Gravitational lensing in WDM cosmologies: the cross-section for giant arcs

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ABSTRACT

The nature of the dark sector of the Universe remains one of the outstanding problems in modern cosmology, with the search for new observational probes guiding the development of the next generation of observational facilities. Clues come from tension between the predictions from Λ cold dark matter (ΛCDM) and observations of gravitationally lensed galaxies. Previous studies showed that galaxy clusters in the ΛCDM are not strong enough to reproduce the observed number of lensed arcs. This work aims to constrain the warm dark matter (WDM) cosmologies by means of the lensing efficiency of galaxy clusters drawn from these alternative models. The lensing characteristics of two samples of simulated clusters in the Λ warm dark matter and ΛCDM cosmologies have been studied. The results show that even though the cold dark matter (CDM) clusters are more centrally concentrated and contain more substructures, the WDM clusters have slightly higher lensing efficiency than their CDM counterparts. The key difference is that WDM clusters have more extended and more massive subhaloes than CDM analogues. These massive substructures significantly stretch the critical lines and caustics and hence they boost the lensing efficiency of the host halo. Despite the increase in the lensing efficiency due to the contribution of massive substructures in the WDM clusters, this is not enough to resolve the arc statistics problem.

Key words: gravitational lensing: strong – Methods: numerical – cosmology: theory – dark matter.

1 INTRODUCTION

The standard model of cosmology (cold dark matter + dark energy) is found to be in good agreement with the observational data on large scales of the Universe (e.g. Cole et al. 2005; Percival et al. 2010; Komatsu et al. 2011; Addison, Hinshaw & Halpern 2013). However, there are a few discrepancies between observations and the cold dark matter (CDM) model. For instance, the Λ cold dark matter (ΛCDM) cosmology predicts that many more satellites must be around Milky Way-sized galaxies than the observed satellites (Klypin et al. 1999; Moore et al. 1999). This problem could be resolved by adopting warm dark matter (WDM) model as an alternative to the standard model (e.g. Colín, Avila-Reese & Valenzuela 2000; Bode, Ostriker & Turok 2001; Lovell et al. 2014).

Gravitational lensing has revealed another discrepancy between observation and the current cosmological model. Galaxy clusters are observed to produce more lensed giant arcs than predicted by the ΛCDM model (for a review see Meneghetti et al. 2013). This discrepancy is known as the arc statistics problem and was firstly addressed by Bartelmann et al. (1998). Further, work from Wambsganss, Bode & Ostriker (2004) showed that the lensing optical depth for giant arcs is strongly dependent on the source redshift due to the very steep cluster mass function. Contrary to Bartelmann et al. (1998), these authors have found that the abundance of giant arcs in the ΛCDM model agrees with that of the observed clusters when using a wider range of source redshift. However, later research by Li et al. (2005) confirmed the arc statistics problem and found that the optical depth increased at a slower rate with the source redshift than reported in Wambsganss et al. (2004). A recent study from Meneghetti et al. (2011) has confirmed the findings of Bartelmann et al. (1998). However the new study shows that the arc...
statistics discrepancy between the lensing efficiency of the observed clusters and the ΛCDM clusters is about a factor of 2 rather than one order of magnitude. We note that the lensing efficiency can be significantly boosted during the cluster mergers, where the lensing cross-section can increase by one order of magnitude or more on a time-scale of ~0.1 Gyr (Torri et al. 2004). It has also been found that the ΛCDM clusters produce smaller Einstein radii than those observed (e.g. Broadhurst & Barkana 2008; Zitrin et al. 2011). The abundance of giant arcs is found to be sensitively dependent on the concentration of matter in their cores, and hence provides a statistical constraint on the density profile of galaxy clusters (e.g. Wu & Hammer 1993; Oguri, Taruya & Suto 2001). The optical depth for giant arcs could be boosted drastically by increasing the normalization of power spectrum (Fedeli et al. 2008). Puchwein & Hilbert (2009) investigated the contribution of structures along the line of sight (los) to the strong lensing efficiency. They found that the abundance of giant arcs for individual clusters could be increased by up to ~50 per cent, showing that its contribution becomes more significant for clusters of lower masses and sources at higher redshifts. Baryonic processes can steepen the central density profile of galaxy clusters and hence increase lensing efficiency by a small factor. However, that is not enough to resolve the discrepancy observed (e.g. Puchwein et al. 2005; Rozo et al. 2008; Wambsganss, Ostriker & Bode 2008; Mead et al. 2010; Killeider et al. 2012).

