>
</tr>
<tr>
<td>l</td>
<td>4:53</td>
</tr>
<tr>
<td>b</td>
<td>+31:69</td>
</tr>
<tr>
<td>(m−M)0</td>
<td>14.9 ± 0.2</td>
</tr>
<tr>
<td>Median E(B−V)</td>
<td>0.23</td>
</tr>
<tr>
<td>Heliocentric distance</td>
<td>9.5 ± 0.9 kpc</td>
</tr>
<tr>
<td>Galactocentric distance</td>
<td>5.0 ± 1.0 kpc</td>
</tr>
<tr>
<td>Width (FWHM)</td>
<td>7.0 arcmin ± 0.8 arcmin (19 ± 2 pc)</td>
</tr>
<tr>
<td>Length</td>
<td>~2.5 (~400 pc)</td>
</tr>
<tr>
<td>MV</td>
<td>−3.0 ± 0.5</td>
</tr>
<tr>
<td>LV</td>
<td>1.4 ± 0.6 × 10^3 L⊙</td>
</tr>
</tbody>
</table>

We find that the stream is exceptionally short (~2.5, i.e. ~400 pc, in projection) compared to all the other streams found so far, that are usually several tens of degrees long. For comparison, the shortest stream known to date is the Pisces–Triangulum stream, which has been traced over ~15° (i.e. ~7 kpc) on the sky (e.g. Bonaca et al. 2012; Martin et al. 2014). We have explored shifting the MSTO selection box in magnitude to account for possible effects of differential reddening residuals and extension of the stream along the line of sight, but failed to detect any overdensity beyond the current extent. This experiment did reveal a possible distance gradient along the stream, with the eastern tail being closer to the Sun, although the apparent change in MSTO magnitude could simply be a consequence of the differential reddening.

Surprisingly, neither the stellar density along the stream nor a careful visual inspection of the images reveals a potential remnant of the progenitor. This suggests that it has already been completely disrupted. Given that the length of a tidal stream is a function of the time since the stars became unbound, a short stream may indicate that the progenitor has been disrupted only recently. However, this scenario is hard to reconcile with the lack of an obvious progenitor in the vicinity of the stream. Another possibility is that we are observing the stars of a fully disrupted cluster at apogalacticon on a highly elliptical orbit: at this point of the orbit, unbound stars tend to clump together because of the slower orbital velocity (e.g. Dehnen et al. 2004; Küpper, Lane & Heggie 2012). The data currently available are not sufficient to reliably trace the orbit of the progenitor, which would help us understand its past evolution and likely fate; radial velocities and proper motions of a sample of stream members will be crucial for this purpose.

Spectroscopic metallicities of stream stars will also shed light on the nature of the progenitor, and help understand how such events may contribute to the stellar populations in the central regions of the Galaxy. For example, recent spectroscopic surveys of the outer Galactic bulge have revealed the presence of a significant number of metal-poor stars ([Fe/H] < −1; e.g. Gonzalez et al. 2011; García Pérez et al. 2013). Their origin may be linked to tidal stripping events such as the one we are witnessing with this stream.

ACKNOWLEDGEMENTS

EJB and AMNF are grateful to Douglas C. Heggie and Anna Lisa Varri for enlightening discussions. The authors would like to thank the anonymous referee for a prompt report and useful comments. This research was supported by a consolidated grant from the Science Technology and Facilities Council. EFS and NFM acknowledge support from the DFG’s grant SFB881 (A3 ‘The Milky Way System’. NFM gratefully acknowledges the CNRS for support through PICS project PICS06183. The research leading to these results has received funding from the European Research Council under the European Union’s Seventh Framework Programme (FP 7) ERC Grant Agreement no. [321035].

The PS1 Surveys have been made possible through contributions of the Institute for Astronomy, the University of Hawai’i, the Pan-STARRS Project Office, the Max-Planck Society and its participating institutes, the Max Planck Institute for Astronomy, Heidelberg and the Max Planck Institute for Extraterrestrial Physics, Garching, The Johns Hopkins University, Durham University, the University of Edinburgh, Queen’s University Belfast, the Harvard–Smithsonian Center for Astrophysics, the Las Cumbres Observatory Global Telescope Network Incorporated, the National Central University of Taiwan, the Space Telescope Science Institute, the National Aeronautics and Space Administration under Grant No. NNX08AR22G issued through the Planetary Science Division of the NASA Science Mission Directorate, the National Science Foundation under Grant No. AST-1238877 and the University of Maryland.
This work has made use of the IAC-STAR Synthetic CMD computation code. IAC-STAR is supported and maintained by the computer division of the Instituto de Astrofísica de Canarias.

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