These studies have all used CDM, however, the fundamental nature of dark matter is still a subject of debate. The dark matter in the Universe must be non-baryonic, but the important question that needs to be answered is: is the dark matter hot, cold or warm? Hot dark matter scenario has been ruled out in the early 1980s as the free streaming scale is found to be too large and hence galaxies would not form (White, Frenk & Davis 1983). On the other hand, the structure formation in the Universe based on the CDM scenario has been a great success at describing the large scales of the Universe (e.g. Blumenthal et al. 1984; Davis et al. 1985). The Λ warm dark matter (ΛWDM) cosmology has started to receive more attention in the last several years as it agrees better with observations on small scales (e.g. Lovell et al. 2012; Schneider et al. 2012; Libeskind et al. 2013). The differences between the structure formation of the CDM and WDM have been outlined by Bode et al. (2001). These authors have pointed out that the WDM structures have larger core radii, lower core densities and less abundance of substructures.

Therefore, it is only natural to study whether WDM is a solution to the arc statistics problem. In this work, we explore both ΛCDM and ΛWDM cosmologies to improve our understanding about the differences of structures between the two cosmologies and how those differences affect the lensing efficiency to produce giant arcs. Naively, one would expect that the WDM clusters produce smaller Einstein radii and cross-sections than their CDM counterparts as they are less concentrated and contain fewer substructures. This should result in smaller convergence and shear fields which in turn lead to a lower lensing efficiency. However, our results show that the WDM clusters are slightly more efficient than the CDM analogues due to some physical differences between these two cosmologies. To confirm these physical differences, we compare in a companion study (Elahi et al. 2014) the mass accretion and internal properties of the WDM and CDM clusters.

Our paper is organized as follows. In Section 2, we describe the numerical methods used in this work. We then present the main results of this study (Section 3) and a brief description on why the WDM clusters are slightly stronger lenses than the CDM counterparts (Section 4). Finally, we discuss our results and conclude in Section 5.

### 2 Numerical Methods

#### 2.1 Simulations

We study 10 pairs of clusters extracted from zoom simulations (see Table 1 for their bulk properties at $z = 0$). These zoom simulations used a parent simulation of $L_{\text{box}} = 150 h^{-1} \text{Mpc}$ containing $128^3$ particles in the ΛCDM model ($h = 0.7, \Omega_0 = 0.3, \Omega_\Lambda = 0.7, \sigma_8 = 0.9$). Clusters with masses of $> 10^{14} h^{-1} \text{M}_\odot$ were identified in the parent simulation using AMIGA’s Halo Finder (AHF), which uses a overdensity threshold approach to identify haloes (AHF; cf. Knollmann & Knebe 2009). The overdensity used to define the halo corresponded to the so-called virial radius $r_{\text{vir}}$, where the mean interior density is $\Delta_{\text{vir}}$ times the critical density of the Universe at that redshift, $\rho_c(z) = 3H^2(z)/8\pi G$, where $H(z)$ and $G$ are the Hubble parameter and the gravitational constant, respectively. For 10 such clusters, we identified all particles within a radius of $\sim 3r_{\text{vir}}$ of the cluster at $z = 0$ in the parent simulation and determined their initial positions using an inverse Zel’dovich transformation to obtain the particle positions at $z = \infty$, from which we determined the spatial extent of the initial Lagrangian volume. This volume defines the central region of a multilevel mask for the high resolution region. The resampled Lagrangian regions are initialized using ΛCDM and ΛWDM cosmologies, which primarily differ in the power included at small scales. We present here a brief description of the ΛWDM model, for more details see Elahi et al. (2014). Our ΛWDM model used a 0.5 keV thermally produced dark matter particle (Bode et al. 2001), which results in a suppression of growth for halo with $M \leq M_{\text{min}} = 2.1 \times 10^{13} \text{M}_\odot$, the so-called half-mode mass scale where the WDM power spectrum is 1/4 that of the CDM one (Schneider et al. 2012). All simulations are run with GADGET2, a TreePM code (Springel 2005) and the zoom simulations used a gravitational softening length based on Power et al. (2003), i.e. $\epsilon_{\text{soft}} = 4 r_{\text{vir}}/\sqrt{N_{\text{vir}}}$.

#### 2.2 Ray tracing method

Mock lensing maps for two samples of simulated clusters in WDM and CDM cosmologies were produced using the ray tracing method. Mapping the light rays from the lens plane to the source plane can be done using the lensing equation:

$$\beta = \theta - \alpha(\theta),$$

### Table 1. Cluster properties at $z = 0$. $M_{\text{vir}}$ is the virial mass of the halo, expressed in units of $10^{15} h^{-1} \text{M}_\odot$, assuming $\Delta_{\text{vir}} = 200$; $r_{\text{vir}}$ is the virial radius, in units of $h^{-1} \text{Mpc}$; and $N_{\text{DM}}$ is the number of dark matter particles within $r_{\text{vir}}$.  

<table>
<thead>
<tr>
<th>$M_{\text{vir}}$ ($10^{15} h^{-1} \text{M}_\odot$)</th>
<th>$r_{\text{vir}}$ ($h^{-1} \text{Mpc}$)</th>
<th>$N_{\text{DM}}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>C1 1.023 57</td>
<td>2.06</td>
<td>681 684</td>
</tr>
<tr>
<td>C2 0.744 62</td>
<td>1.85</td>
<td>507 416</td>
</tr>
<tr>
<td>C3 0.622 94</td>
<td>1.74</td>
<td>423 850</td>
</tr>
<tr>
<td>C4 0.521 51</td>
<td>1.64</td>
<td>357 024</td>
</tr>
<tr>
<td>C5 0.495 27</td>
<td>1.61</td>
<td>322 581</td>
</tr>
<tr>
<td>C6 0.438 71</td>
<td>1.55</td>
<td>309 802</td>
</tr>
<tr>
<td>C7 0.422 37</td>
<td>1.53</td>
<td>298 265</td>
</tr>
<tr>
<td>C8 0.406 59</td>
<td>1.51</td>
<td>297 888</td>
</tr>
<tr>
<td>C9 0.403 26</td>
<td>1.51</td>
<td>269 994</td>
</tr>
<tr>
<td>C10 0.413 51</td>
<td>1.52</td>
<td>292 005</td>
</tr>
</tbody>
</table>
where $\beta$ and $\theta$ are the angular positions of light rays on the source and lens planes, respectively, and $\alpha(\theta)$ is the deflection angle of light rays on the lens plane, which is given by

$$
\alpha(\theta) = \frac{1}{\pi} \int \kappa(\theta') \frac{\theta - \theta'}{|\theta - \theta'|^2} d^2\theta',
$$

where $\kappa$ is the convergence. For gravitational lenses under the thin lens approximation, the convergence is simply proportional to the projected surface mass density, $\kappa = \frac{\kappa_{\text{crit}}}{\kappa_{\text{max}}}$, where $\kappa_{\text{crit}}$ is the critical surface mass density that depends on the angular distance between observer and lens $D_L$, between observer and source $D_S$ and between lens and source $D_{LS}$:

$$
\kappa_{\text{crit}} = \frac{c^2}{4\pi G} \frac{D_S}{D_L D_{LS}}.
$$

Note that the thin lens approximation hold as clusters, the gravitational lenses studied here, are much smaller than the above-mentioned distances. For this study, we place our lens and source planes at redshifts of $z_L = 0.3$ and $z_S = 2.0$, respectively.

The deflection angle is evaluated by multiplying the convergence map with the kernel from equation (2) in Fourier space, that is

$$
\alpha(\theta) = \frac{1}{\pi} [\kappa(\theta) \ast K(\theta)],
$$

where

$$
K(\theta) = \frac{\theta}{|\theta|^2}.
$$

We implement a zero-padding method to overcome the periodicity of Fourier transform when convolving the convergence with the kernel function.

For our clusters, we use only particles within $= 2.5$ Mpc of the cluster centre to calculate the convergence by projecting them on to two 2D grids: the particles within 0.5 Mpc are projected on to a small high resolution grid of $1024 \times 1024$ such that the angular resolution is 0.22 arcsec and the rest of the particles are projected on to a larger lower resolution grid of 5 Mpc of $1024 \times 1024$. The projected surface mass density is smoothed using a truncated Gaussian kernel of size of 5 $h^{-1}$ kpc in order to overcome the numerical noise due to the discreetness of $N$-body simulation. The contribution of particles in the outer grid to the deflection angle and shear fields of the inner grid has been taken into consideration by implementing a bilinear interpolation scheme between both maps. This technique was also used in Kiledar et al. (2012) and its advantage is that it actually produces high resolution lensing maps for the core of clusters where the strong lensing regime can happen.

The magnification of an image is characterized by the Jacobian matrix that can be written in terms of the convergence and the two components of shear $\gamma_1$ and $\gamma_2$ as follows:

$$
\mathbf{A} = \frac{\partial \beta}{\partial \theta} = \frac{1}{2} \begin{pmatrix}
\frac{\partial^2 \psi}{\partial \theta_1^2} & \frac{\partial^2 \psi}{\partial \theta_1 \partial \theta_2} \\
\frac{\partial^2 \psi}{\partial \theta_2 \partial \theta_1} & \frac{\partial^2 \psi}{\partial \theta_2^2}
\end{pmatrix}
= \begin{pmatrix}
1 - \kappa - \gamma_1 & -\gamma_2 \\
-\gamma_2 & 1 - \kappa + \gamma_1
\end{pmatrix},
$$

where $\psi$ is the lensing potential:

$$
\psi = \frac{1}{\pi} \int \kappa(\theta') \ln |\theta - \theta'| \, d^2\theta'.
$$

The total magnification of an image is given by $A = \det(\mathbf{A})$. Regions of high magnification on the lens plane occur whenever the determinant of Jacobian matrix vanishes. The two curves associated with this criteria are the radial critical curve where the radial eigenvalue $\lambda_r = 1 - \kappa + \gamma$ goes to zero and images are radially elongated with respect to the curve and the tangential critical curve (a.k.a. Einstein curve) where the tangential eigenvalue $\lambda_t = 1 - \kappa - \gamma$ vanishes and images are tangentially elongated with respect to the curve. The corresponding curves to the critical lines on the source plane are the radial and tangential caustics.

3 RESULTS

3.1 Mock lensing maps

Fig. 1 shows the convergence maps for all clusters. Here the green and blue lines are the critical lines on the lens plane and the caustics on the source plane, respectively. This figure illustrates the effect of the larger substructures in the WDM version of some clusters.
Figure 1. The convergence maps for CDM and WDM clusters by considering the rotation matrix that produces the median value of cross-section for WDM clusters, the green lines are the critical lines on the lens plane and the blue lines represent the corresponding caustics on the source plane. The side length of each grid is 1 Mpc and the colour bar has been set to be in the range [0:3] for all panels.
on the critical lines and caustics. Note that here we have chosen to show the orientation that corresponds to the median value of cross-section for WDM clusters in the following figures. Similar maps can be done for the lensing shear and deflection angle.

### 3.2 Lensing cross-sections and Einstein radii

The length-to-width ratio $L/W$ of the corresponding images is measured by means of the eigenvalues of the magnification tensor $\lambda_r$ and $\lambda_t$. Two types of images are expected: images elongated in the radial direction of critical line which occur when $|\lambda_t|$ is larger than $|\lambda_r|$ and images elongated in the tangential direction with respect to critical line that happen when $|\lambda_r|$ is larger than $|\lambda_t|$. So, any image with $|\lambda_t|$ or $|\lambda_r|$ greater than some physically meaningful threshold $\eta$, would be classified as a tangential or radial arc, respectively. This technique for computing the lensing cross-section was firstly proposed by Fedeli et al. (2006). The elongation threshold throughout this work is chosen to be $\eta = 7.5$ as previously considered in the literature (e.g. Puchwein et al. 2005; Redlich et al. 2012). Higher elongation thresholds lead to a smaller number of giant arcs and less elongation thresholds are very sensitive to the ellipticity of sources (Killedar et al. 2012).

Fig. 2 shows the probability density function of cross-sections of both CDM and WDM samples. The upper panel shows the probability of cross-section due to all arcs while the middle and lower panels show the probabilities of cross-section due to tangential and radial arcs, respectively. The error bars represent the statistical error for each bin and the subplots show the difference between the two distributions. This figure illustrates that the WDM clusters are slightly stronger lenses than the CDM counterparts even though they are less concentrated and contain fewer substructures. The figure also illustrates that the WDM clusters have a more significant tail of large cross-sections for radial arcs, which means that the radial arcs are more common in the WDM clusters than in the CDM counterparts.

This difference is also seen in the distribution of Einstein radii, as shown in Fig. 3. We see that the WDM clusters produce slightly larger Einstein radii than their CDM analogues.

As we use the same orientations with respect to the los for both cosmologies, we plotted the lensing cross-sections and Einstein radii of the WDM version as a function of those two lensing quantities for the CDM counterpart, as shown in Figs 4 and 5. Again, these two figures illustrate that the WDM version of our sample of clusters produces higher lensing efficiency than their CDM analogues for most of the 1500 los as they lie above the unity line. We found that 64 and 60 per cent of the los produce higher lensing cross-sections and Einstein radii, respectively, in the WDM version of clusters.

In Fig. 6, we show the tight correlation between the lensing cross-sections and Einstein radii for CDM and WDM clusters. The colour bar of the distributions represents the mean shear on the grid for each los. The figure illustrates that the WDM clusters can produce stronger shear field than the CDM counterparts even they have fewer substructures. This figure also illustrates that the lensing cross-section and Einstein radius are strongly dependent on the lensing shear, we also found a similar behaviour when the distributions are coloured with the mean convergence on the grid. For this reason, the lensing properties of simulated clusters must be studied by considering different los. Each los for a particular cluster has a unique convergence and shear fields due to the triaxiality of haloes.
Gravitational lensing in WDM cosmologies

Figure 3. The probability density function of Einstein radii for CDM sample (blue) and WDM sample (red). The blue, red and black arrows point to the median Einstein radius of the CDM and WDM samples as well as the observed sample in Meneghetti et al. (2011). The error bars represent the statistical error for each bin and the subplot shows the difference between the two distributions.

Figure 4. The lensing cross-sections of the 1500 los for the WDM version of all clusters versus the cross-sections of the CDM counterparts, the black line is the unity line.

and the existence of substructures, hence we get a unique Einstein radius and cross-section for individual los. The green squares in the lower panel of Fig. 6 represent the observed sample of clusters from Meneghetti et al. (2011). Despite the slight enhancement in the lensing efficiency when considering WDM cosmologies, it appears that none of our simulated clusters can reproduce the observed data of four clusters in Meneghetti et al. (2011) which have very high Einstein radius and cross-section for giant arcs.

The median cross-section and Einstein radius over the 150 los along with the 16th and 84th percentiles are plotted against the virial mass of clusters as shown in Figs 7 and 8. We found that eight of the WDM clusters produce higher lensing cross-sections than their CDM counterparts. We also found that seven of the WDM clusters produce larger Einstein radii than the CDM analogues. These figures show that the WDM version of most of the clusters produced significantly higher lensing efficiency in comparison to its CDM counterpart. The figures also demonstrate that the variations of the lensing cross-sections and Einstein radii of most clusters are significantly wide due to the triaxiality of clusters and existence of big subhaloes. This means that the number of los used in this paper are fairly enough to study the lensing characteristics of galaxy clusters. We also think that the clusters that have narrow distributions are probably more spherical and have fewer big subhaloes. The WDM version of the cluster c10 which produces a higher lensing cross-section, has actually a slightly smaller Einstein radius than its CDM counterpart. This is most likely due to the effect of far away substructures from a cluster’s core, if these subhaloes are massive enough, they can develop their own critical lines and caustics, such that they can contribute to the cross-section of the clusters, while their contribution to Einstein radius is negligible. We found that the cross-section and Einstein radius of the WDM version of our sample could be boosted by up to ~29 and ~24 per cent, receptively, in comparison to its CDM counterpart.

We perform a least-squares fit using log ($\sigma_{7.5}$) = $m$ log ($\theta_E$) + $b$ to the data of both samples. The slope, intercept and the correlation coefficient of the linear fitting are shown in Fig. 6. The black lines in Fig. 6 show a linear fit to each data set. We found a shallower slope and higher normalization in the WDM cosmology in comparison to the CDM one. This suggests that the contribution of the source ellipticity to the lensing cross-section is more important for the WDM version of the clusters, which is another evidence for the effect of larger size subhaloes in the WDM clusters. These large substructures significantly modify the shear field such that the effect of sources ellipticity is more significant.

We compare our fit for the CDM sample with those from Meneghetti et al. (2011) and Kiledar et al. (2012) as shown in green and red lines, respectively, in the upper panel of Fig. 6. Our fit agrees with that of Meneghetti et al. (2011), even though we still get a shallower slope and higher normalization in the WDM cosmology in comparison to the CDM one. This suggests that the contribution of the source ellipticity to the lensing cross-section is more important for the WDM version of the clusters, which is another evidence for the effect of larger size subhaloes in the WDM clusters. These large substructures significantly modify the shear field such that the effect of sources ellipticity is more significant.

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fitting different geometrical figures to the images, they considered ellipses, circles, rectangles and rings, and then they pick up the figure whose circumference matches best with that of the image under consideration. A comparison between the eigenvalues technique for computing the lensing cross-section and that adopted in Meneghetti et al. (2011) has been done by Redlich et al. (2012). These authors have found that the geometrical figures fitting method in Meneghetti et al. (2011) fits 99 per cent of the lensed arcs to rectangles and that method gives identical results to the eigenvalues method if only ellipses are fitted to the lensed arcs. This means that the geometrical fitting method gives a higher lensing cross-section than the eigenvalues method by a factor of $\frac{4\pi}{\pi}$. The huge disagreement between our relation and that from Killedar et al. (2012) by considering their subsample of relaxed clusters at redshift $z_L = 0.25$ can most likely be due to some physical differences between our clusters and theirs (e.g. different abundance of substructures).

4 WDM (SLIGHTLY) STRONGER LENSES THAN CDM?

This result might seem initially count intuitive since WDM clusters contain less substructures and CDM clusters are more concentrated. A more centrally concentrated mass distribution will produce stronger convergence field, and more substructures will increase the shear field. Based on this naive prediction, the CDM clusters should produce larger cross-sections and Einstein radii than the WDM counterparts. However, our results show that eight of the WDM clusters produce higher lensing efficiency than the CDM version of the clusters. We found that the key difference in the WDM clusters is the existence of more massive and more extended substructures than those in the CDM counterparts. Those subhaloes significantly enhance the shear field of the host cluster if they are orthogonally aligned with the los and they increase the convergence if they are aligned along the los. We also found that these WDM subhaloes significantly perturb the radial critical line such that the lensing cross-section for radial arcs is higher in the WDM clusters than the CDM ones.
To pin down the effect substructures have on the lensing distribution, we examine their internal properties and mass growth. In the companion study Elahi et al. (2014), we identify all bound (sub)haloes of every snapshot of our simulated clusters using VELOCIraptor (Elahi, Thacker & Widrow 2011). This code identifies haloes using a 3D friends-of-friends algorithm with a linking length of 0.2 times the interparticle spacing, then identifies dynamically distinct phase-space structures residing within each halo. We briefly present here some of the salient findings concerning substructure here (for more details see the companion paper Elahi et al. 2014). The upper panel of Fig. 9 shows the maximum circular velocity radius as a function of enclosed mass for the subhaloes associated with the most massive halo in the 10 clusters of CDM (cyan crosses) and the WDM (violet crosses) versions. The lower panel on the other hand shows the same plot for the CDM subhaloes that have analogues in the WDM version of clusters. The cross-catalogue is made by identifying each (sub)halo in catalogue A that shares particles with a (sub)halo in catalogue B and the merit of the initial matches is determined from $M_{ij} = N_{A,B}^2/(N_A N_B)$ using the halo merger tree code of VELOCIraptor, which is a particle correlator (see Srisawat et al. 2013 for more details of this code). The figure illustrates that on average for a given mass, WDM subhaloes are more likely to have larger sizes than the CDM analogues. Fig. 10 shows the ratio between WDM and CDM subhalo counterparts of the mass enclosed at the maximum circular velocity radius. The error bars represent the median, 16th and 84th percentiles of several mass bins. The figure illustrates that the median of the logarithmic ratio is larger than zero, which means the WDM subhaloes are more massive than those in the CDM clusters. Fig. 11 shows the ratio between WDM and CDM subhalo counterparts of the maximum circular velocity radius. The error bars represent the median, 16th and 84th percentiles of several radial bins. Again, the figure demonstrates the possibility of having more extended WDM subhaloes than the CDM counterparts.

In order to check whether the boost in the lensing efficiency of the WDM clusters comes from the subhaloes, we compute the lensing properties of both versions of the most massive cluster (i.e. c1) with and without projecting the particles in the subhaloes on the lens plane. Fig. 12 shows the cross-section of 75 los of the WDM version versus the cross-section of the CDM counterpart by considering...
subhaloes (red circles) and excluding them (blue triangles). The figure demonstrates that by including the substructures, the WDM version produces higher lensing efficiency for 84 per cent of the los. This fraction drops off to 25 per cent when excluding the particles associated with substructures. This clearly proves that the higher lensing efficiency of the WDM clusters is due to the contribution of subhaloes.

5 DISCUSSION AND CONCLUSIONS

The lensing properties for two samples of simulated clusters in the WDM and CDM cosmologies have been studied. Based on the characteristics of WDM clusters, one would expect them to be weaker lenses than the CDM counterparts as they are less centrally concentrated and have fewer satellites than CDM analogues. The results of this study show that the WDM clusters are slightly more efficient lenses at producing giant arcs than the CDM counterparts. We found that the key difference in the WDM clusters that significantly enhances their lensing efficiency is the existence of more extended and more massive substructures than those in the CDM clusters. Those more massive substructures significantly enhance the shear and convergence fields which in turn lead to a higher lensing efficiency.

However, despite the enhancement in the lensing efficiency, WDM alone cannot account for the differences seen between theory and observation. A more robust comparison regarding the effect of small-scale substructures can be done by adopting recursive-TCM projection algorithm that has recently been proposed by Angulo et al. (2013). The advantage of this method is that there is no need for Gaussian smoothing, and therefore small-scale substructures that are washed away by smoothing remain. However, the contribution of such small-scale substructures to the lensing efficiency is probably not significant.

We found that eight of the WDM clusters produce higher cross-sections and Einstein radii than their CDM counterparts. However, we found that the WDM version of the cluster c10 produces higher cross-section but smaller Einstein radius than its CDM counterpart. This can be attributed to the existence of some substructures which are strong enough and relatively far away from the centre of a cluster to develop their own critical lines and caustics. In such a case, these substructures contribute to the lensing cross-section but not for Einstein radius as we compute Einstein radius only for the core of clusters.

We have fit the function log (σ_T5) = m log (θ_E) + b to our data and found that the relation for the WDM sample has a shallower slope and higher normalization than the CDM counterpart. This is again due to the effect of substructures which substantially affect the shear field. The source ellipticity becomes more significant for the lensing cross-section of the WDM lenses due to their larger subhaloes. By comparing our results to those of Meneghetti et al. (2011) and Killer et al. (2012), we find a good agreement between our fitting and that from Meneghetti et al. (2011). However, our fit disagrees with that from Killer et al. (2012), this can be due to some physical differences of our clusters and theirs.

We conclude that despite the slight enhancement in the lensing efficiency of WDM clusters, they fail to explain the arc statistics problem as none of our clusters can reproduce the observed cross-sections and Einstein radii of four clusters in Meneghetti et al. (2011). As a possible solution, we are going to study the effects of baryonic physics on the lensing properties of WDM clusters. Furthermore, we are planning to analyse the lensing characteristics of clusters of galaxies drawn from other cosmologies, such as coupled dark matter–dark energy models.

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